

Pre-Inflationary Relics and Large Scale Anomalies

Ely D. Kovetz
University of Texas at Austin
Aspen Winter 2012, January 31st

Based on:

A. Ben-David, EDK and N. Itzhaki, “Parity in the CMB: Space Oddity”, ApJ, to appear 2012. arXiv:1108:1702

EDK, A. Ben-David and N. Itzhaki, “Giant Rings in the CMB Sky”, ApJ 724, 2010. arXiv:1005.3923

A. Fialkov, N. Itzhaki and EDK, “Cosmological Imprints of Pre-Inflationary Particles”, JCAP 2010, arXiv:0911.2100

Outline

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- Motivation: Pre-Inflationary Physics

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- Simplest Relic: Pre-Inflationary Particle

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- Pre-Inflationary Particle - Signatures

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- What to look for? (symmetries)
- Observational Tests: CMB Rings, Bulk Flow, Mirror Parity.
- Fitting a model to the data

Λ CDM Universe

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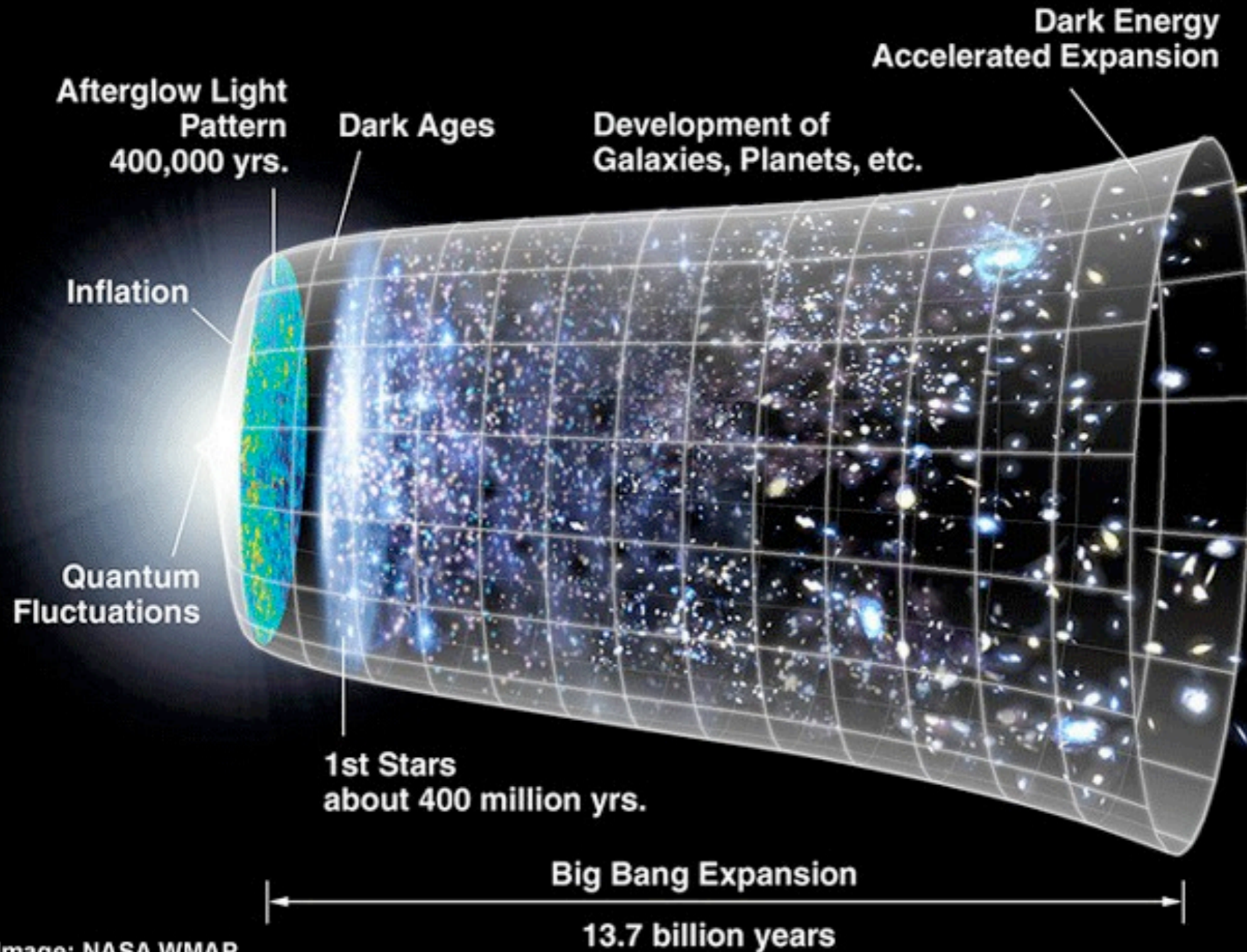


Image: NASA WMAP

Pre-Inflationary Particle - Model

(Fialkov, Itzhaki, EDK, JCAP 2010)

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- Therefore, it induces a contribution to the e.o.m of the inflaton pert. of the form:

$$\delta\ddot{\phi} + 3H\delta\dot{\phi} - \frac{1}{a(t)^2} \nabla^2 \delta\phi = \lambda \frac{\delta^3(x)}{a(t)^3}$$

$$\lambda = \frac{1}{2} \frac{V'}{V} m_{PIP} - \frac{\partial}{\partial\phi} m_{PIP}$$

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PIP - Signatures

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- Creates an (overdense) Spherically Symmetric Cosmic Defect (SSCD).

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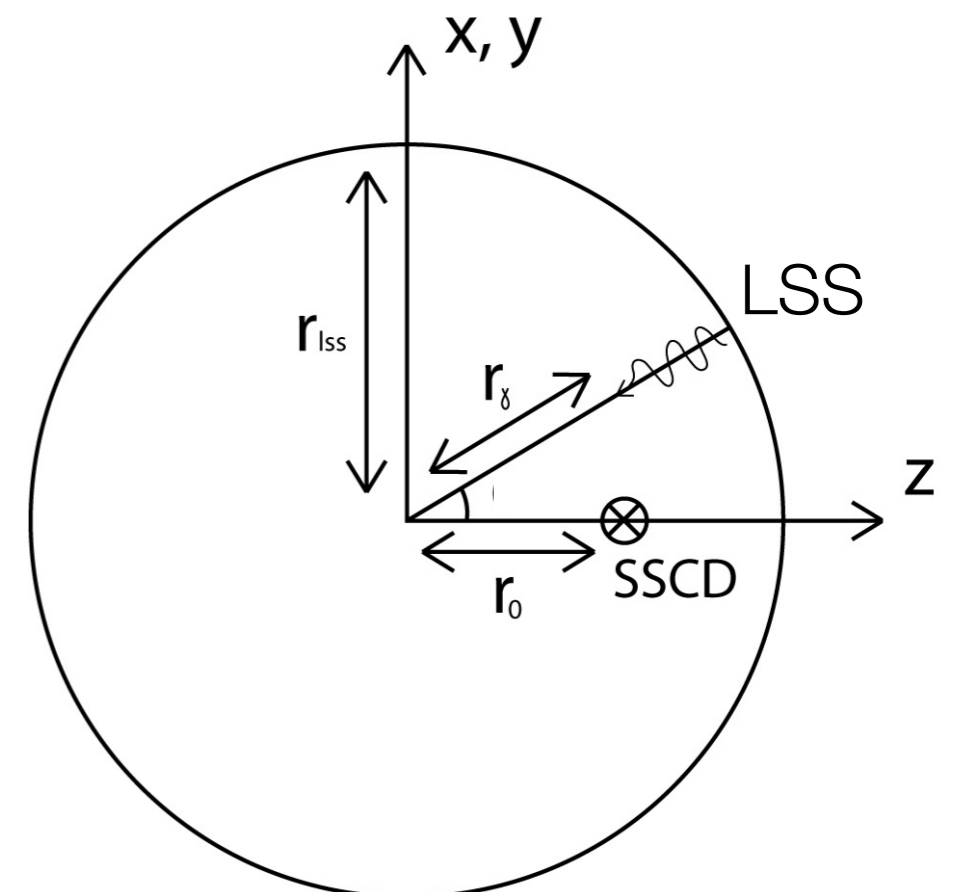
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- Induces (SW + ISW effects) azimuthally symmetric 2nd order CMB anisotropies.

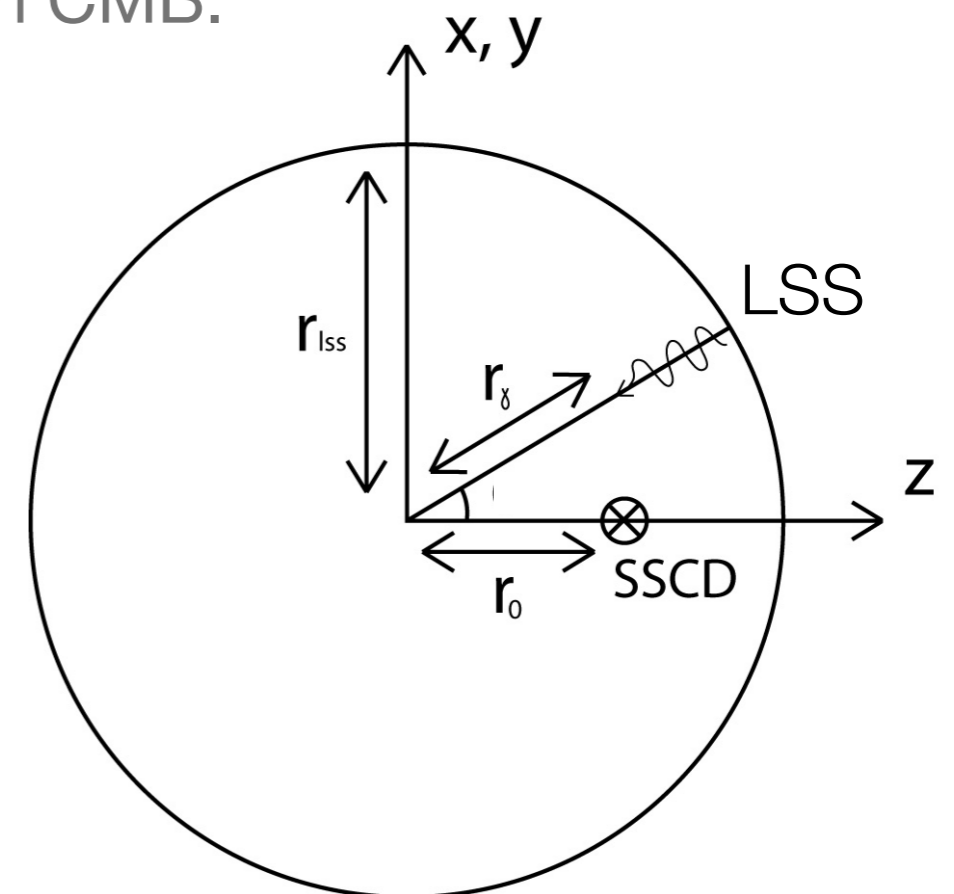
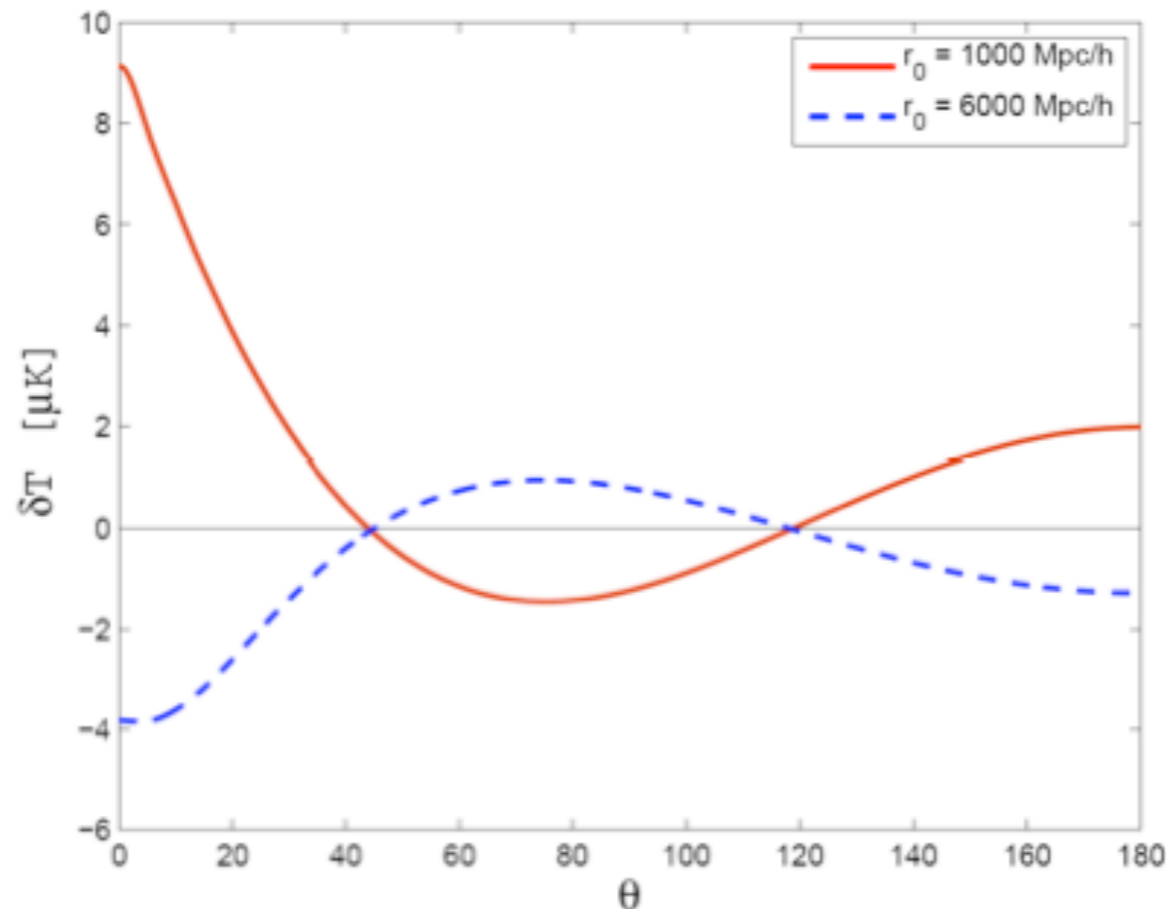


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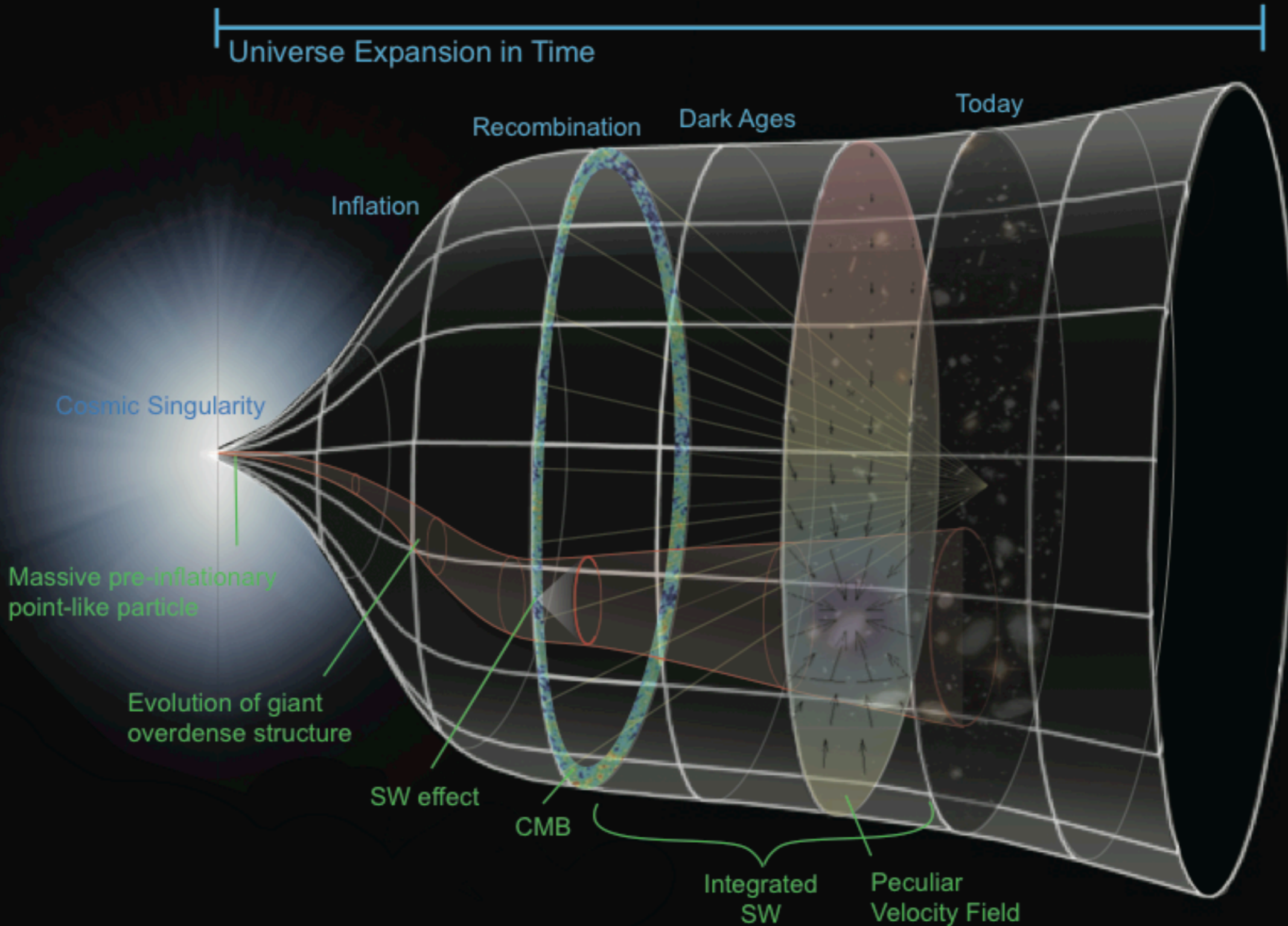
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- Generic imprint: concentric hot and cold rings in CMB.



Λ CDM Universe + Pre-Inflationary Particle (PIP)



Where to look? Large Scales

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Where to look? Large Scales

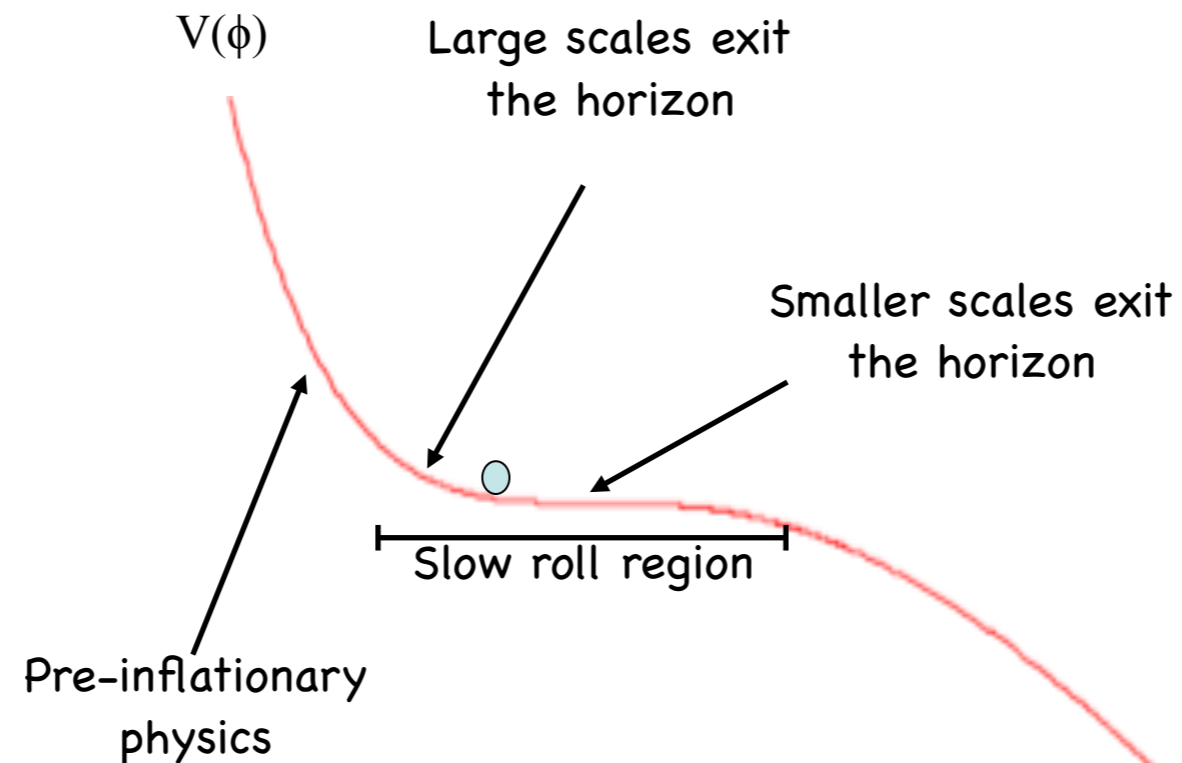
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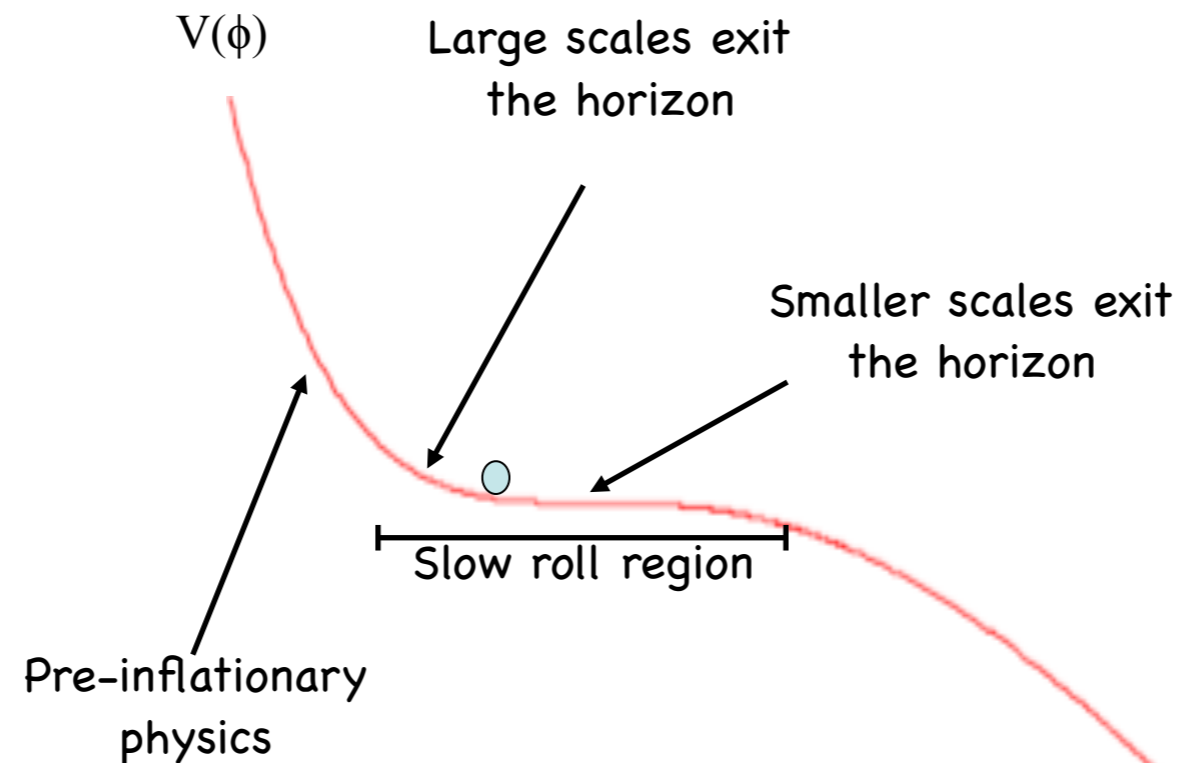
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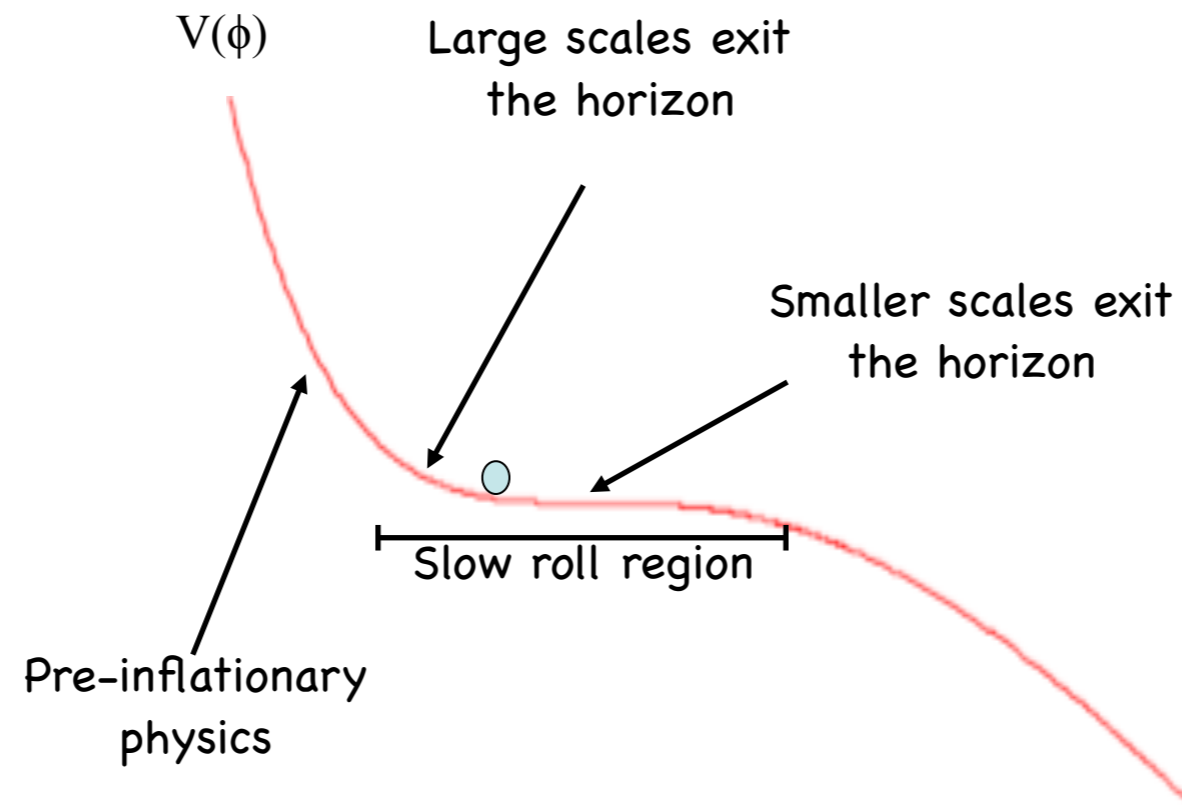


- If relic couples to the inflaton, effect will start at largest scales and drop with ℓ .

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- If relic couples to the inflaton, effect will start at largest scales and drop with ℓ .
- Logarithmic gravitational potential \rightarrow long-range effect.

What to look for? Symmetries

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▶ Test: CMB Giant Rings!

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- Overdense region with long-range gravitational potential:

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- Overdense region with long-range gravitational potential:

▶ Test: Bulk flow (towards this “great attractor”)!

Rings Score - Definition

(EDK, Ben-David and Itzhaki, ApJ 2010)

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Q: Looking for $U(1)$ symmetry - are there unusual rings in the CMB?

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A: We check this using a dedicated score for each direction \hat{n}

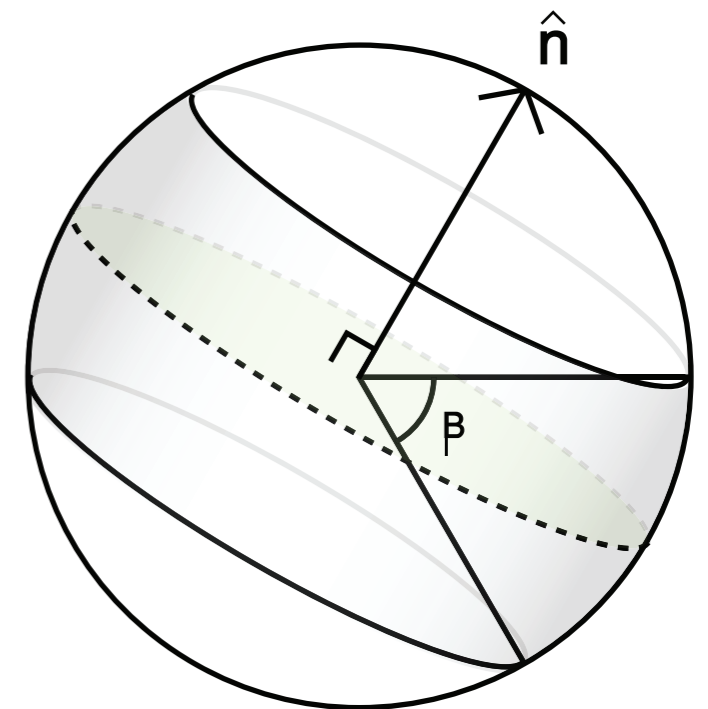
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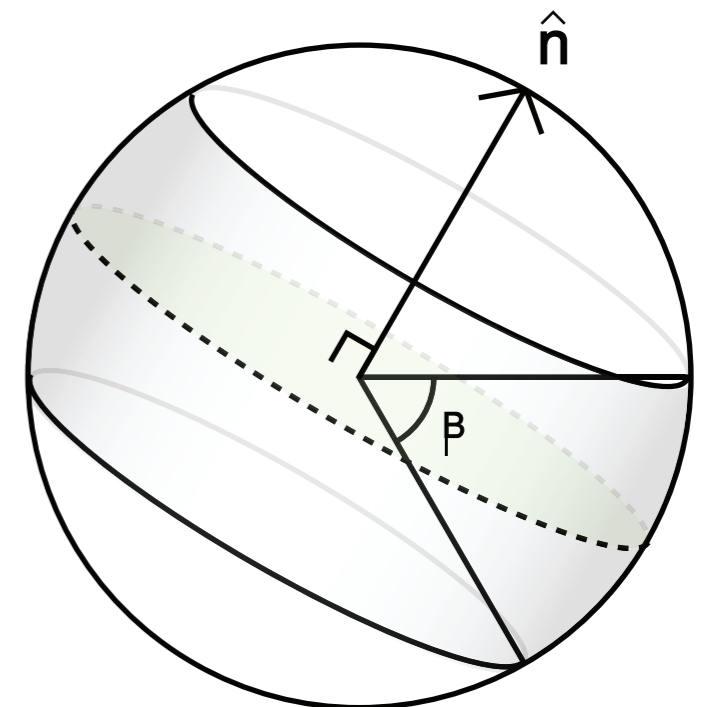
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- Focus on the large rings



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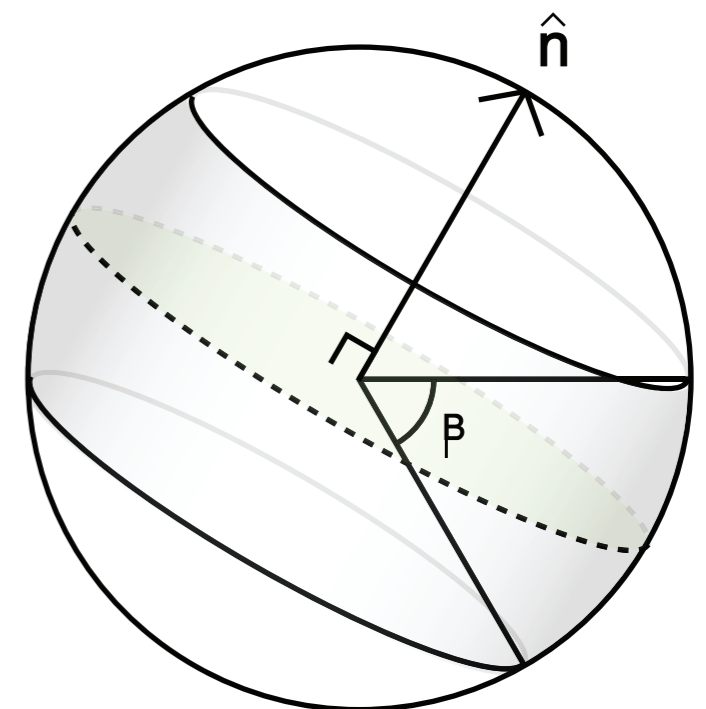
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- Focus on the large rings
- Choose a band of width β around $\theta = \pi/2$



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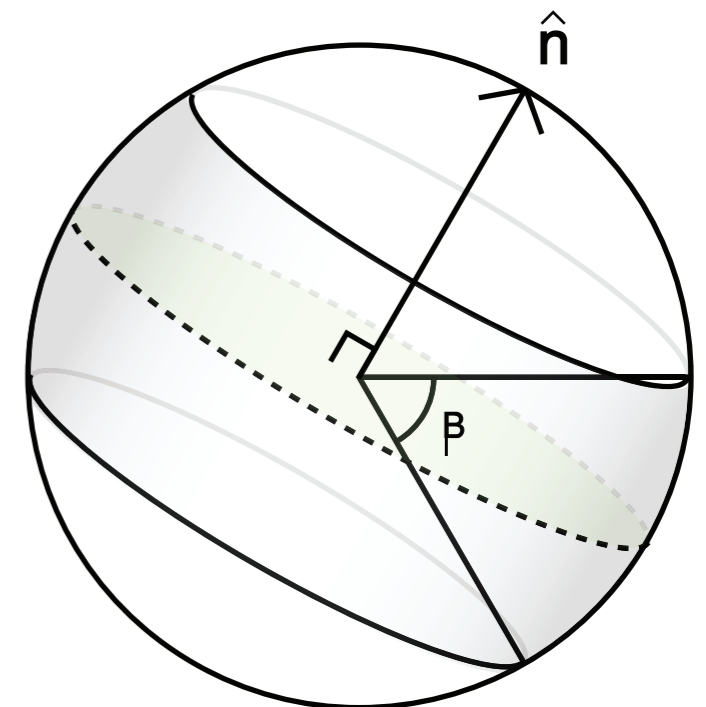
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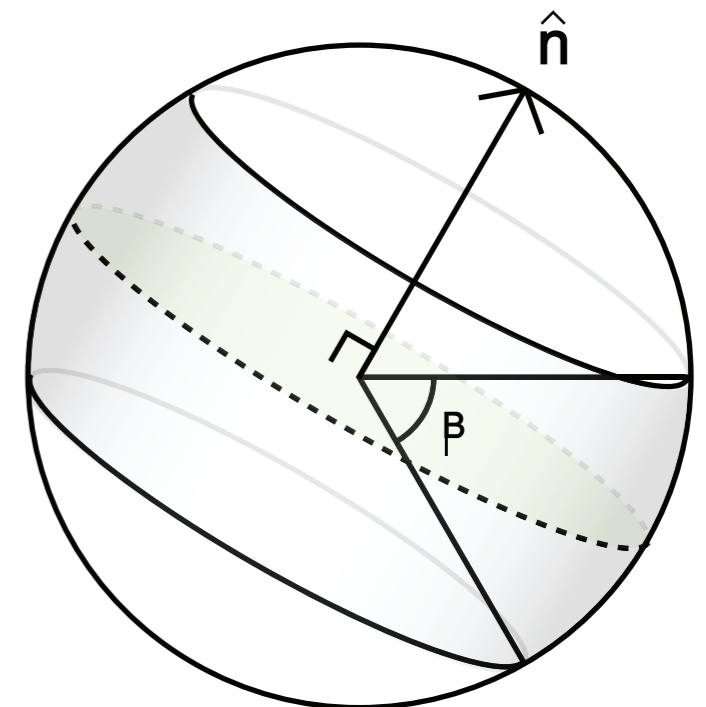
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- T_0 = mean of the total map



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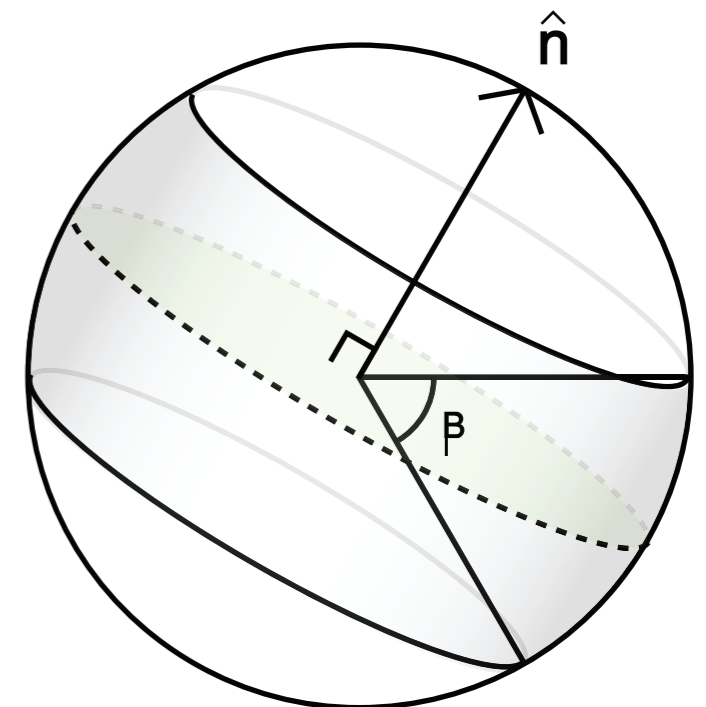
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- Focus on the large rings
- Choose a band of width β around $\theta = \pi/2$
- $T(\theta, \hat{n})$ = mean temperature of an infinitesimal ring
- T_0 = mean of the total map
- Calculate the following:

$$R(\beta, \hat{n}) = \int_{\frac{\pi-\beta}{2}}^{\frac{\pi+\beta}{2}} d(\cos \theta) \tilde{T}^2(\theta, \hat{n}),$$

$$\tilde{T}(\theta, \hat{n}) = T(\theta, \hat{n}) - T_0$$



Rings Score - Results

(EDK, Ben-David and Itzhaki, ApJ 2010)

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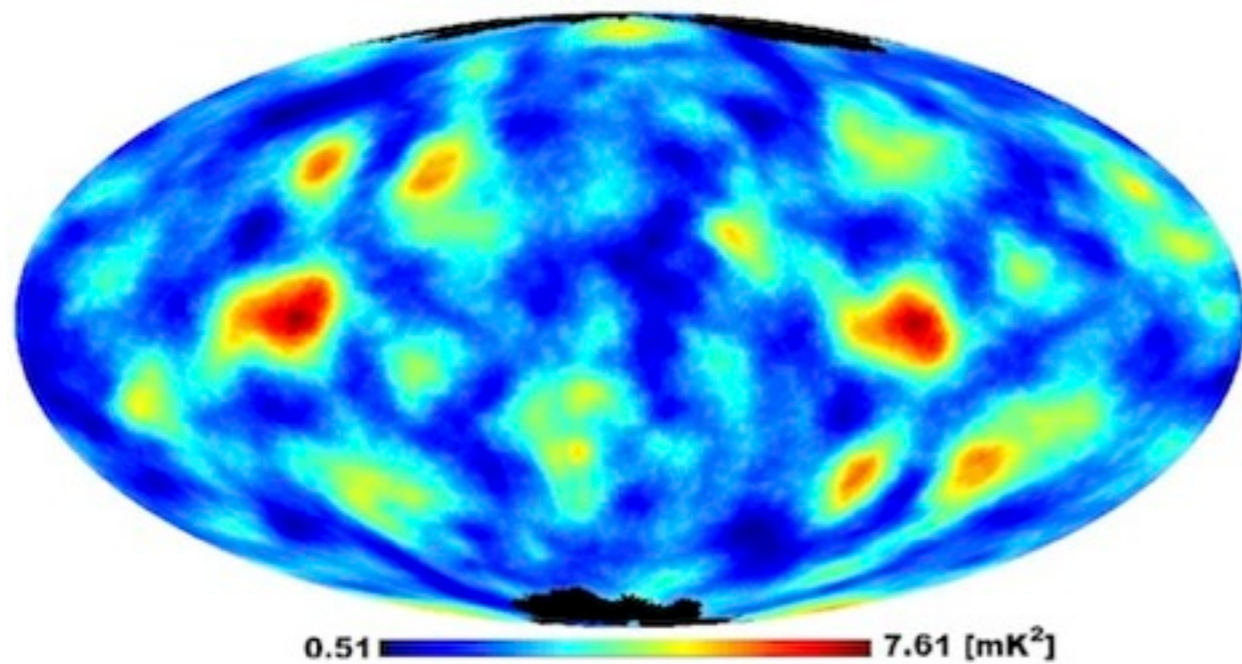
Our findings for WMAP7 ILC masked with KQ75:

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$$\beta = 90^\circ$$

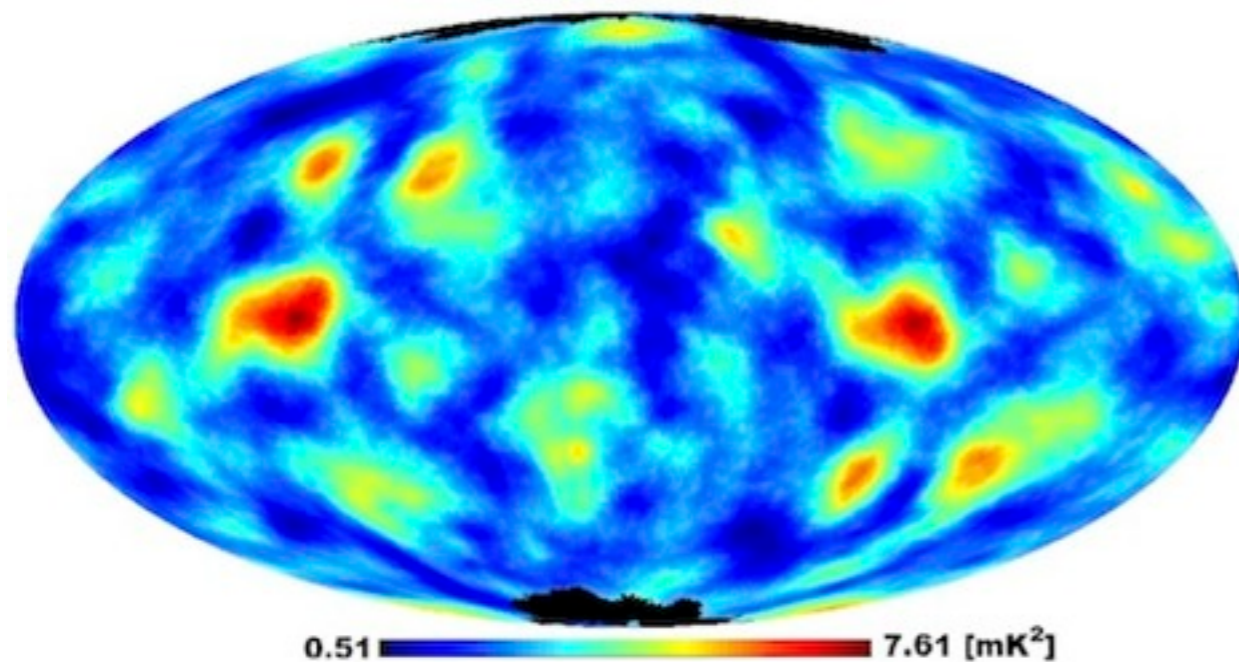
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Our findings for WMAP7 ILC masked with KQ75:

- We see a peak at around $(l, b) = (276^\circ, -1^\circ)$.



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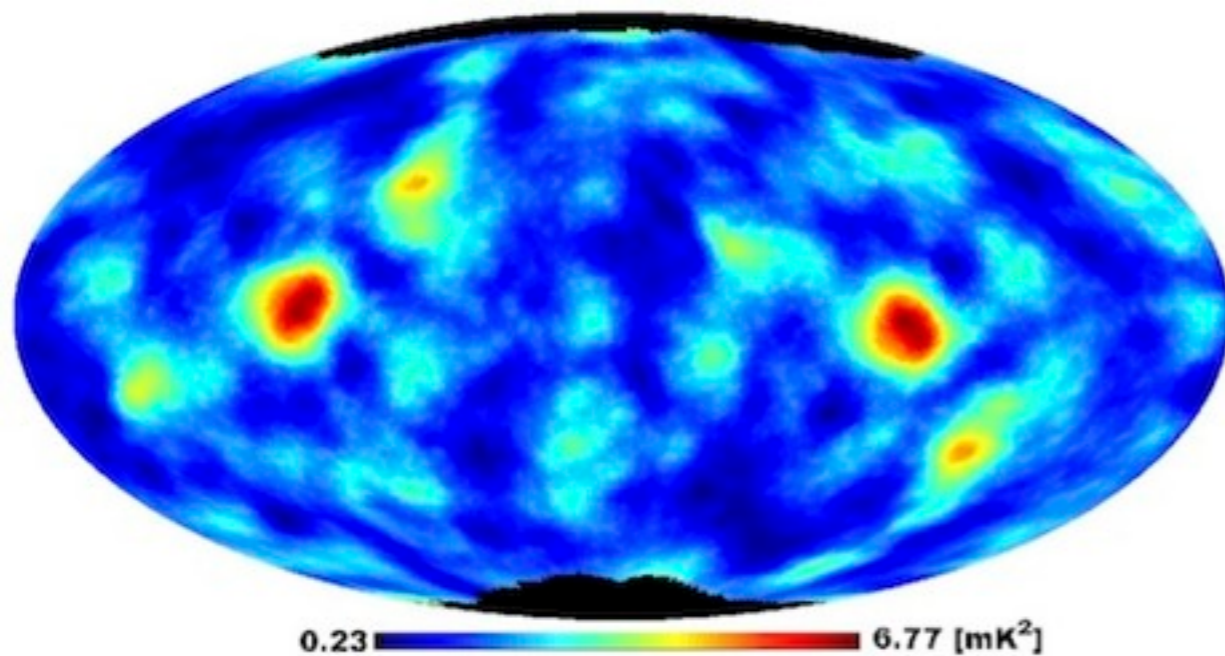
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$$\beta = 60^\circ$$

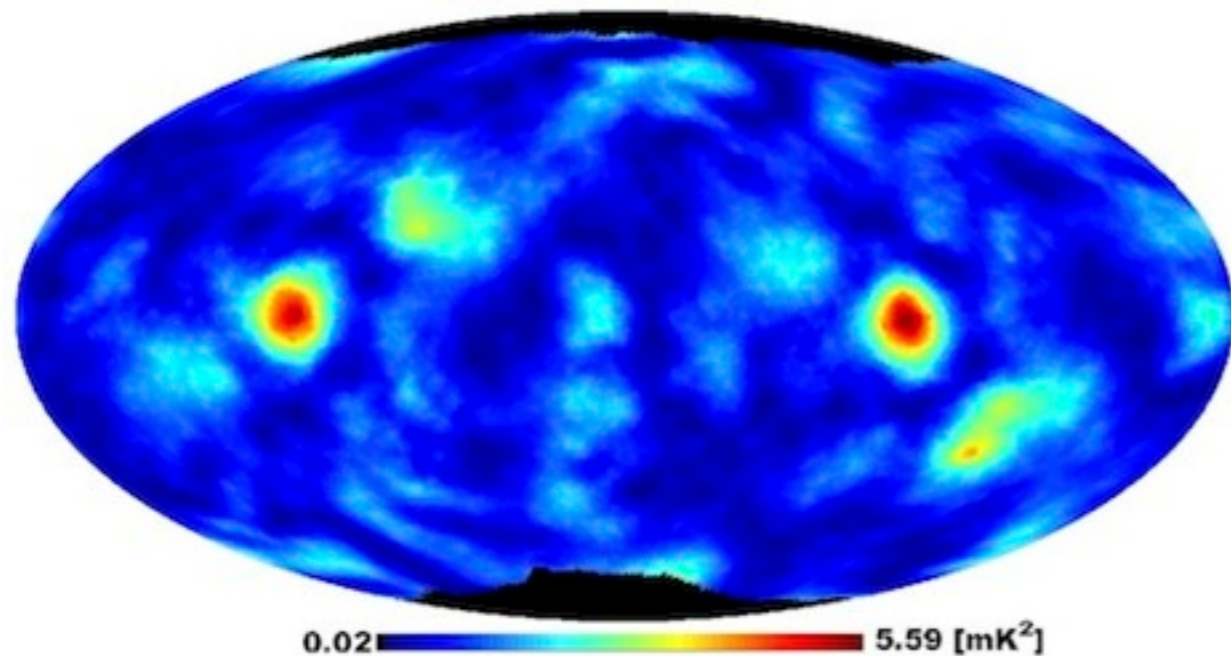
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$$\beta = 30^\circ$$

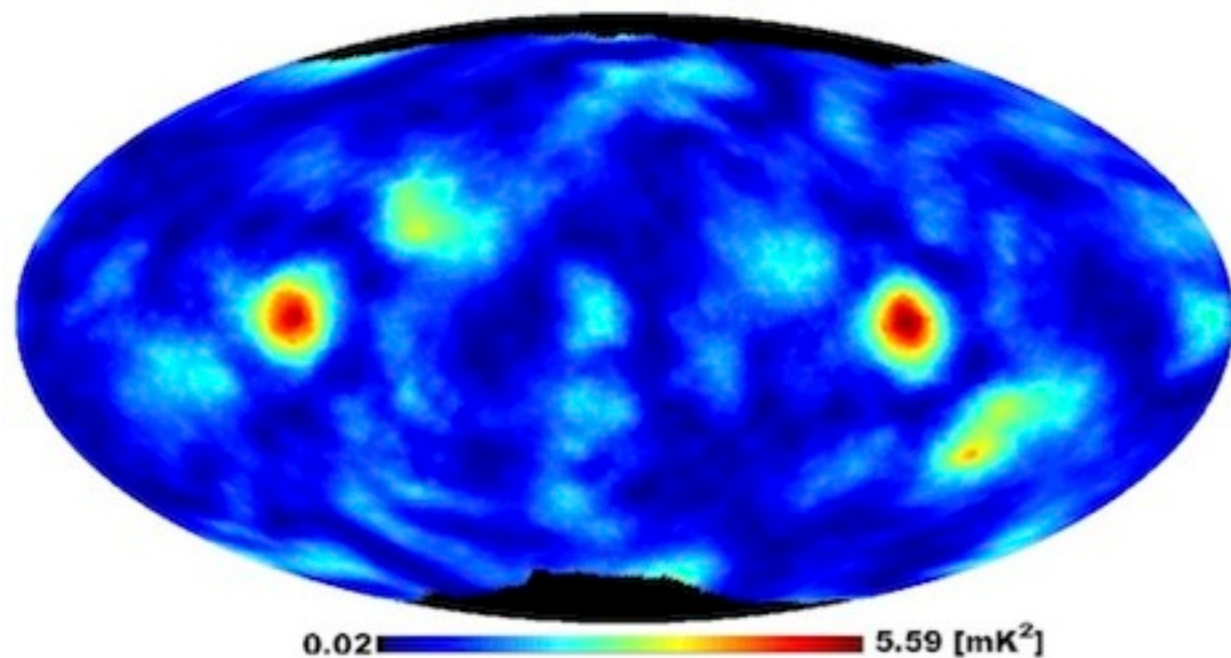
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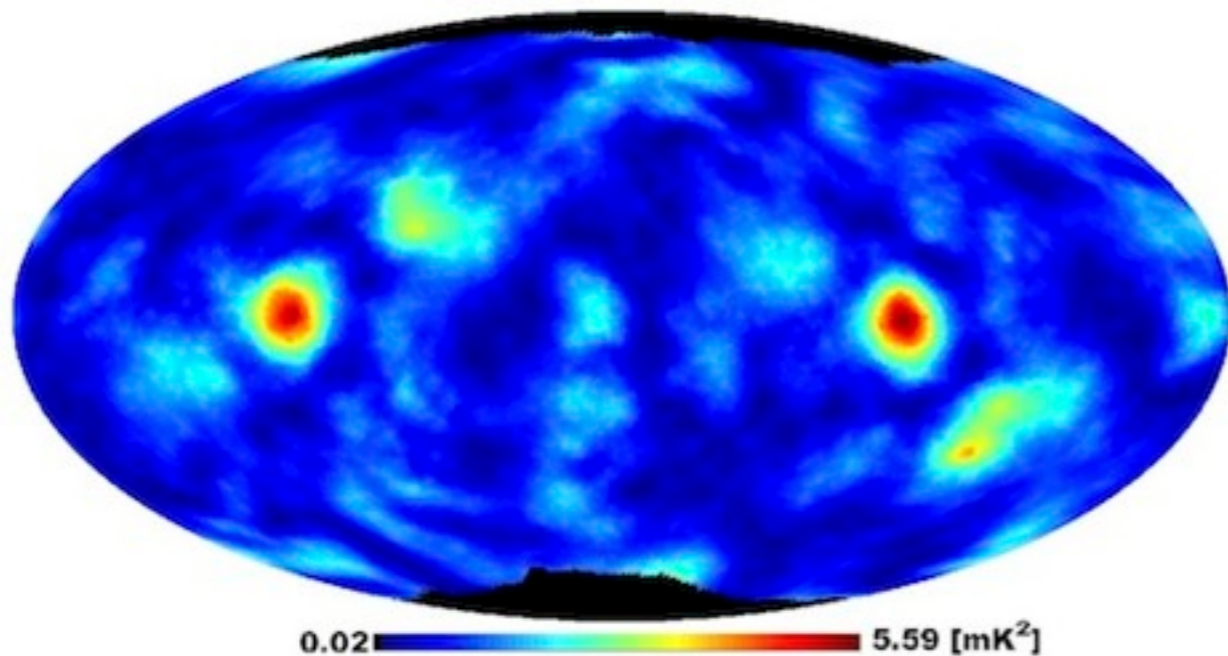
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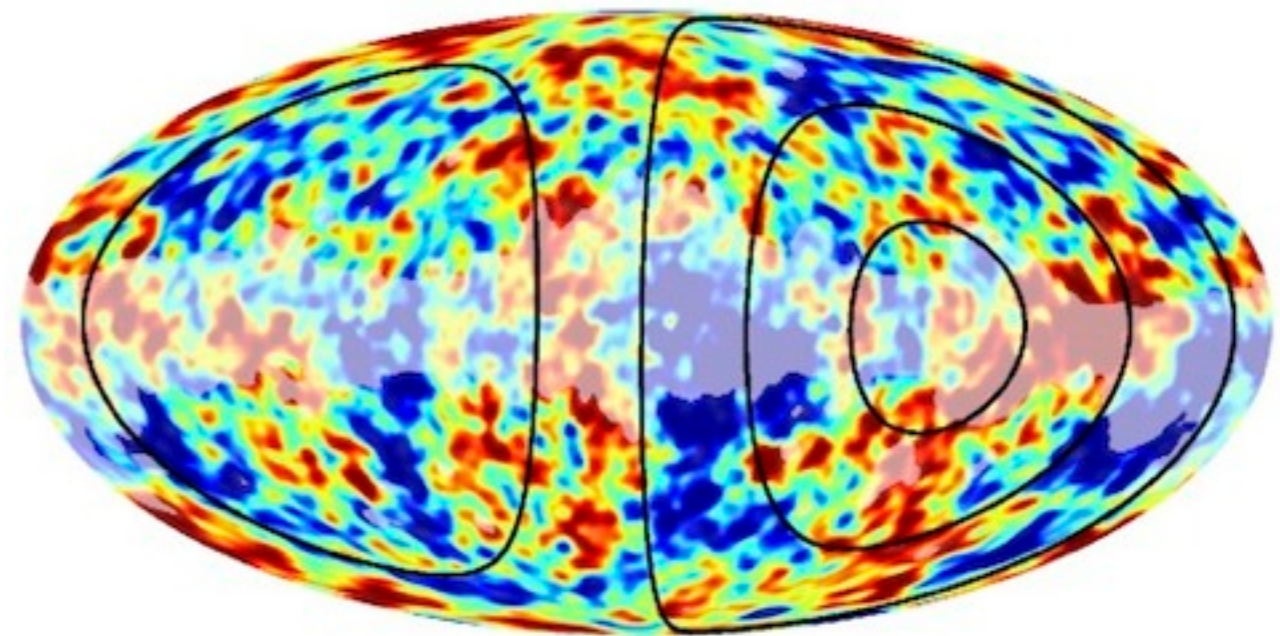
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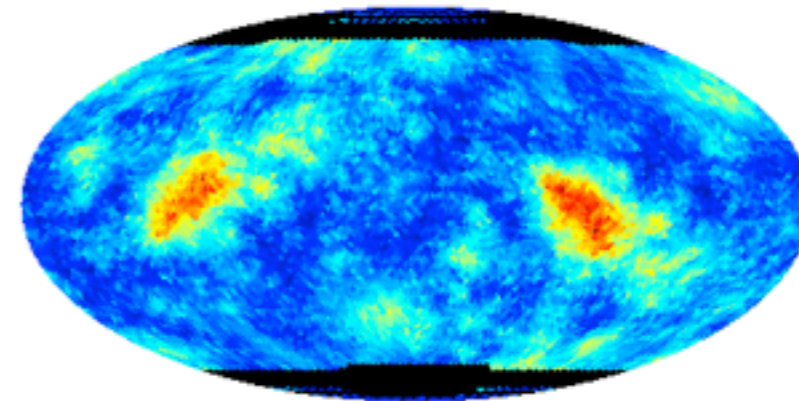
Rings in 7-year ILC map


Rings Score - COBE-DMR

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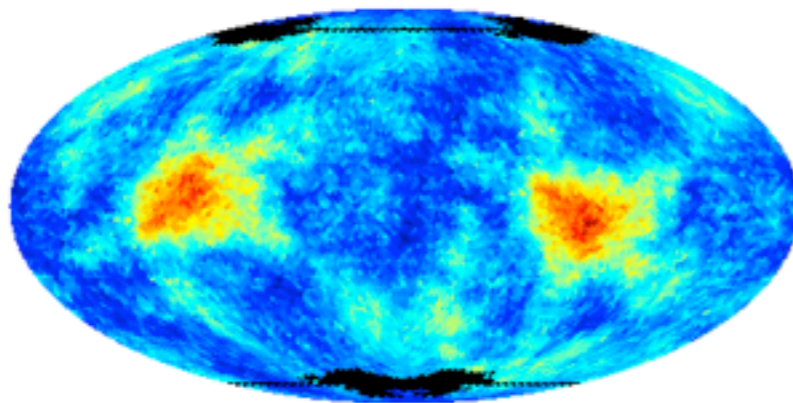
Low Resolution...


DMR f = 90GHZ, $\beta=60$



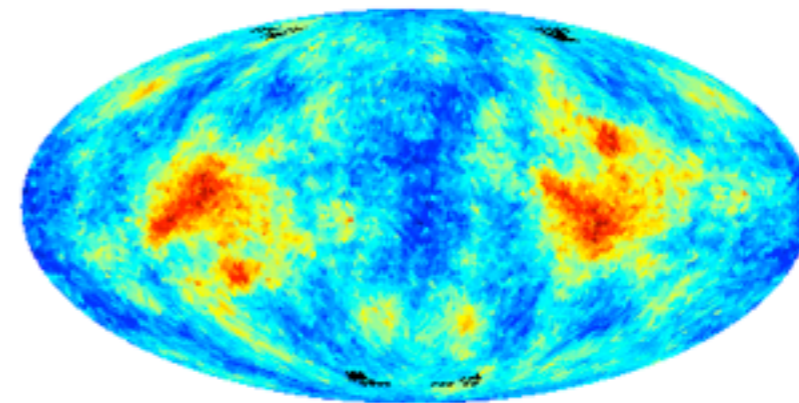
0.16  4.76
[mK²]


DMR f = 90GHZ, $\beta=90$



0.53  6.31
[mK²]

DMR f = 90GHZ, $\beta=120$



0.81  6.69
[mK²]

Bulk Flow - Direction

(EDK, Ben-David and Itzhaki, ApJ 2010)

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- Recent peculiar velocity “Bulk Flow” findings by Feldman et al. (MNRAS 2010):

$$v = 416 \pm 78 \text{ km/s} \quad (l, b) = (282^\circ \pm 11^\circ, 6^\circ \pm 6^\circ)$$

(Amplitude has been recently disputed by: Nusser, Branchini & Davis 2011)

Bulk Flow - Direction

(EDK, Ben-David and Itzhaki, ApJ 2010)

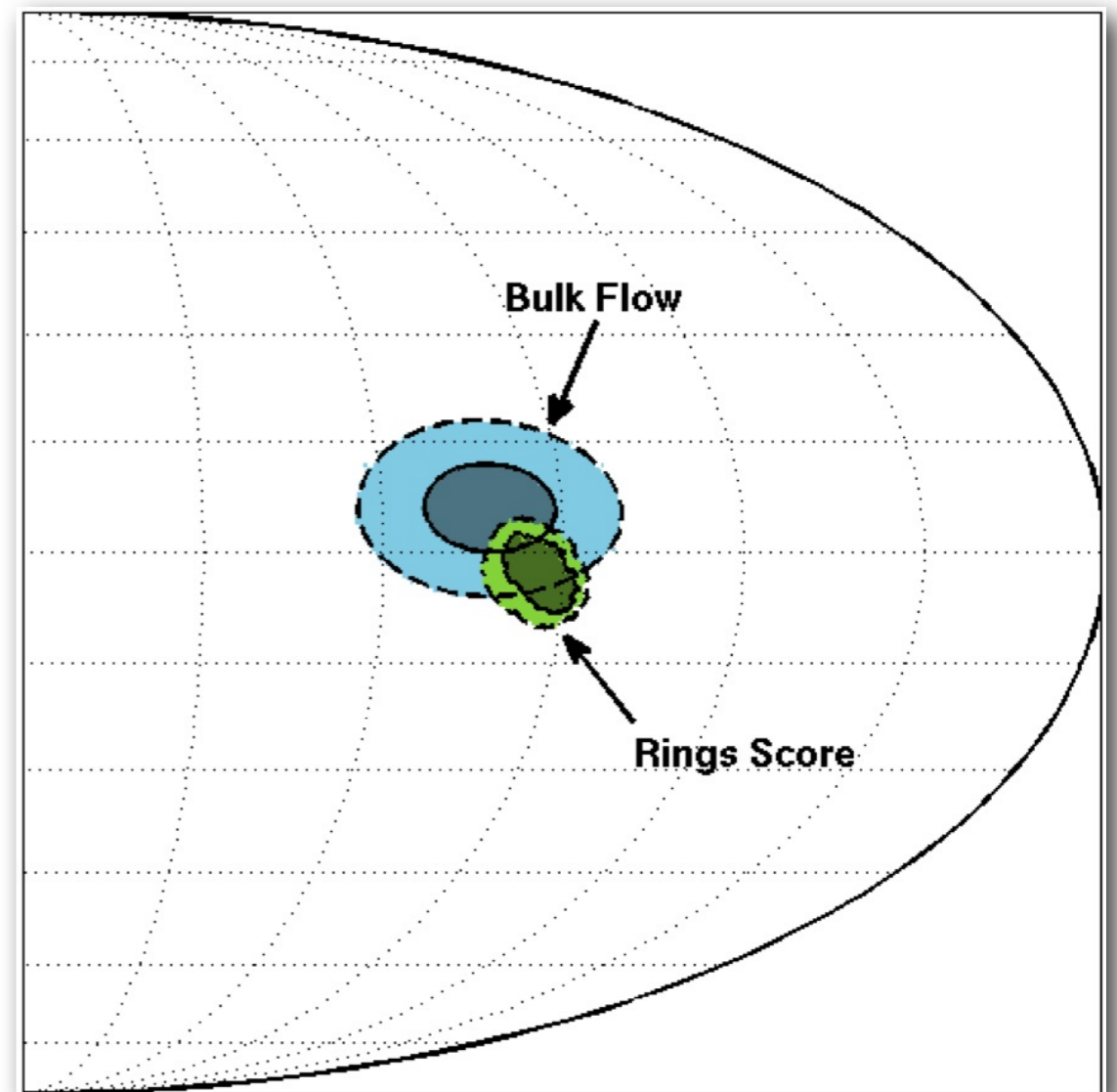
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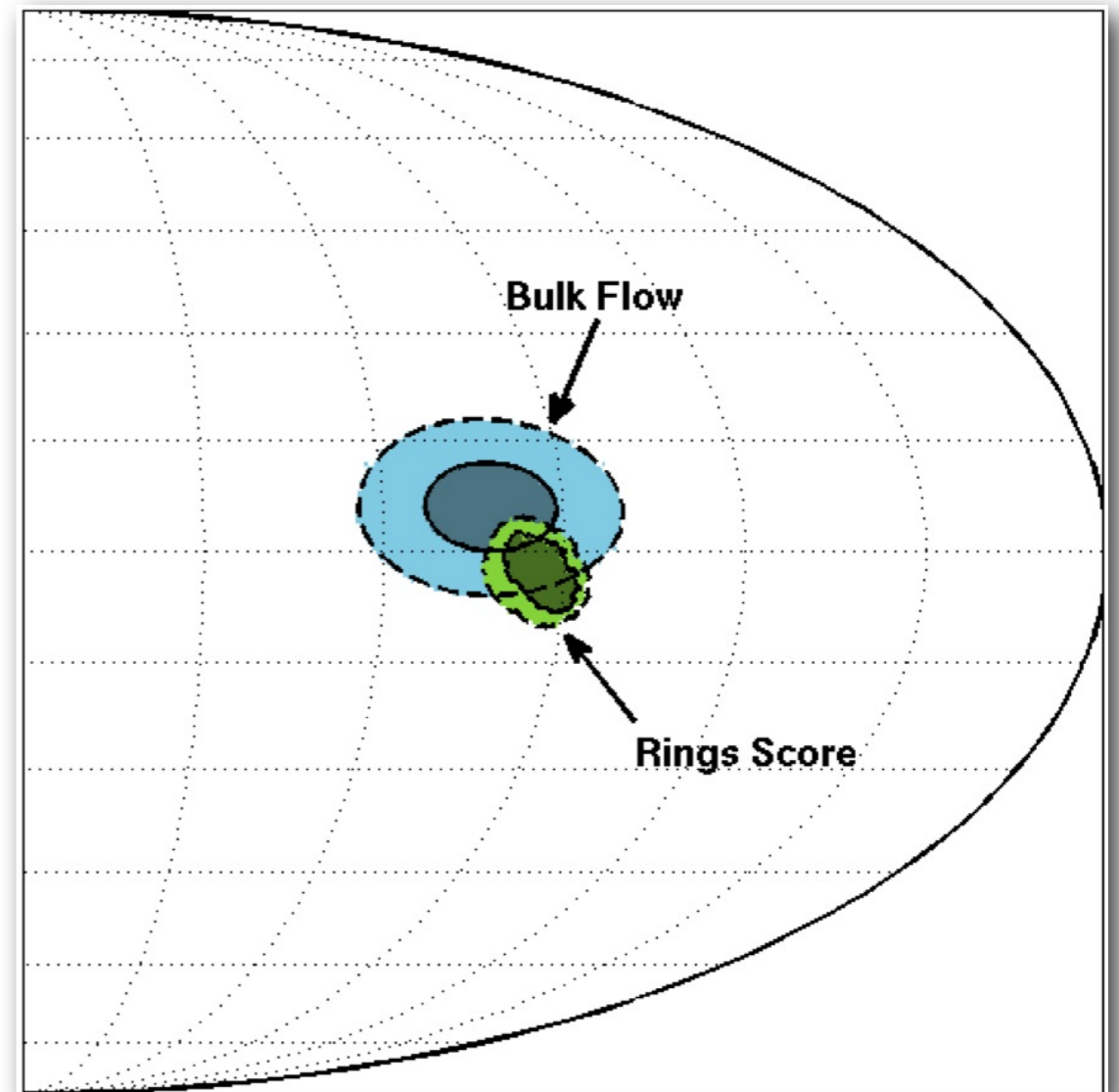
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~1% chance alignment
with the center of the rings!



Parity - “S-Statistic”

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- Looking for mirror parity, use the “S-Statistic”: (de Oliveira-Costa, Smoot & Starobinsky, 1995)

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Parity - “S-Statistic”

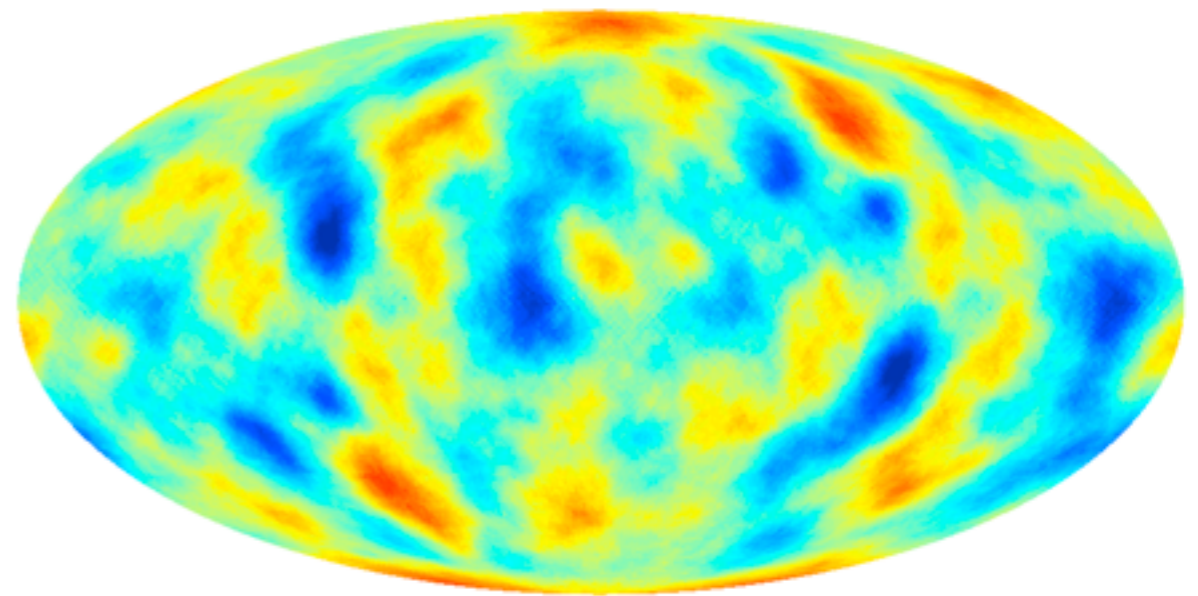
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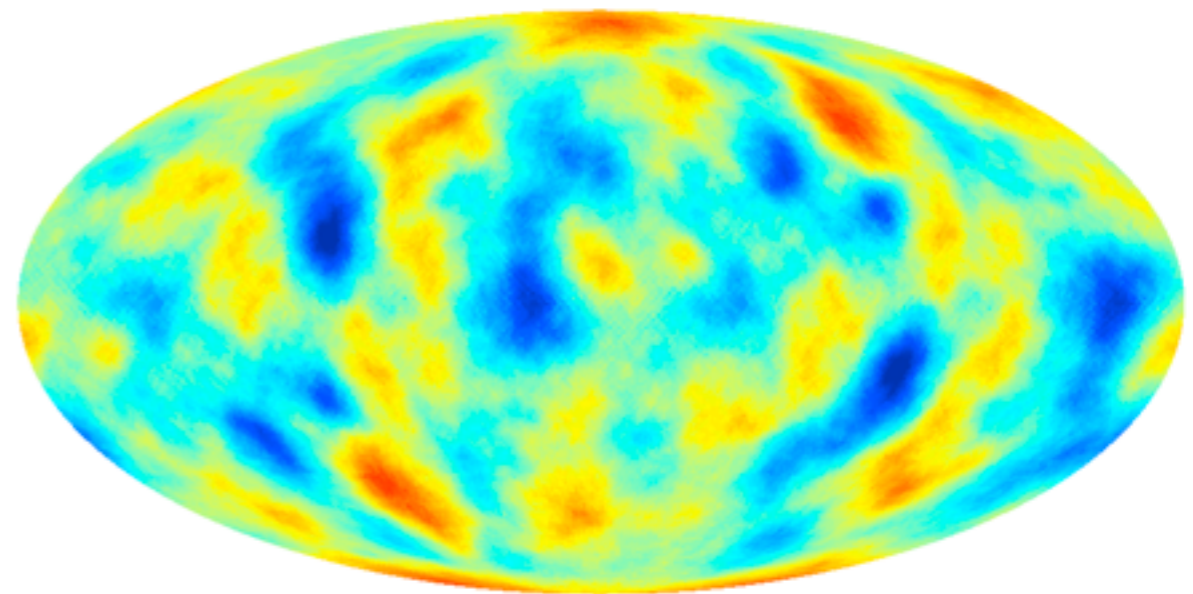
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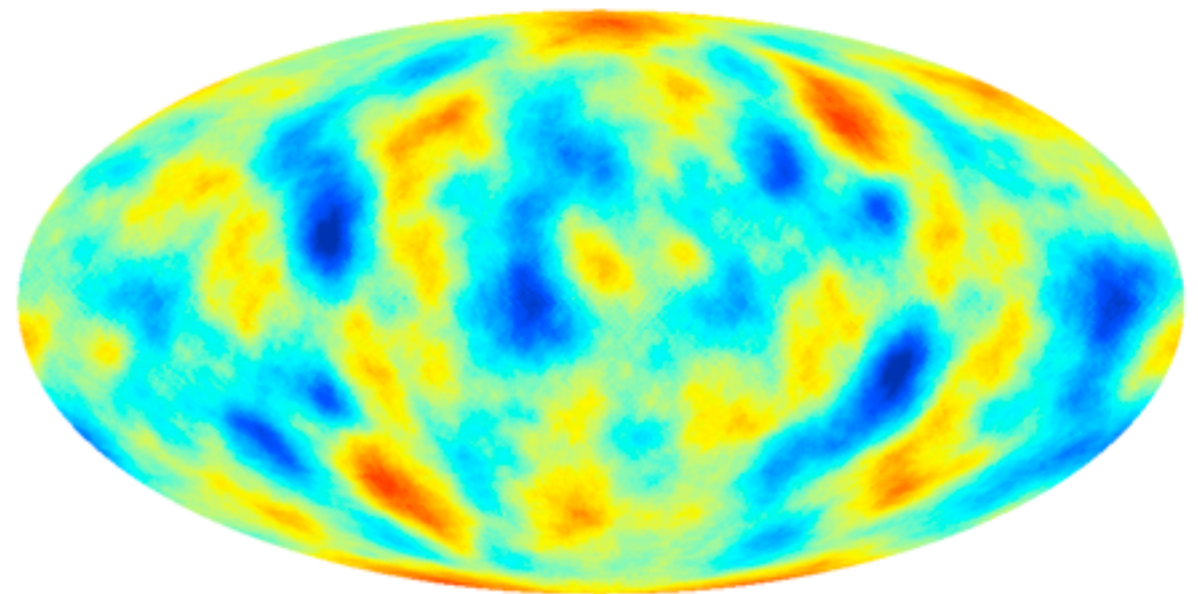
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Parity - “S-Statistic”

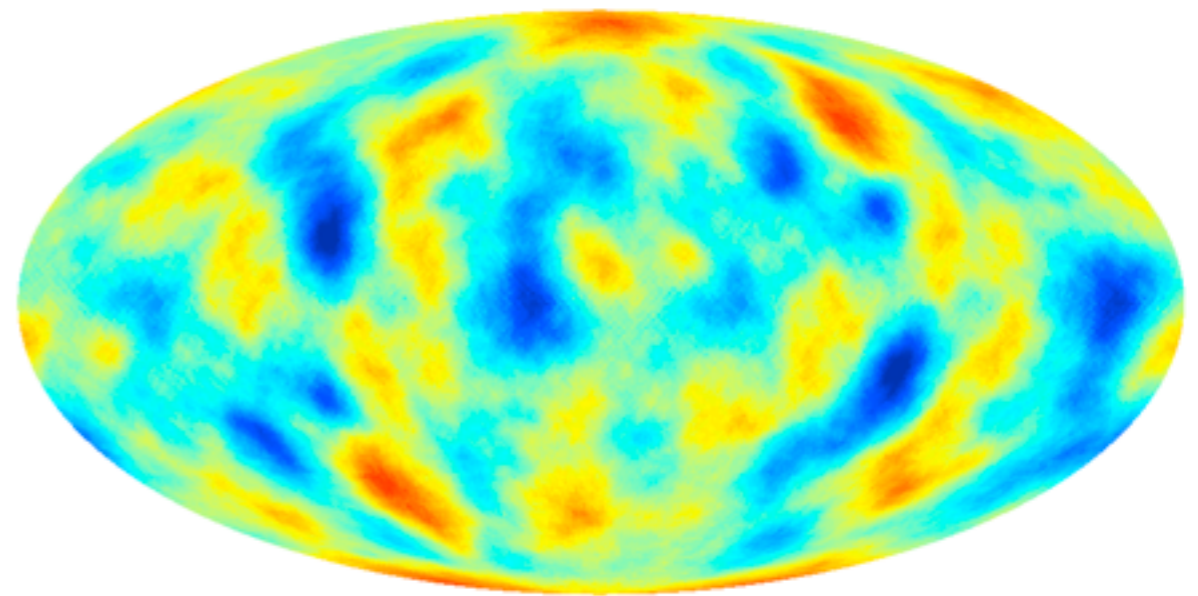
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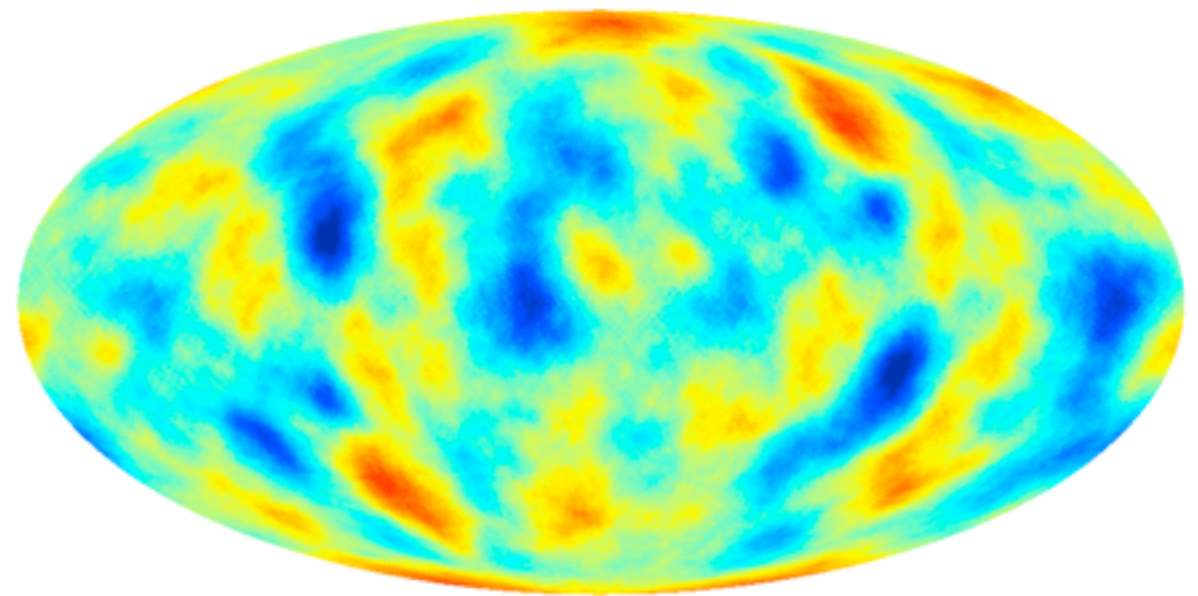
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- Masking the galactic plane results in strong bias of the S-Statistic.

Parity - Even/Odd Multipoles

(Ben-David, EDK and Itzhaki, ApJ 2012)

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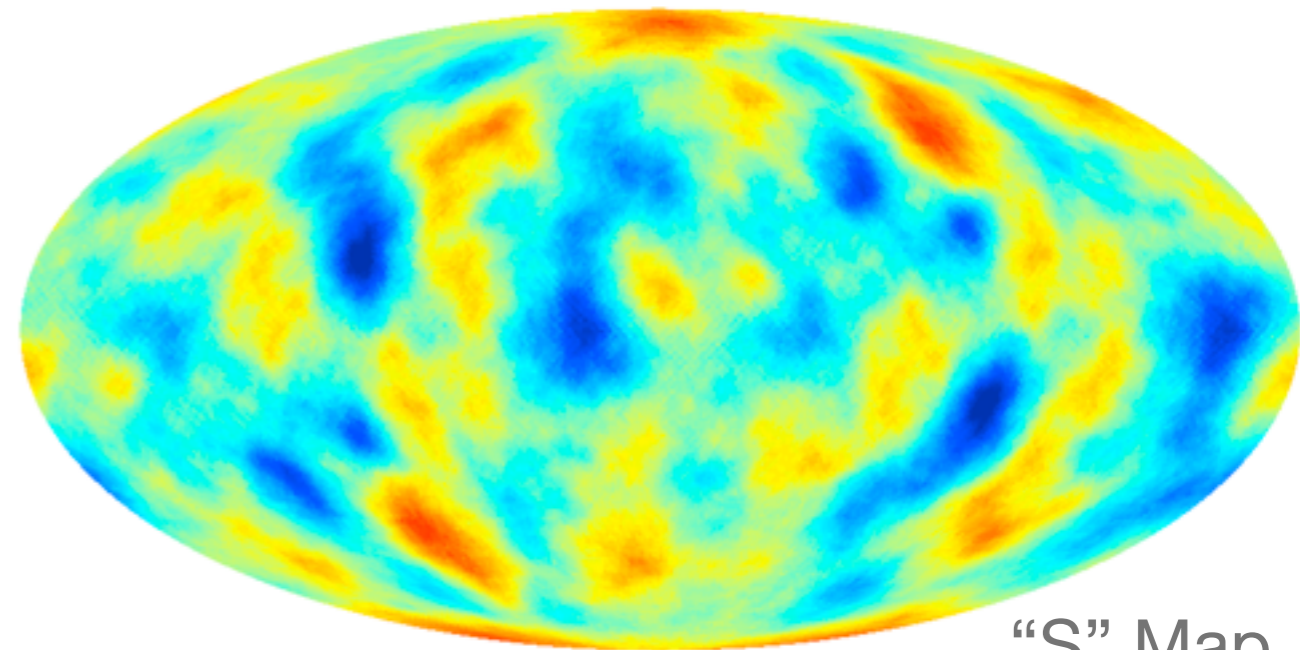
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“S” Map

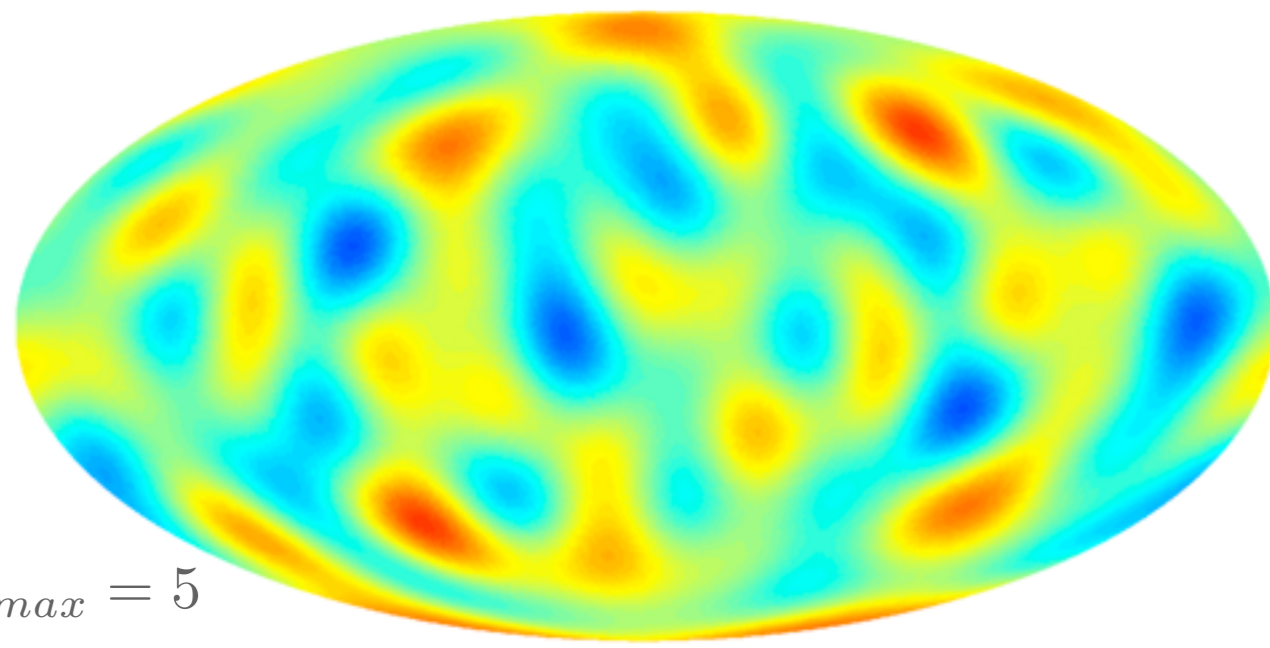
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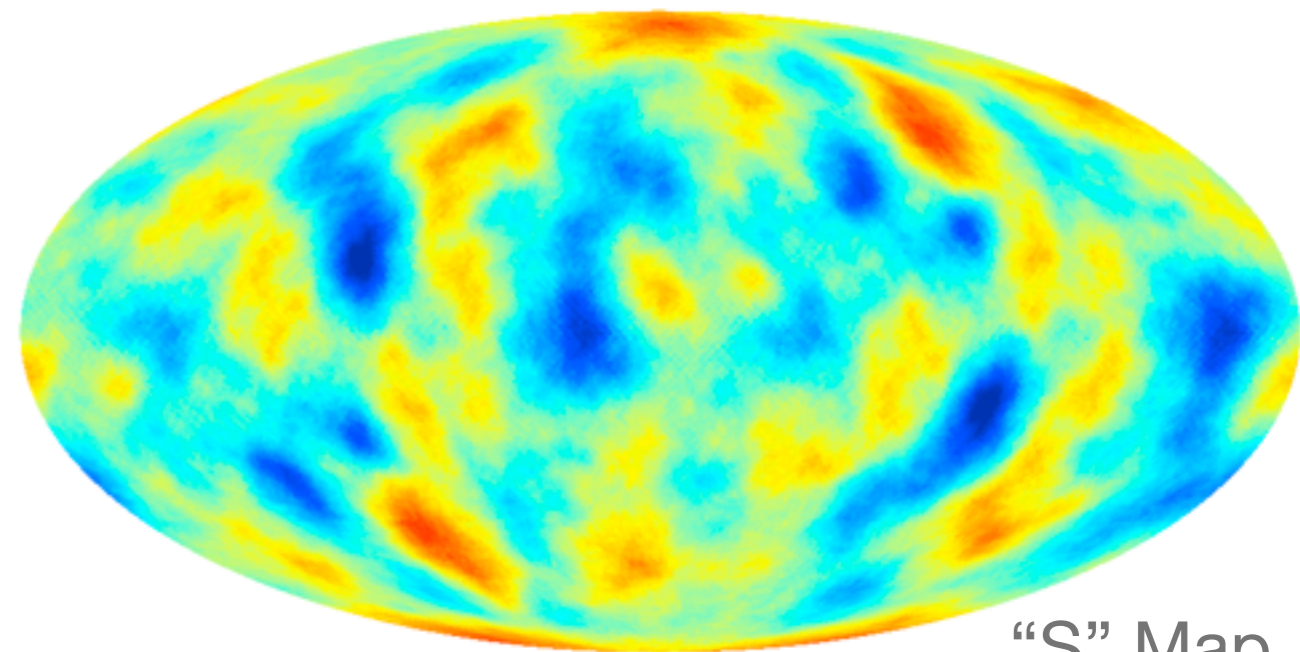
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$\ell_{\max} = 5$

Parity Map



"S" Map

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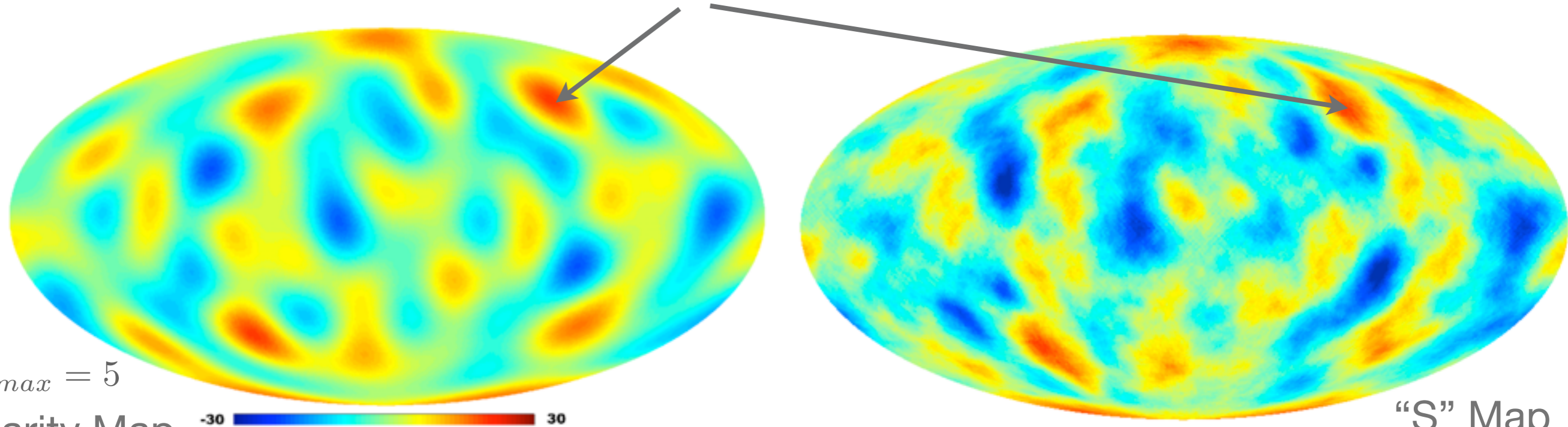
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- ▶ Maximum: $(l, b) \simeq (260^\circ, 60^\circ)$, near “axis of evil” direction. (de Oliveira-Costa et al. 0307282)



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Parity Map

-30 30

“S” Map

Parity - Even/Odd Multipoles

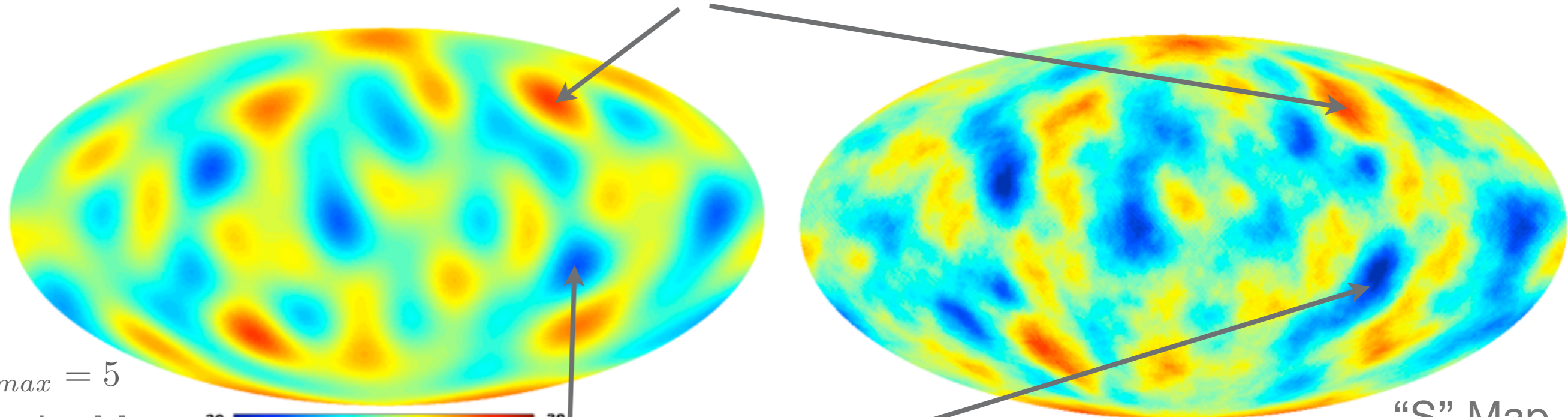
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Parity Map

“S” Map

- ▶ Minimum: $(l, b) \simeq (266^\circ, -19^\circ)$.

Parity - Masking and Reconstruction

(Ben-David, EDK and Itzhaki, ApJ 2012)

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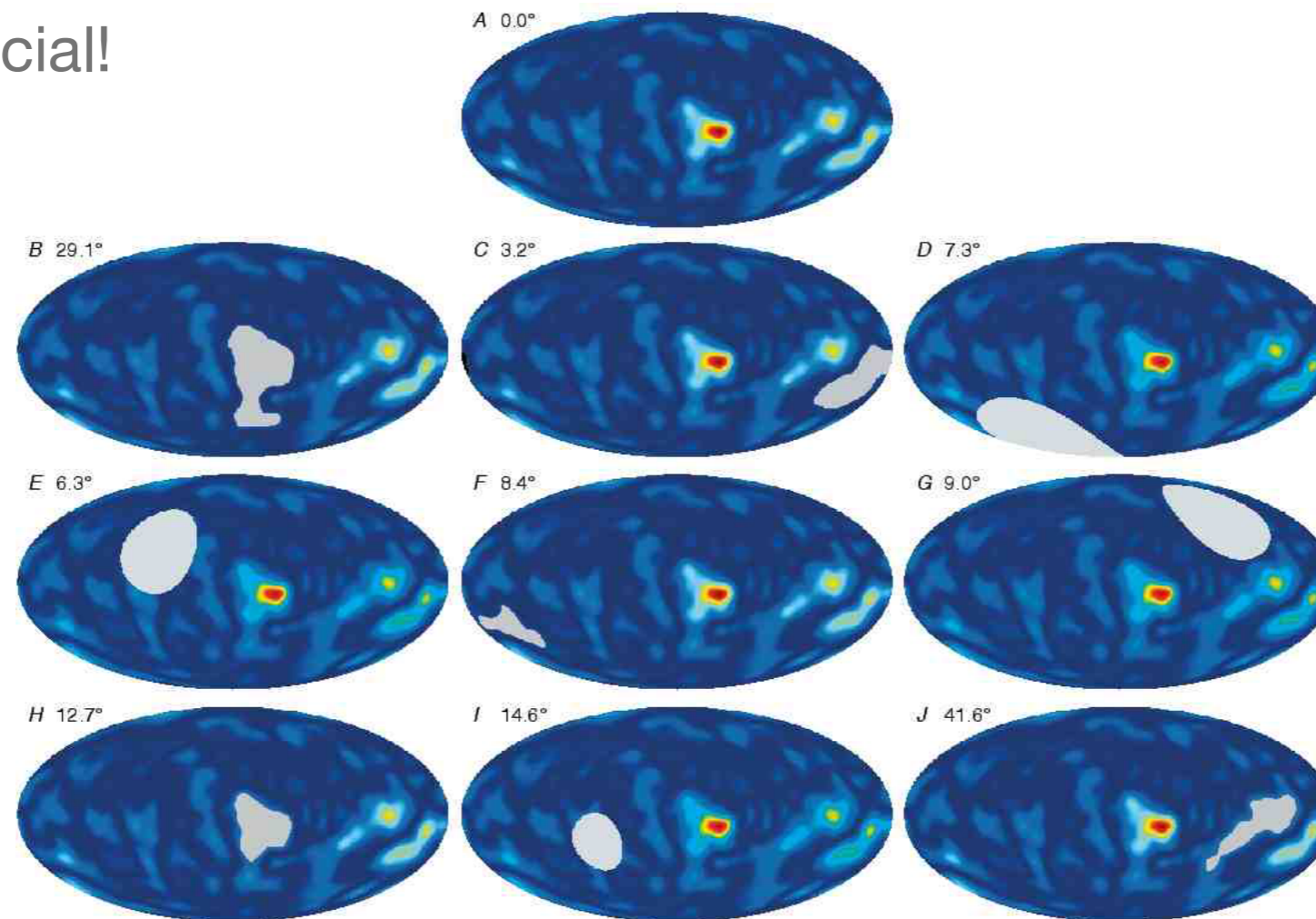
- Masking is crucial!

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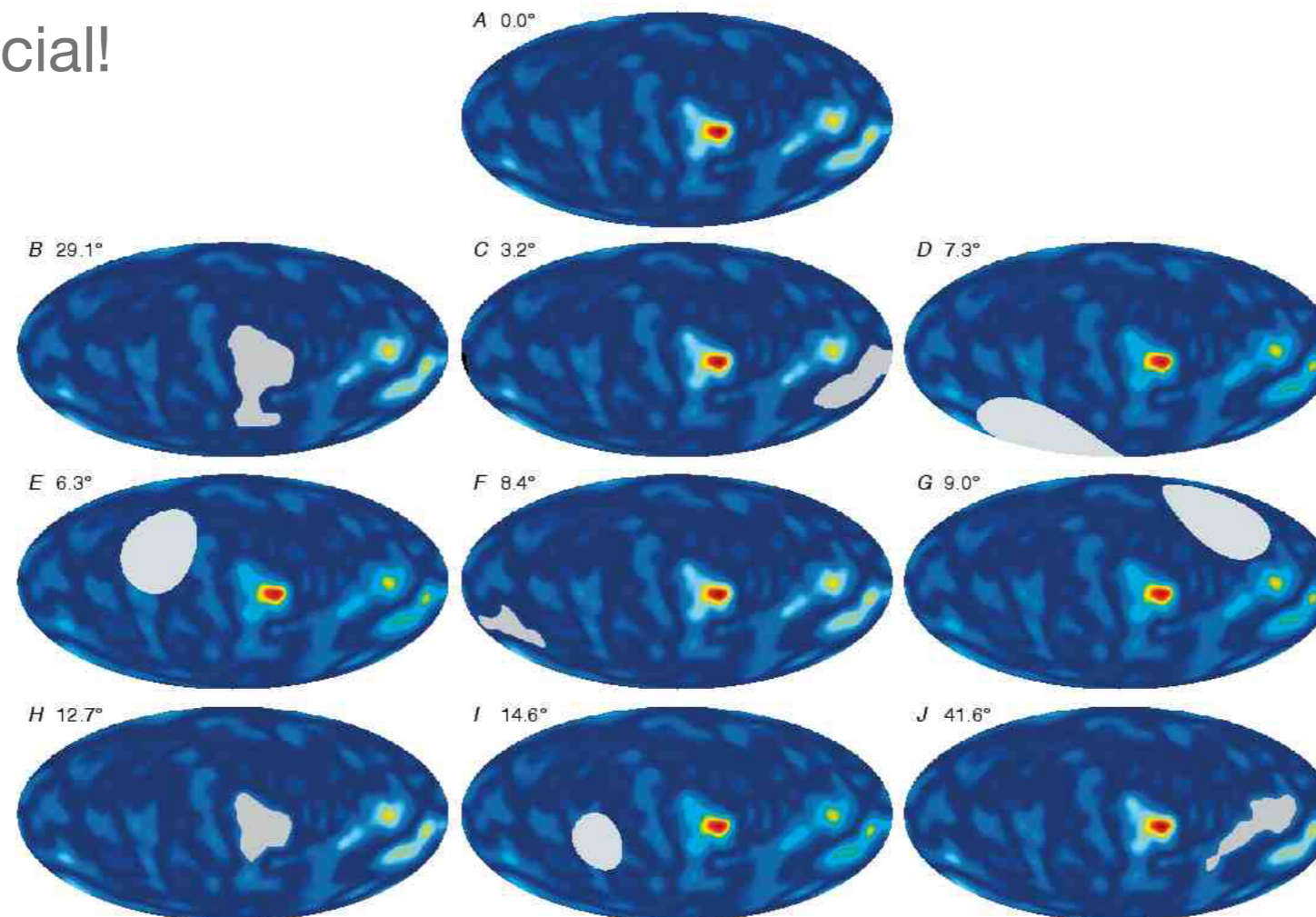
Bennet et al. 2010

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Bennet et al. 2010

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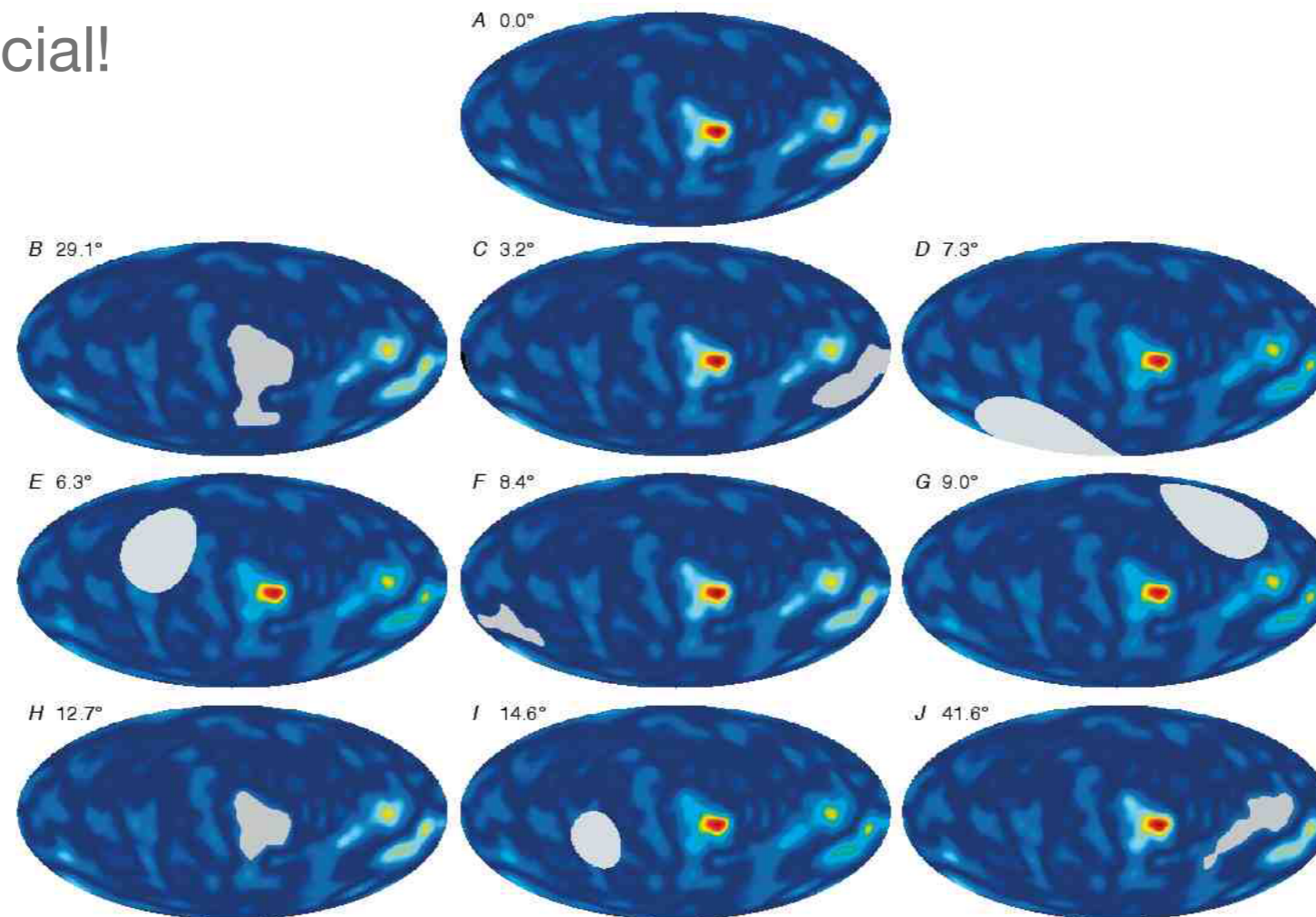
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Efstathiou et al. 2009
Aurich and Lustig 2010)

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Aspen Winter 2012

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Bennet et al. 2010

- a_{lm} Reconstruction:

Decomposition $\mathbf{x} = \mathbf{Y}\mathbf{a} + \mathbf{n}$ where: $\mathbf{Y}_{ij} = Y_{\ell_j m_j}(\hat{\mathbf{r}}_i)$

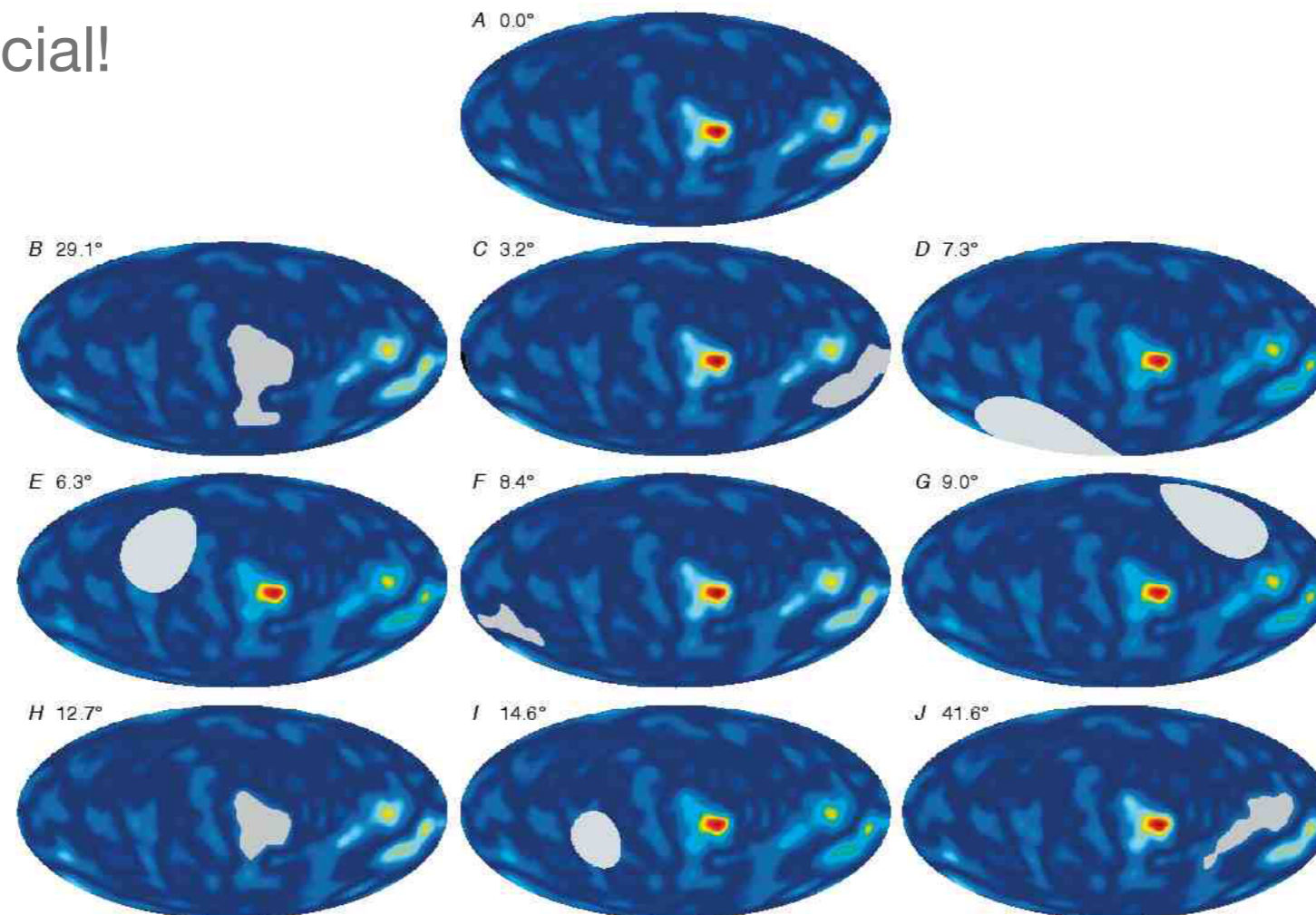
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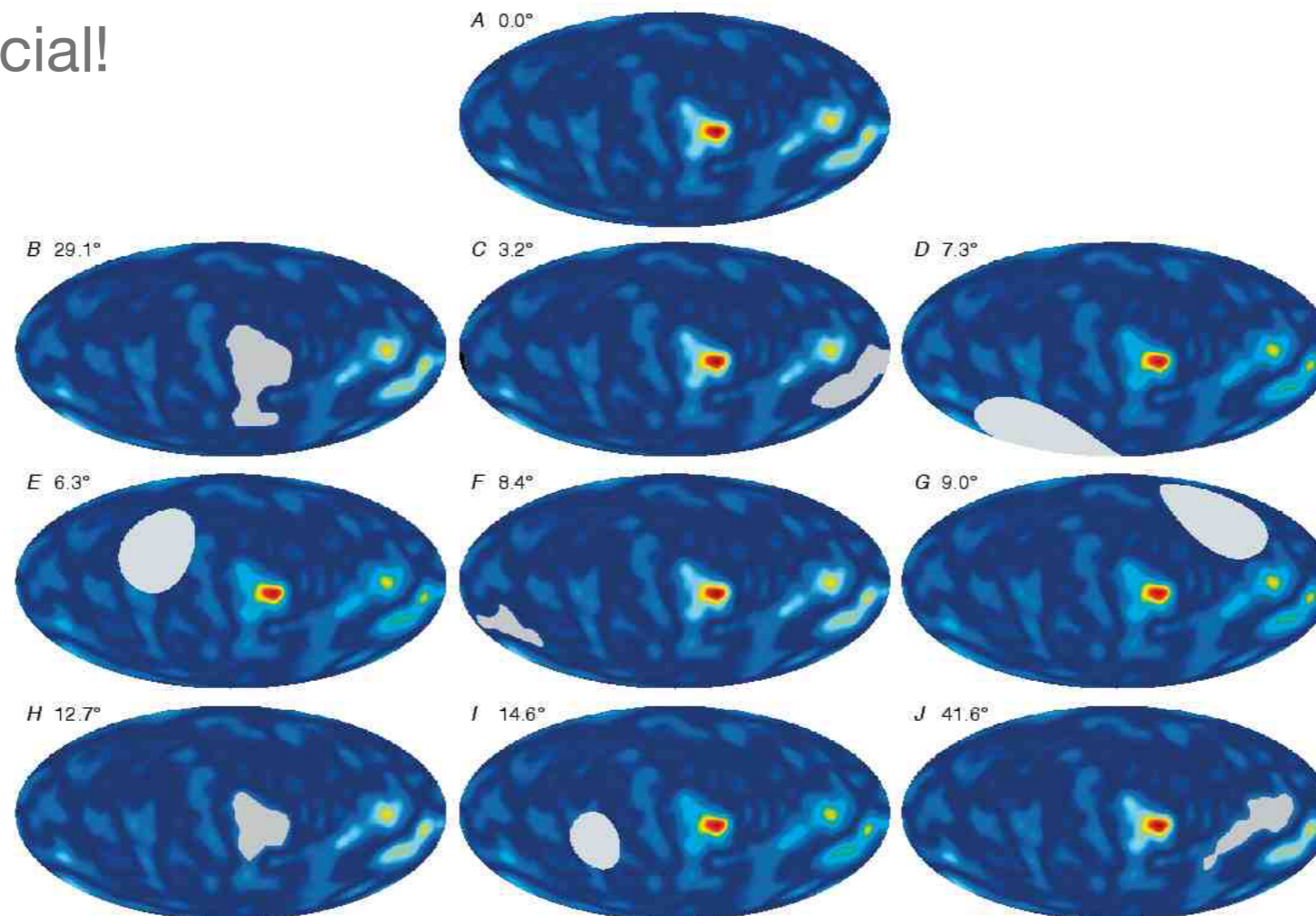
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Parity - Cut-Sky Results

(Ben-David, EDK and Itzhaki, ApJ 2012)

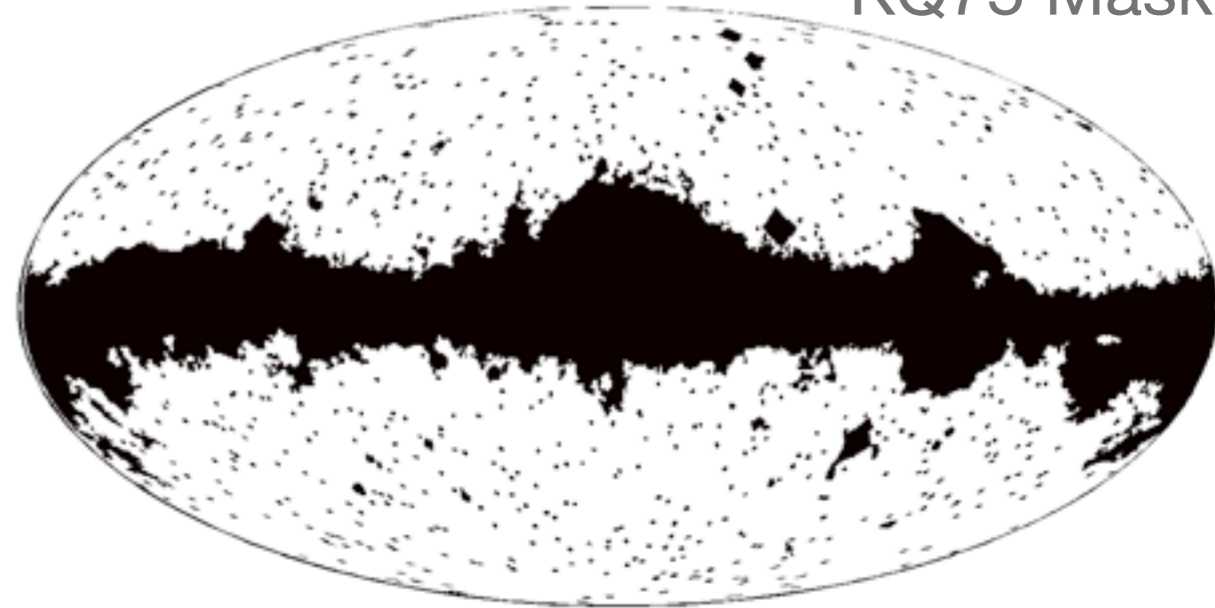
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Parity - Cut-Sky Results

(Ben-David, EDK and Itzhaki, ApJ 2012)

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KQ75 Mask



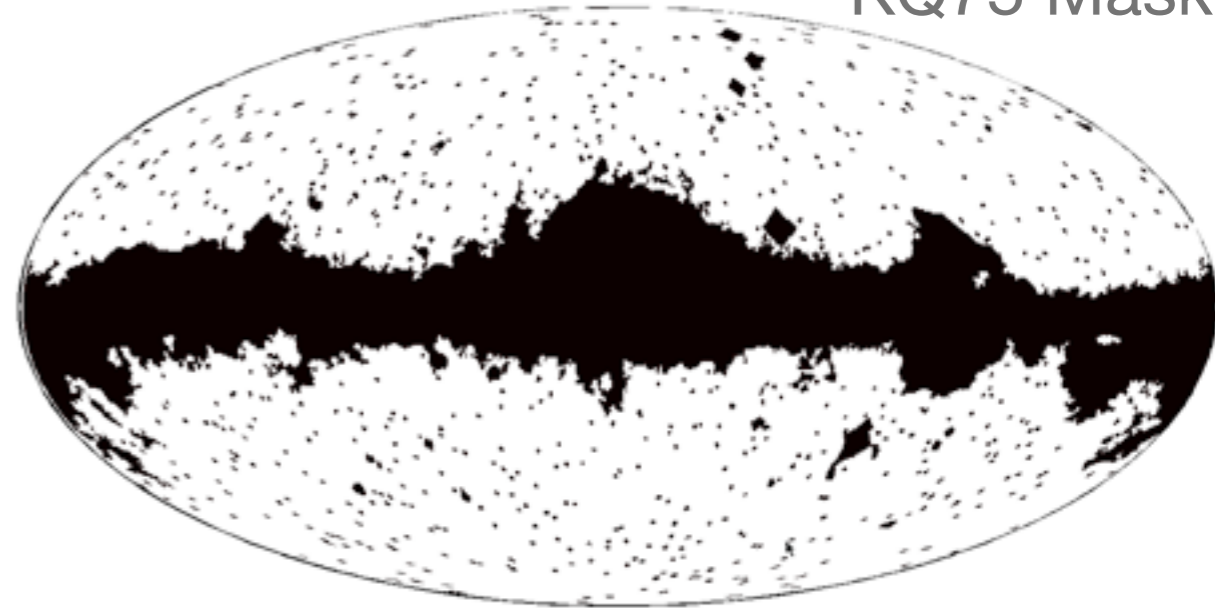
Parity - Cut-Sky Results

(Ben-David, EDK and Itzhaki, ApJ 2012)

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KQ75 Mask

- Mask out 2.5%-10% outlying pixels:

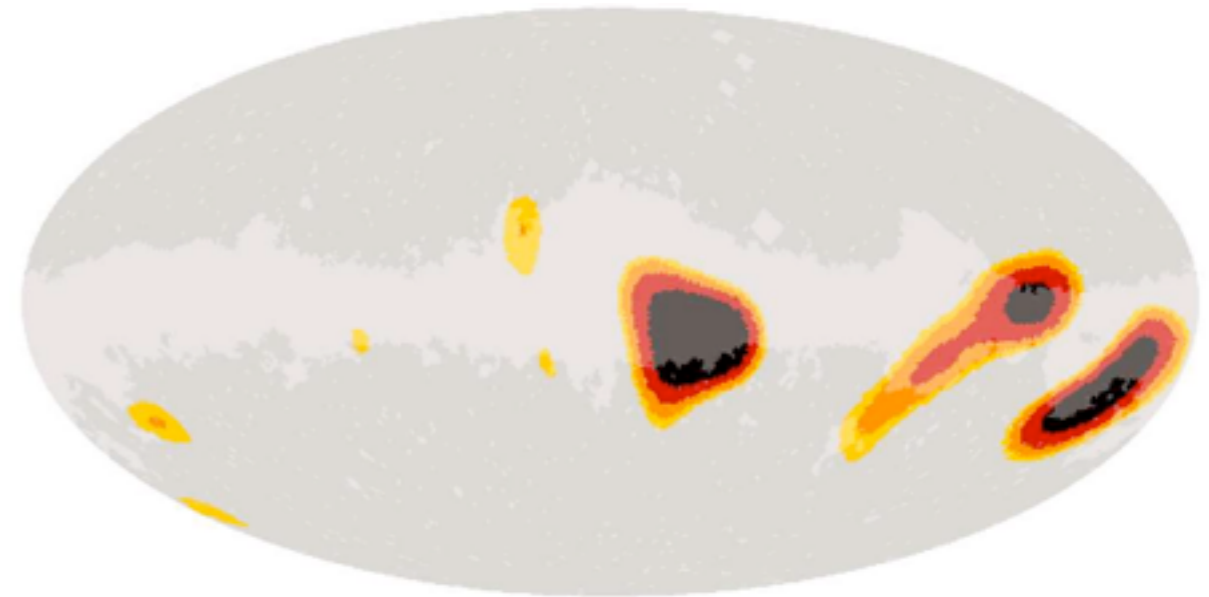


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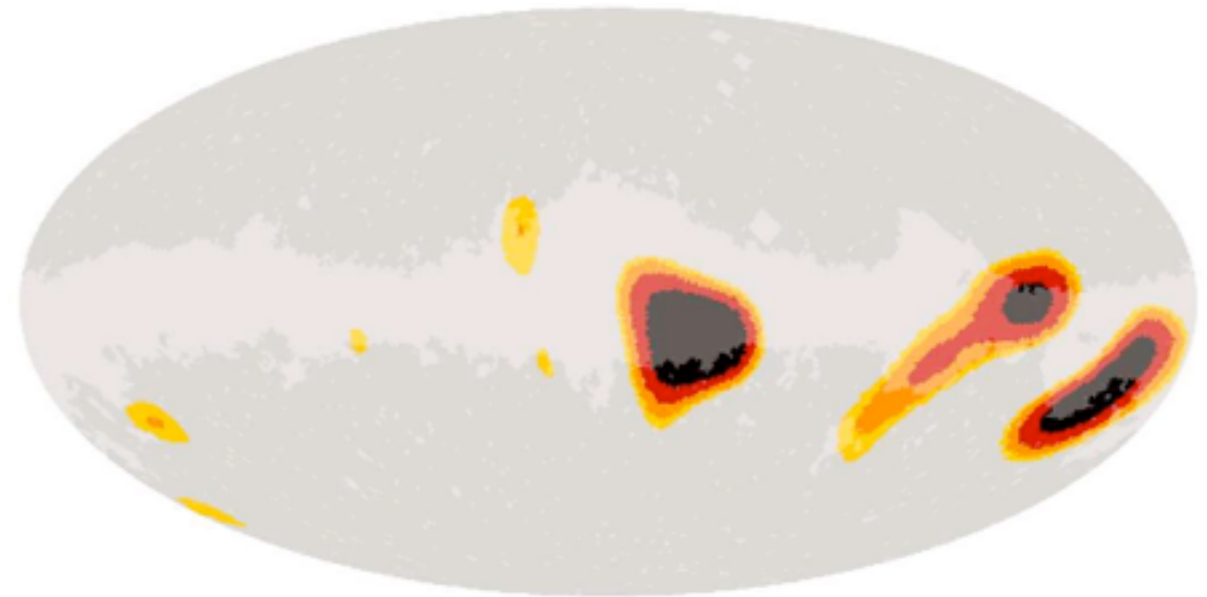


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 - ▶ Masked regions near the Galactic plane.

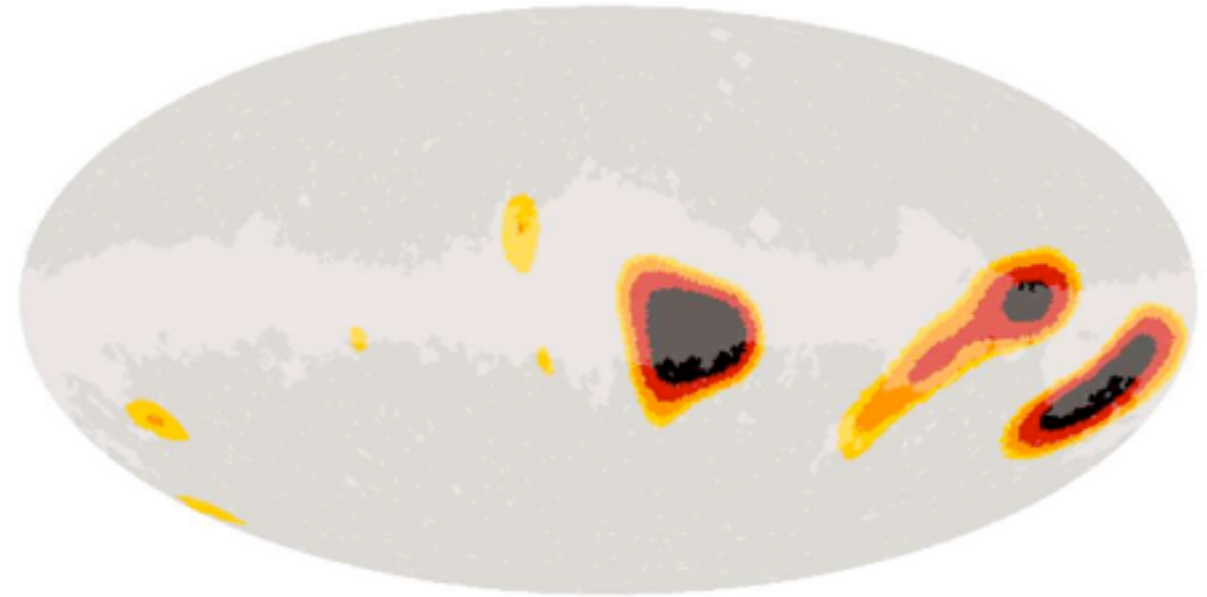


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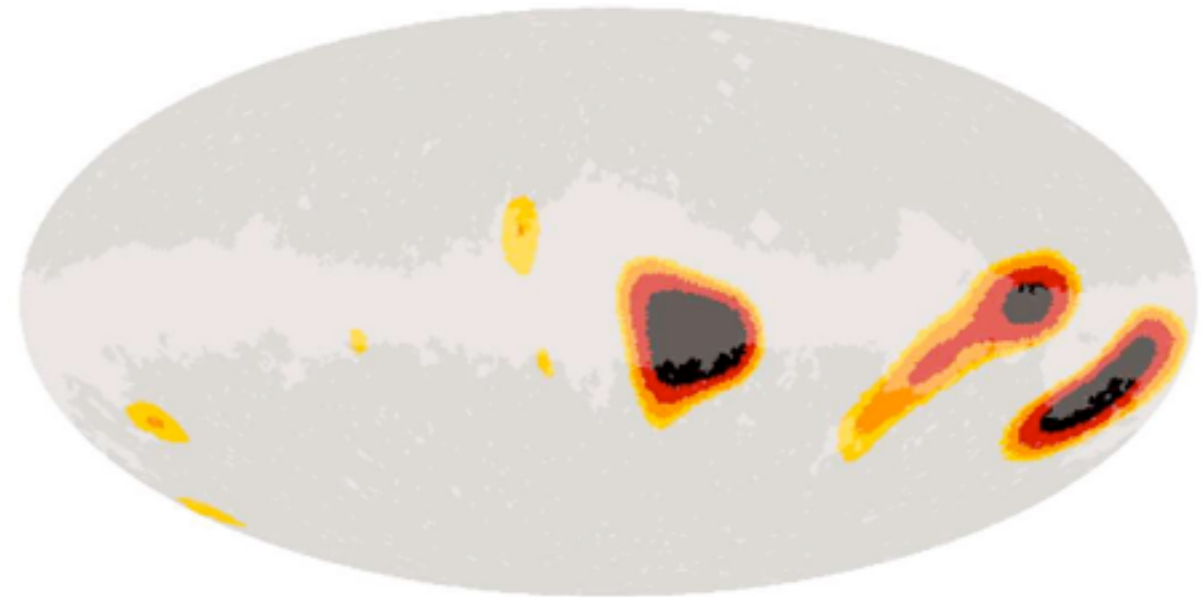


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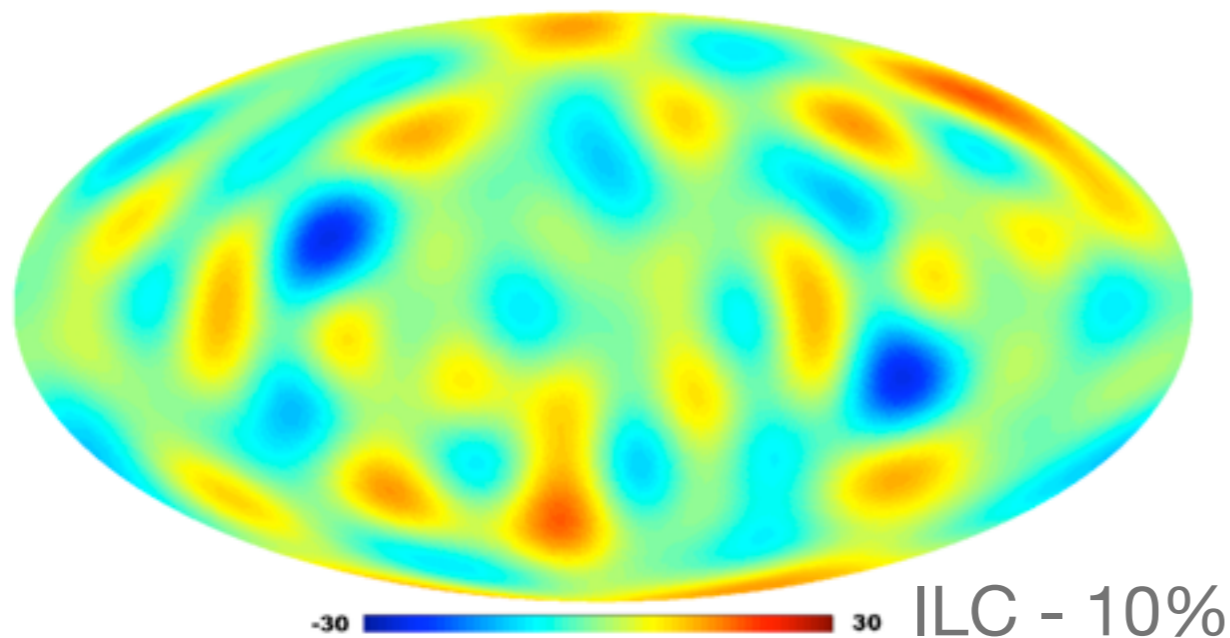
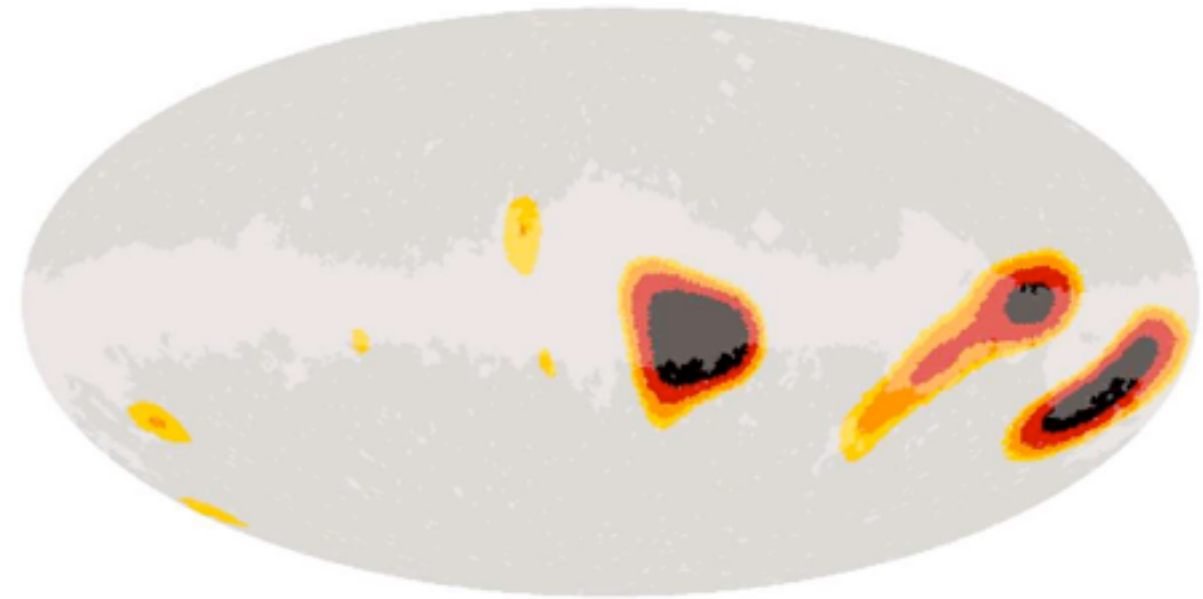


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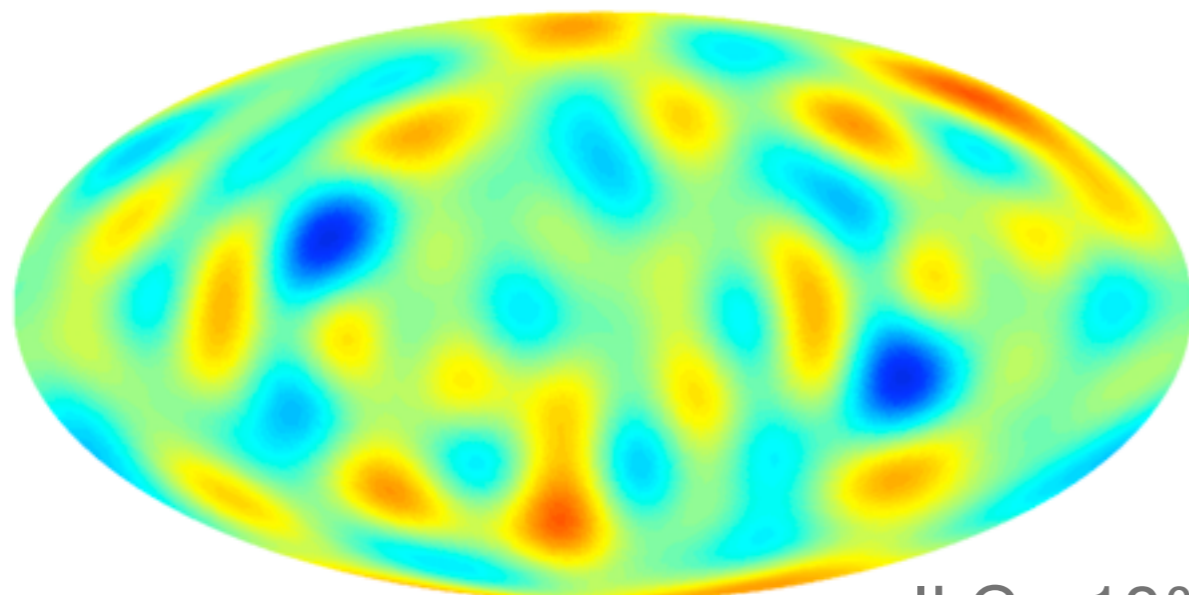
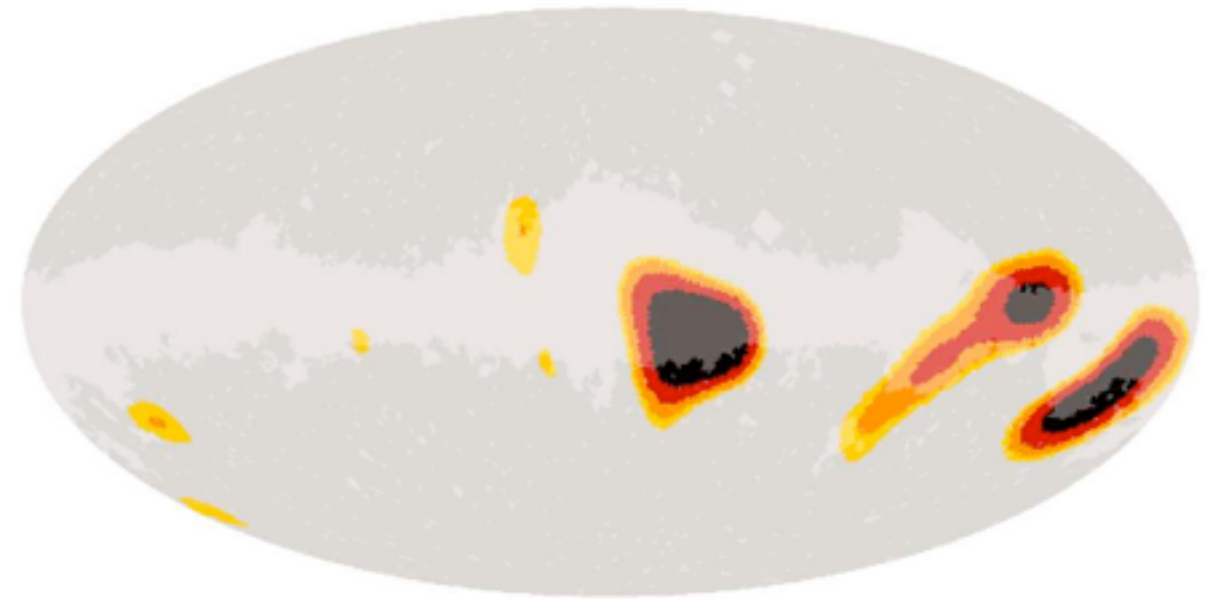


Parity - Cut-Sky Results

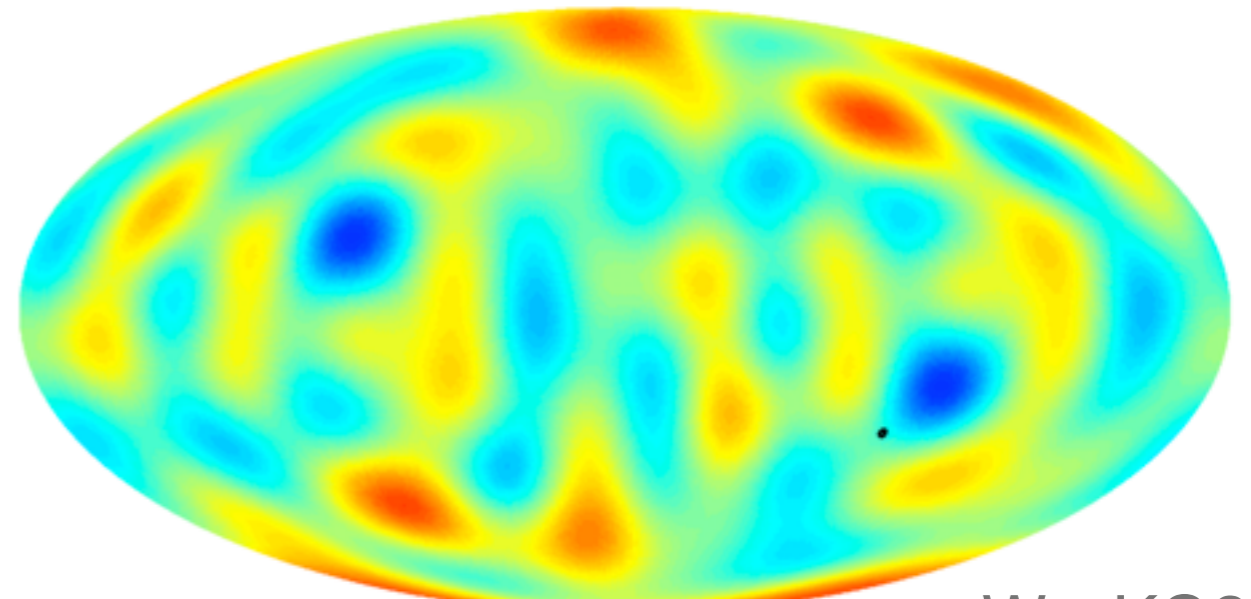
(Ben-David, EDK and Itzhaki, ApJ 2012)


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-30  30 ILC - 10%



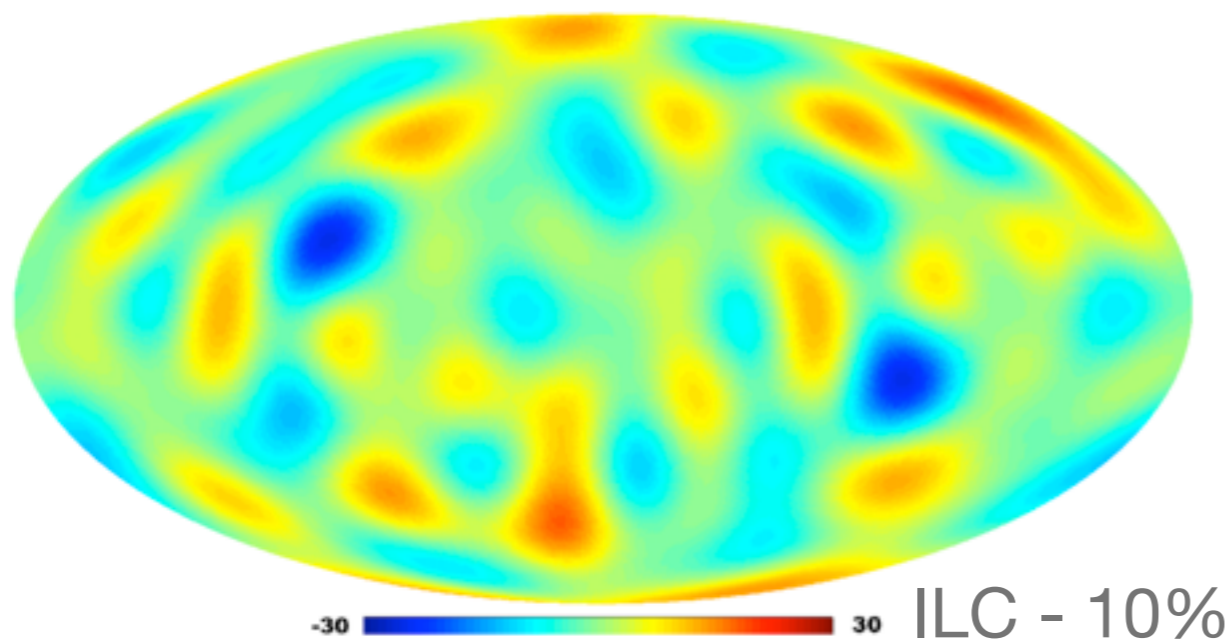
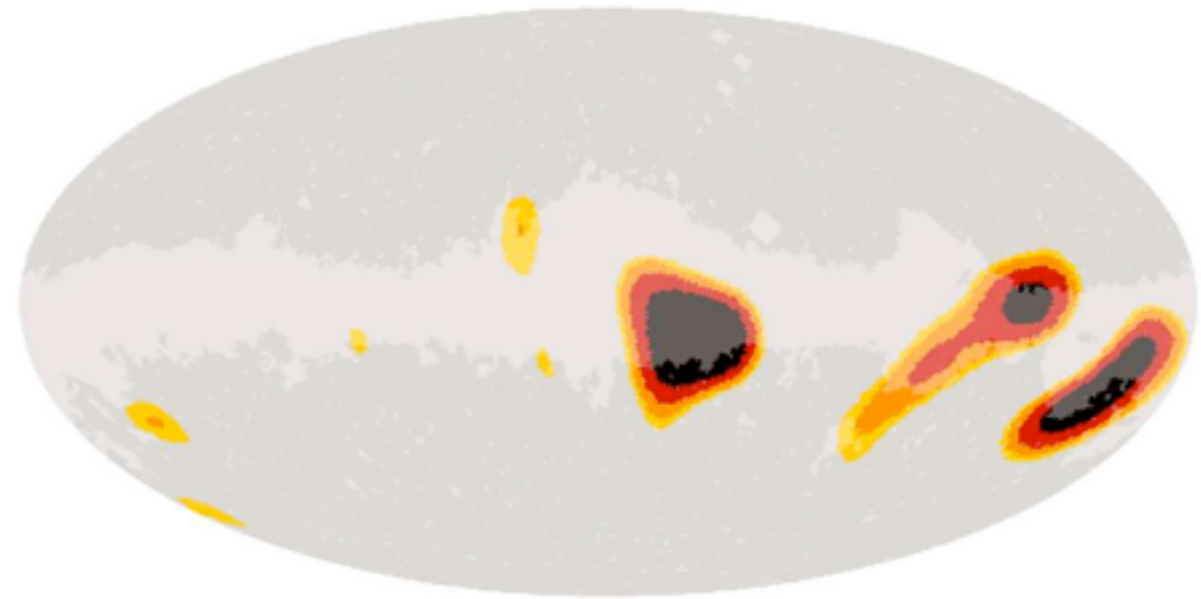
-30  30 W - KQ85

Parity - Cut-Sky Results

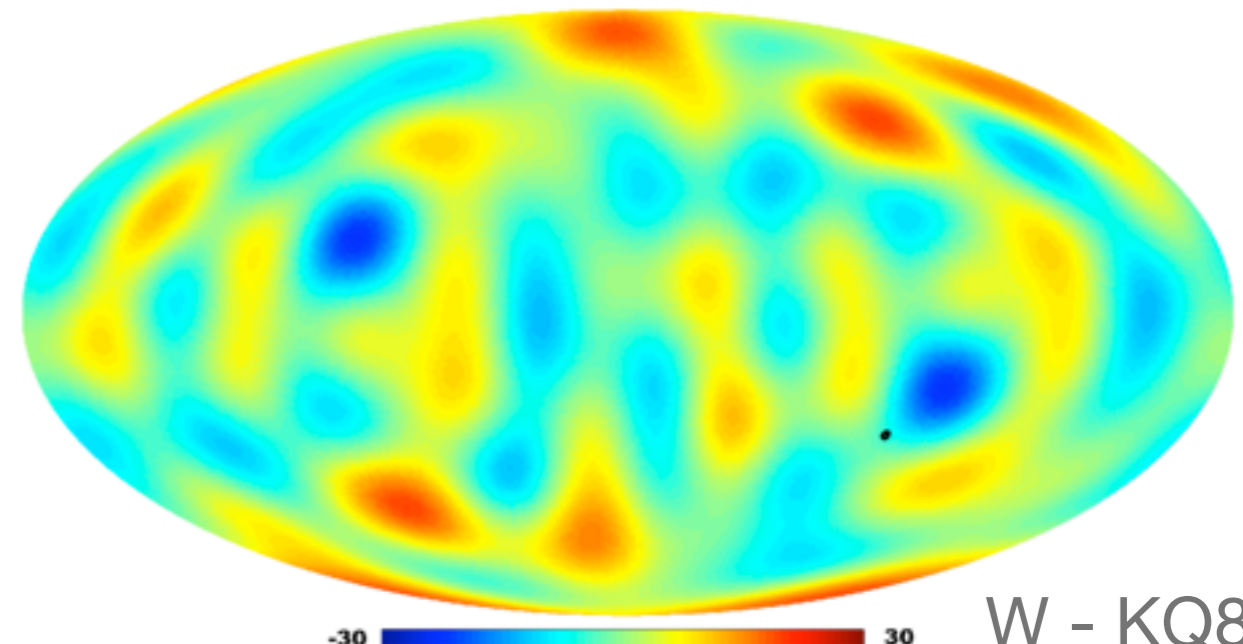
(Ben-David, EDK and Itzhaki, ApJ 2012)

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ILC - 10%



W - KQ85

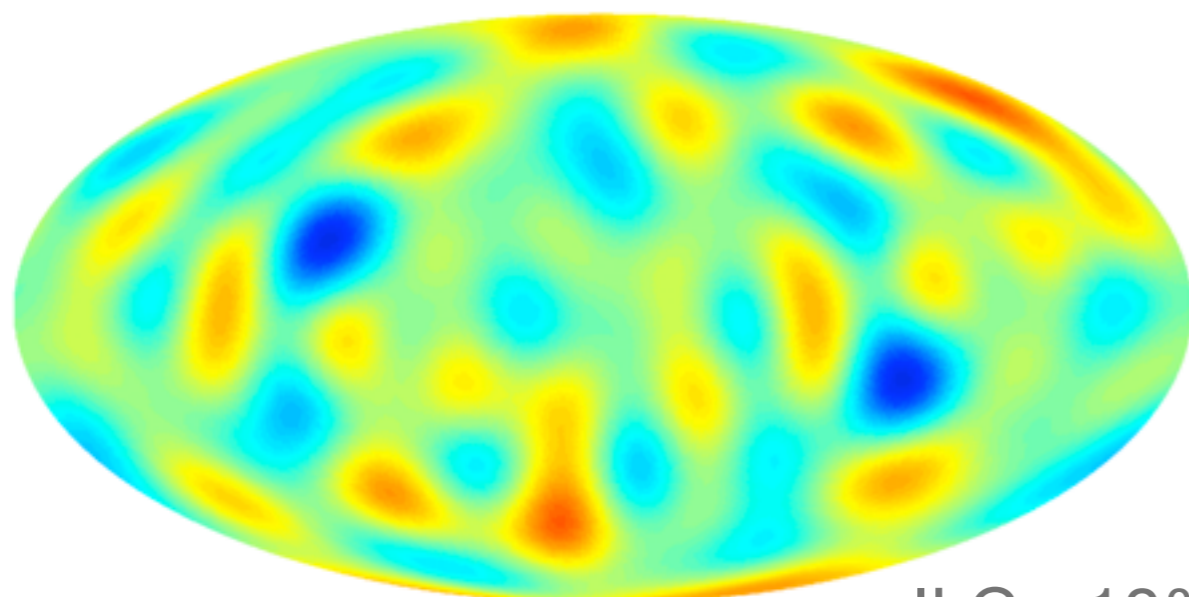
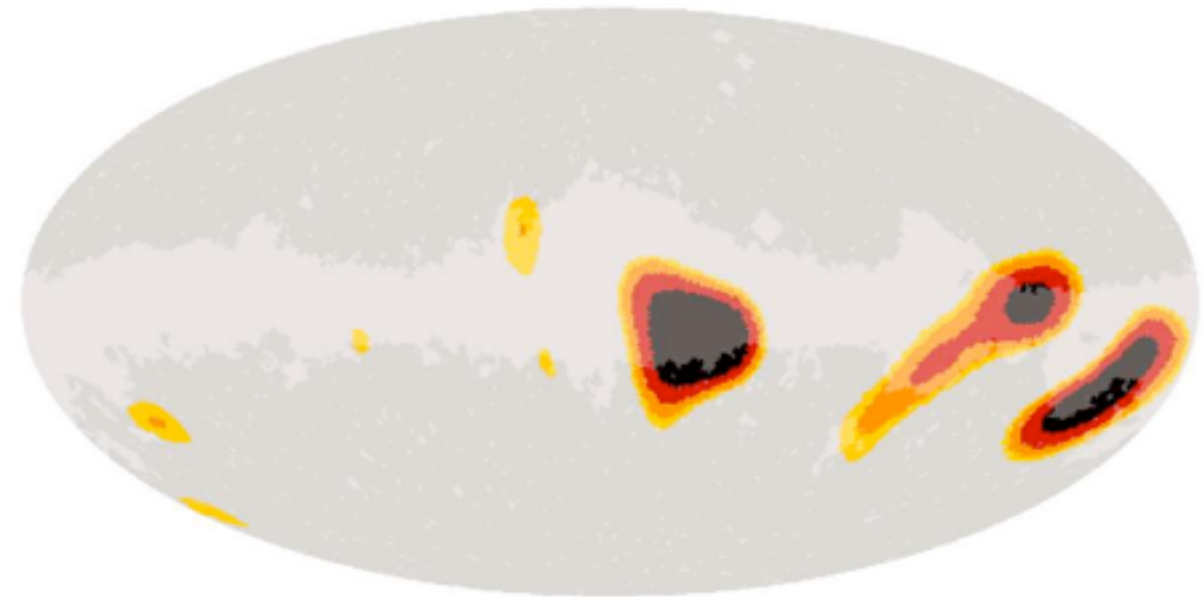
- ▶ Even parity insignificant.

Parity - Cut-Sky Results

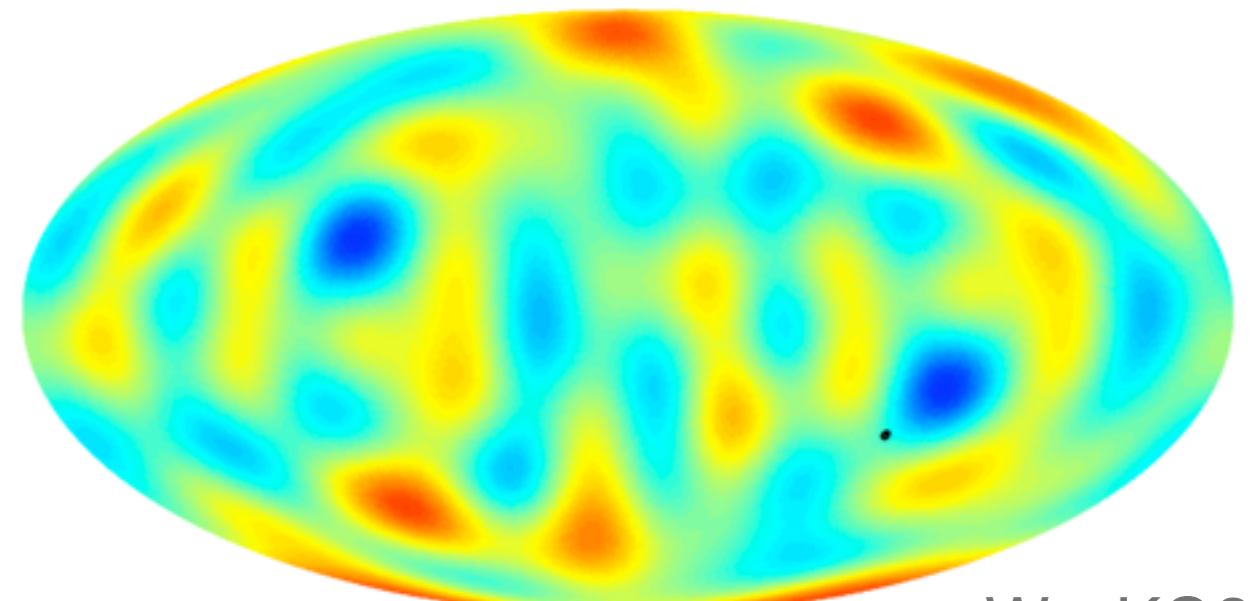
(Ben-David, EDK and Itzhaki, ApJ 2012)

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- Mask out 2.5%-10% outlying pixels:
 - ▶ Masked regions near the Galactic plane.
 - ▶ Procedure applicable to random maps.
- With all masks (including KQ75/85), we find:



-30 30 ILC - 10%



-30 30 W - KQ85

- ▶ Odd parity at $(l, b) \simeq (264^\circ, -18^\circ)$ is significant: $\sim 0.01\%$ for $\ell_{max} = 5 - 7$.

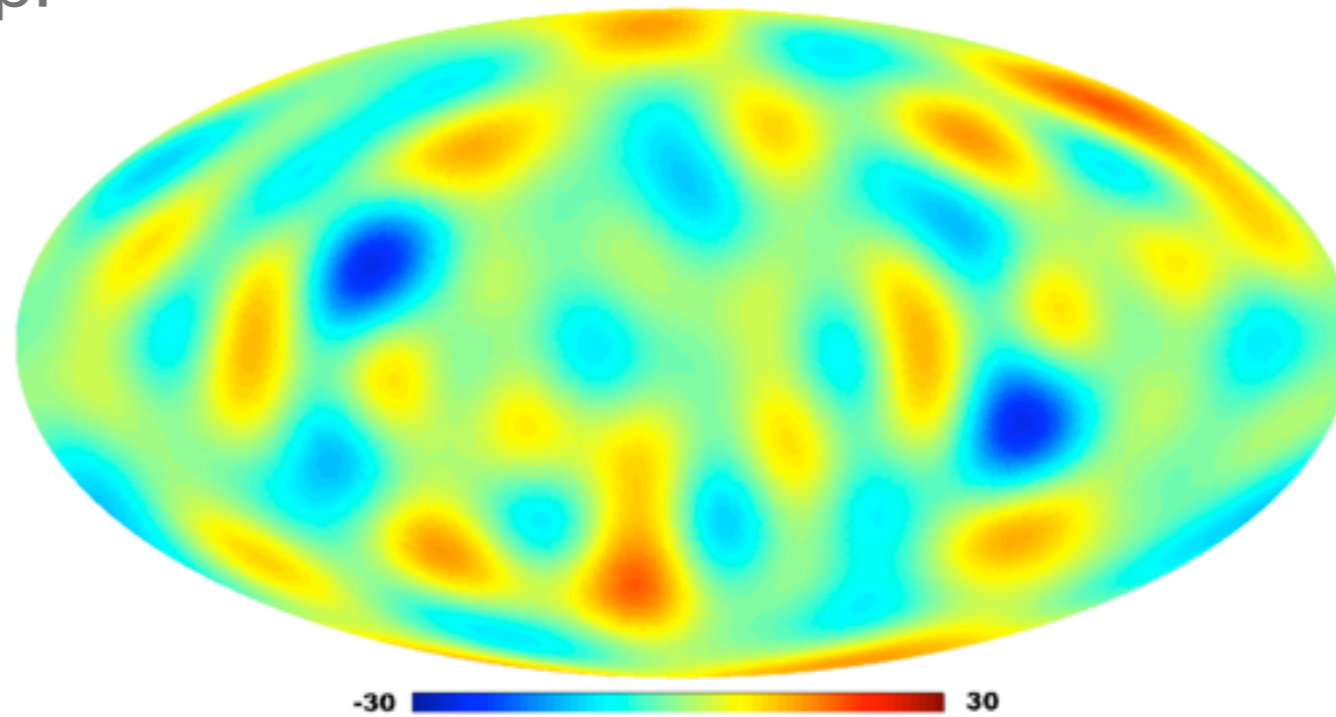
Parity - COBE-DMR

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Parity - COBE-DMR

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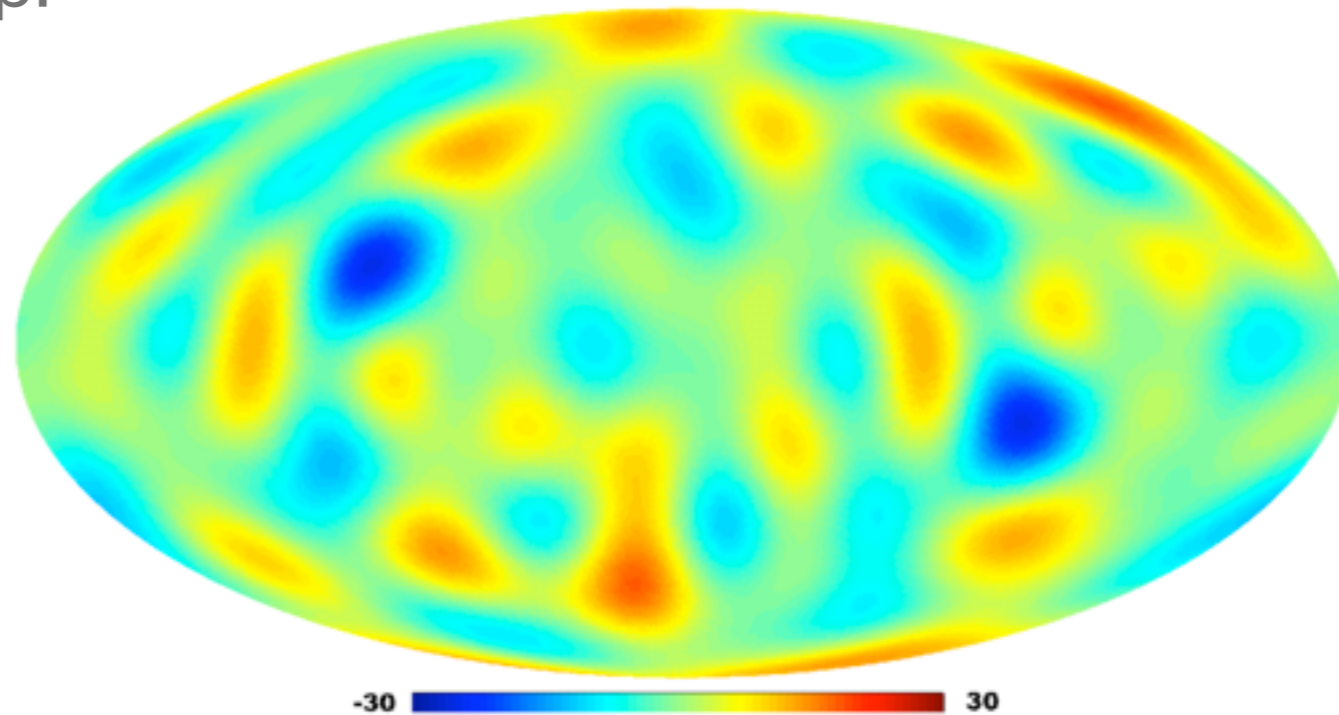
- WMAP parity map:



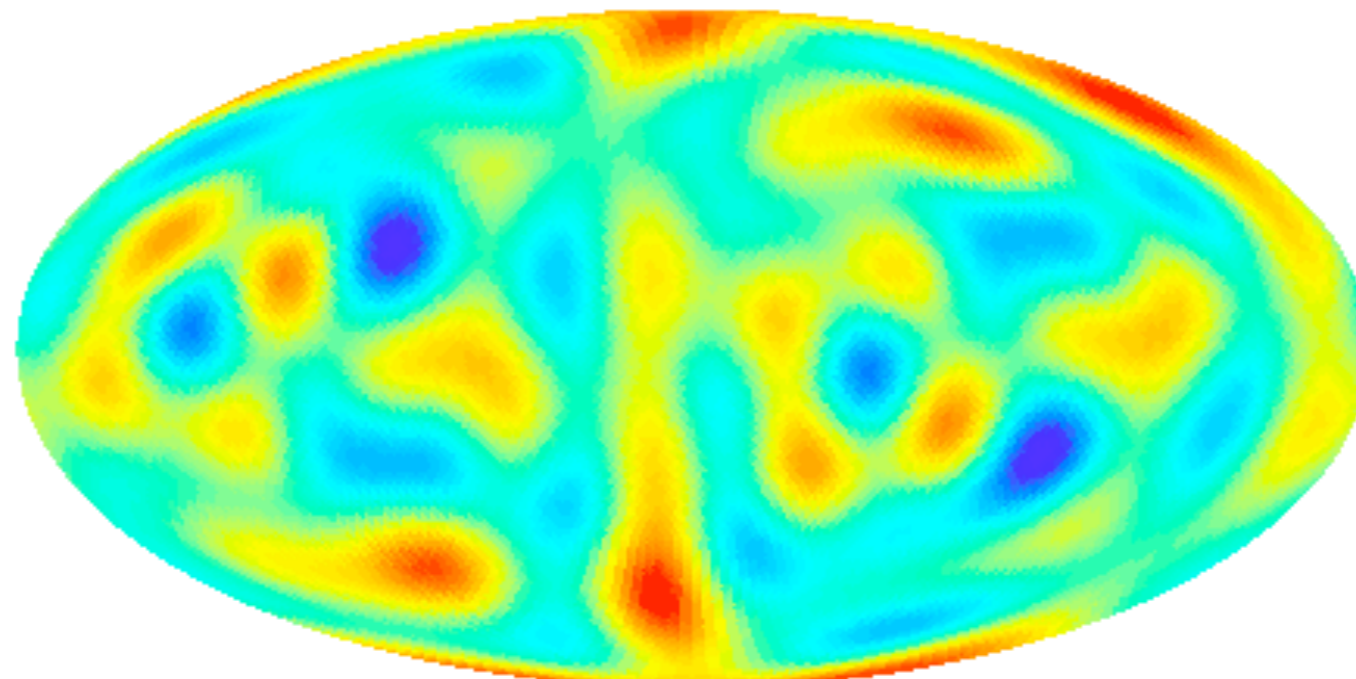
Parity - COBE-DMR

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- WMAP parity map:



- COBE parity map:



Simulation vs. data: Moving PIP

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Simulation vs. data: Moving PIP

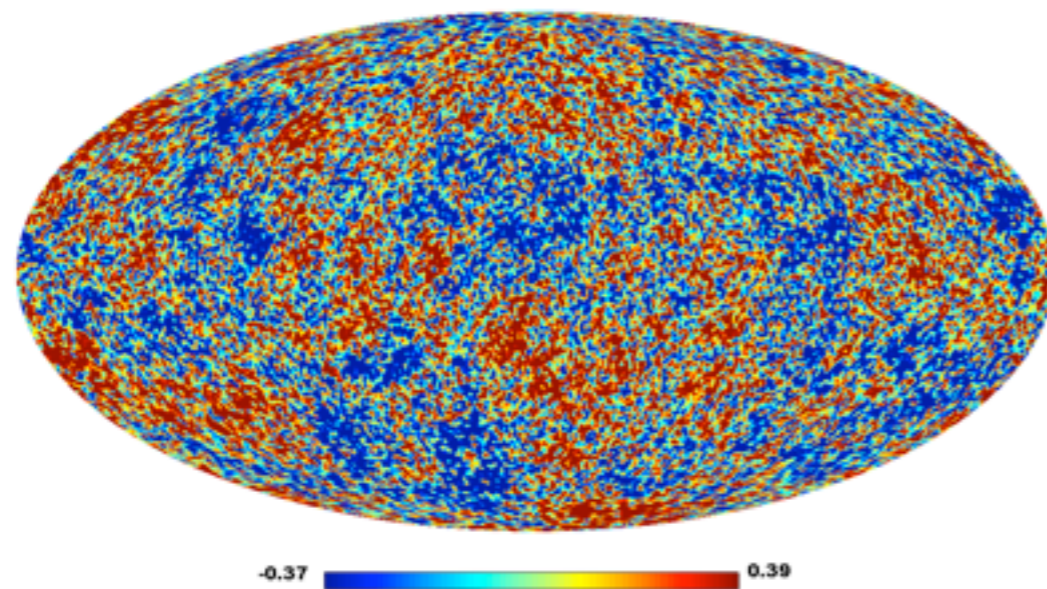
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Simulating LCDM + a moving PIP (with fine-tuned location and velocity):

Simulation vs. data: Moving PIP

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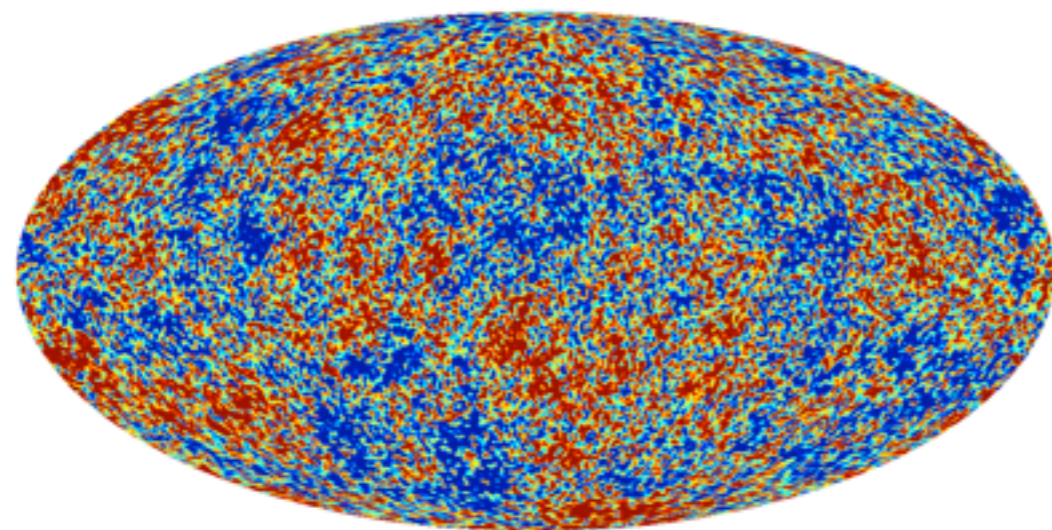
Simulating LCDM + a moving PIP (with fine-tuned location and velocity):



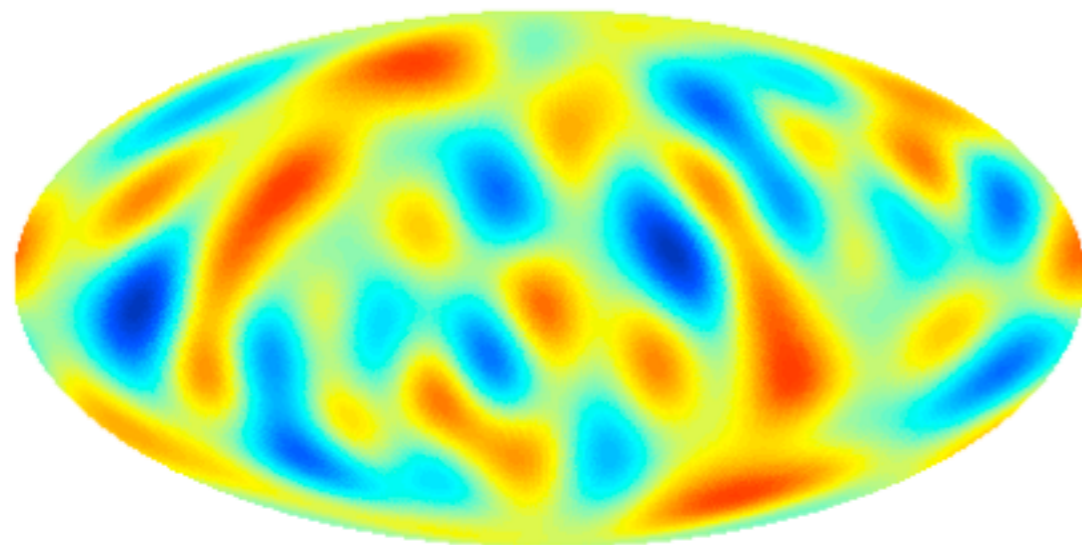
Simulation vs. data: Moving PIP

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Simulating LCDM + a moving PIP (with fine-tuned location and velocity):



-0.37 0.39

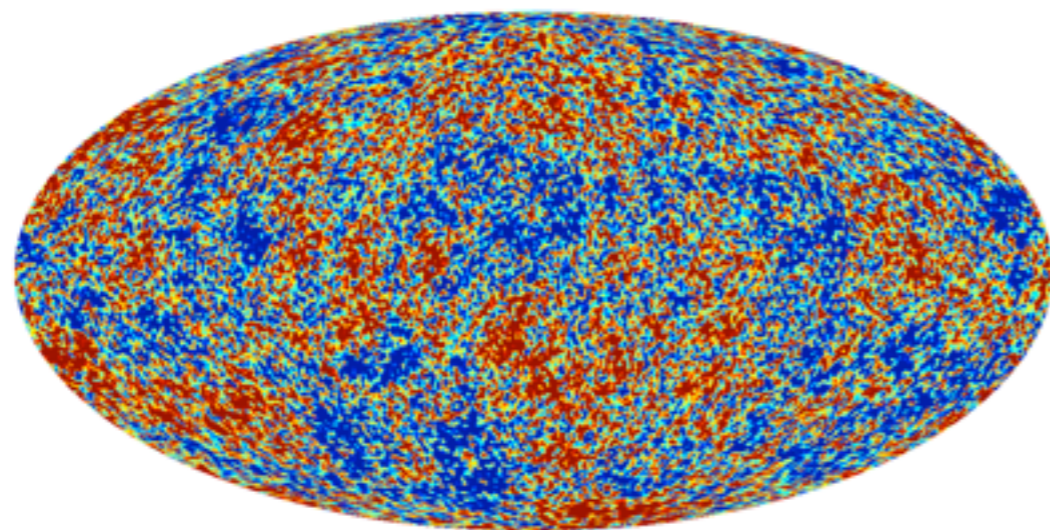


-23.00 21.00

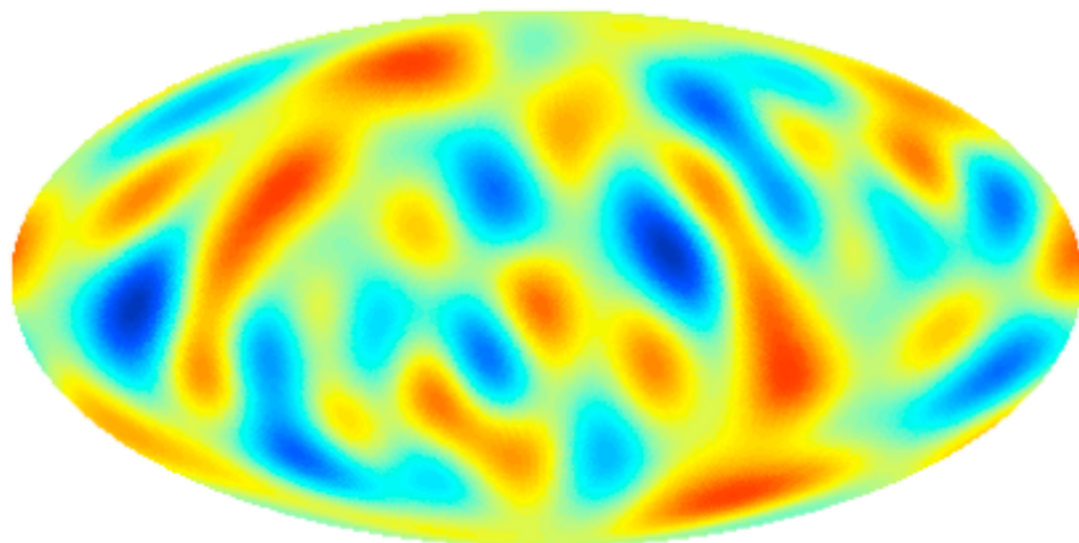
Simulation vs. data: Moving PIP

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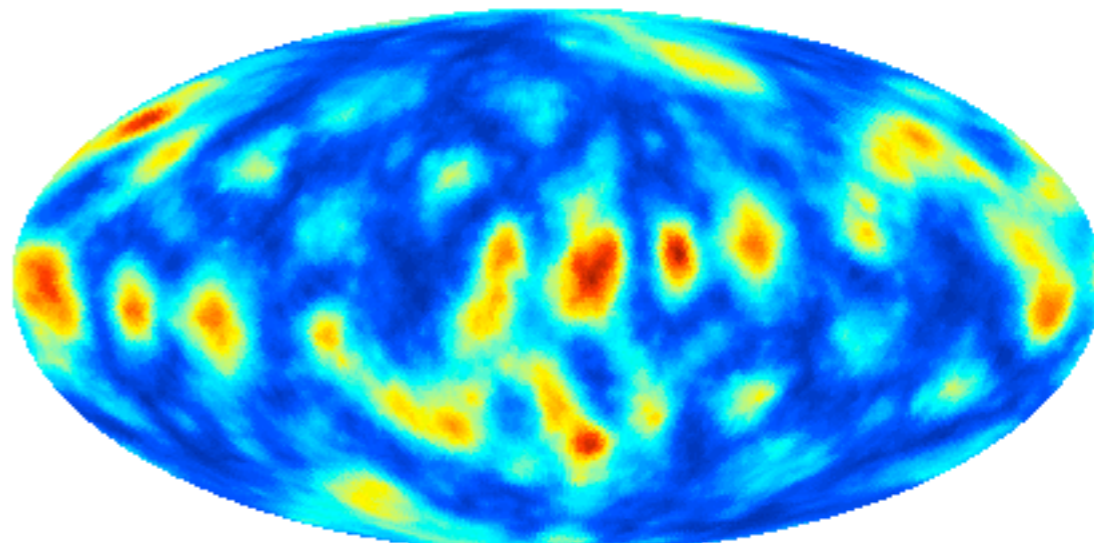
Simulating LCDM + a moving PIP (with fine-tuned location and velocity):



-0.37 0.39



-23.00 21.00

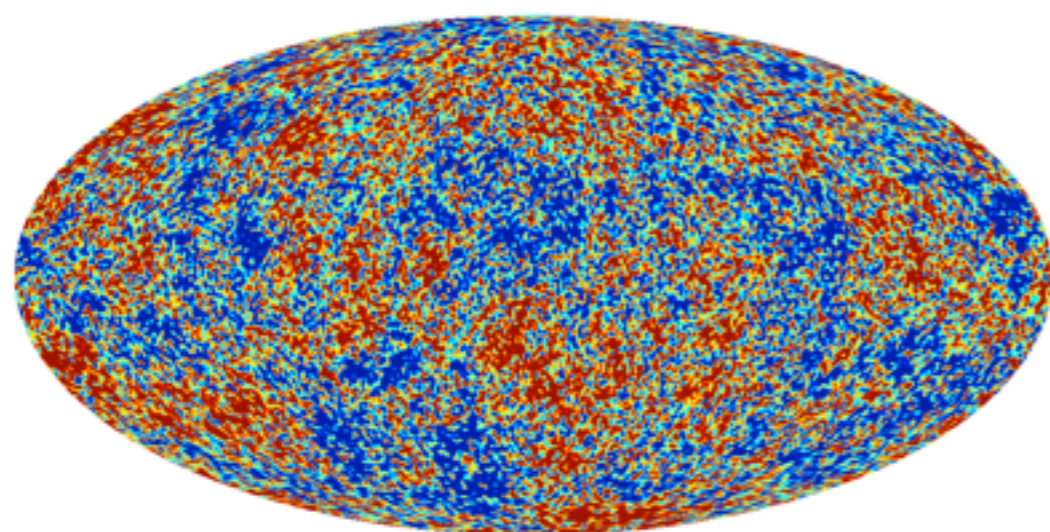


0.05 3.95

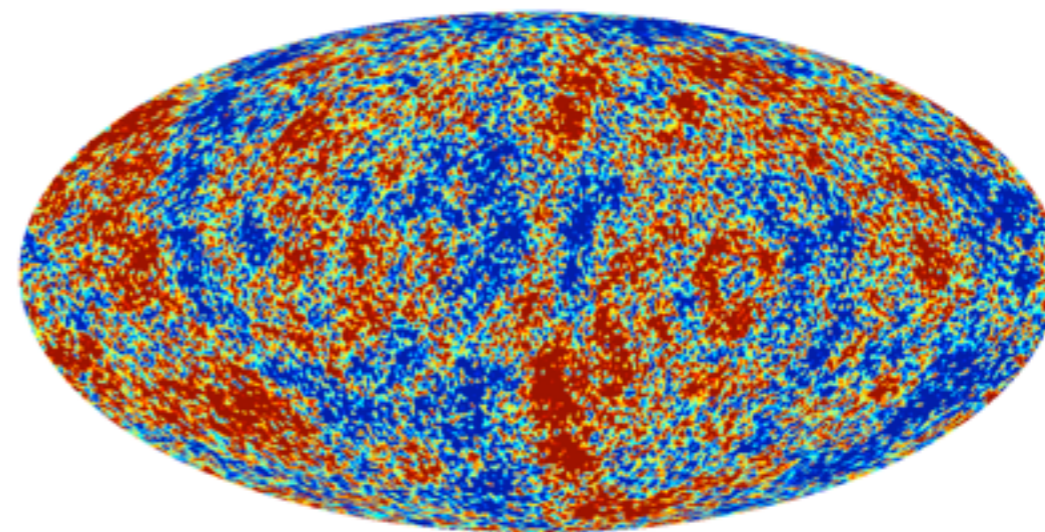
Simulation vs. data: Moving PIP

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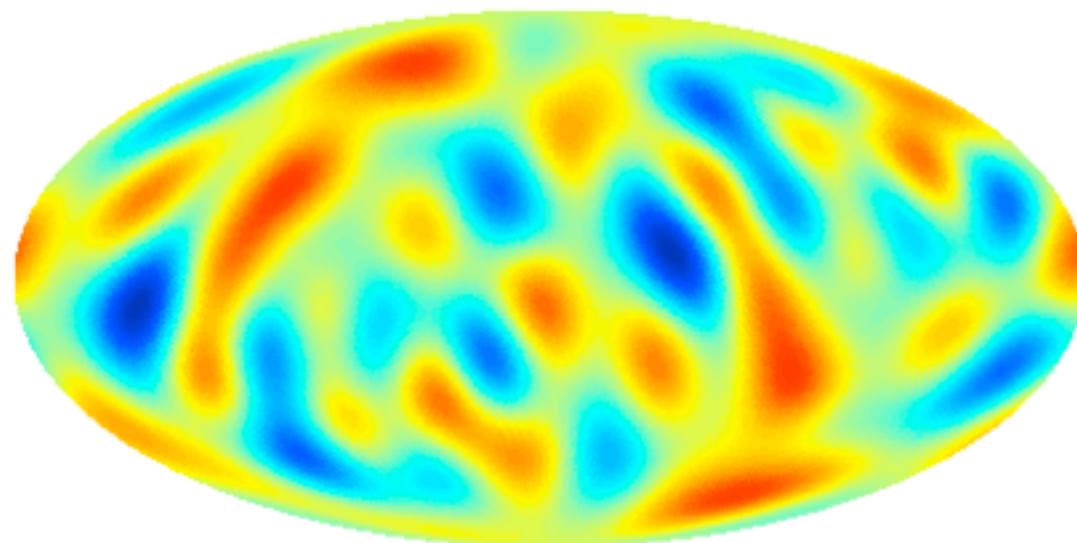
Simulating LCDM + a moving PIP (with fine-tuned location and velocity):



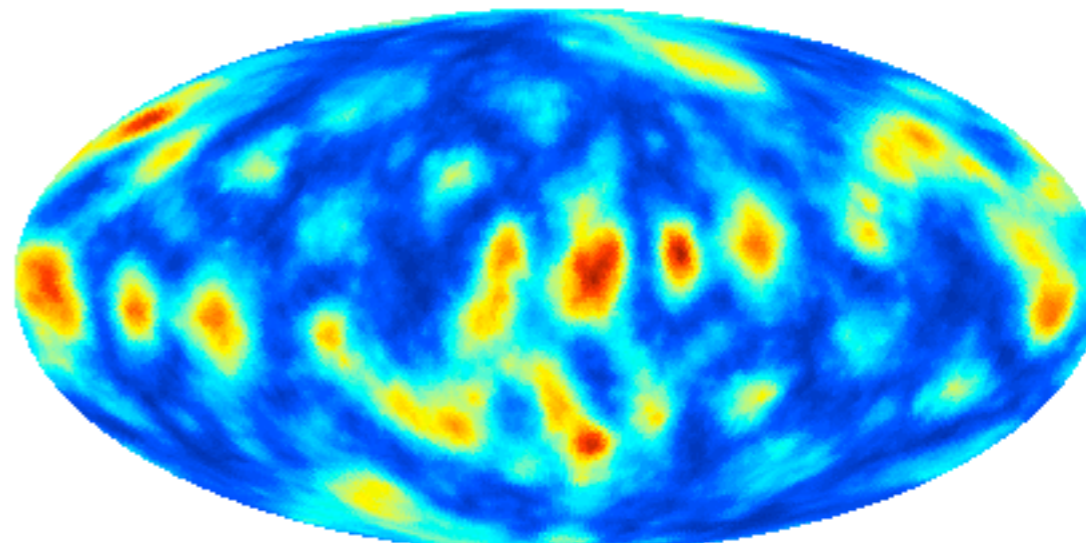
-0.37 0.39



-0.37 0.44



-23.00 21.00

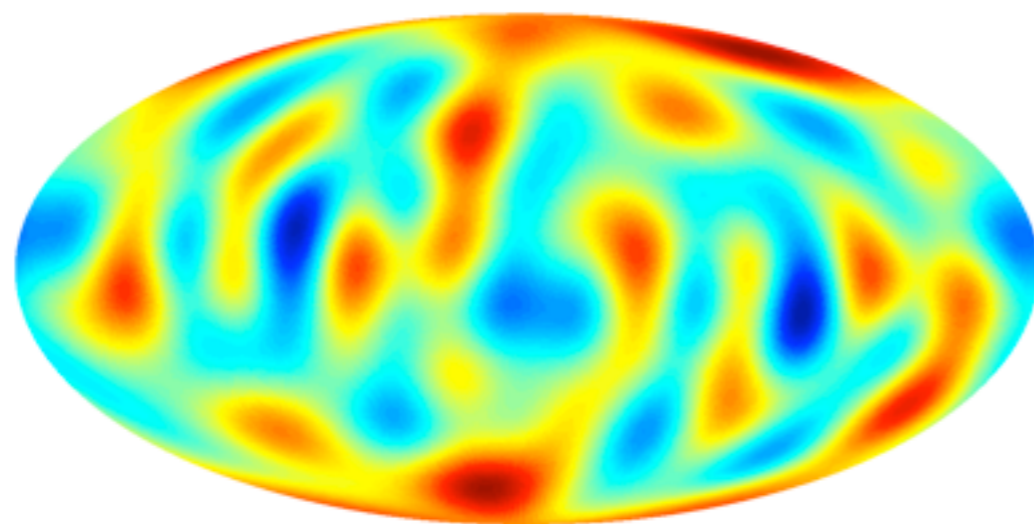


0.05 3.95

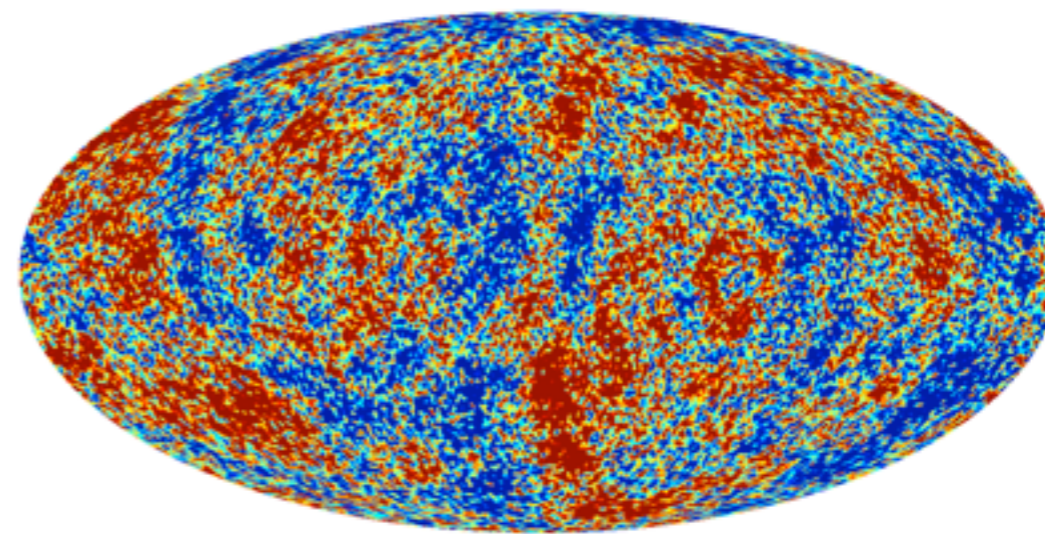
Simulation vs. data: Moving PIP

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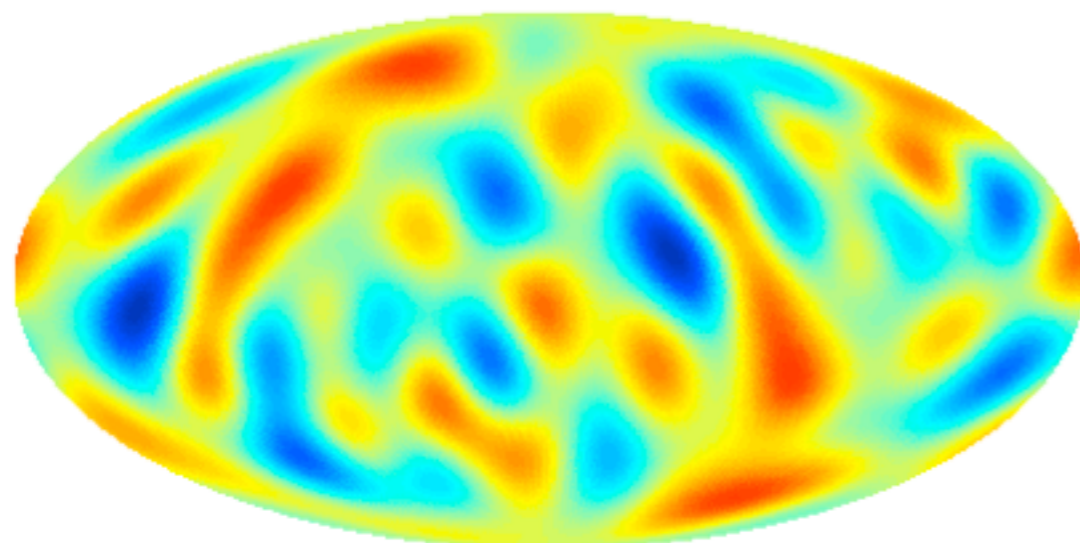
Simulating LCDM + a moving PIP (with fine-tuned location and velocity):



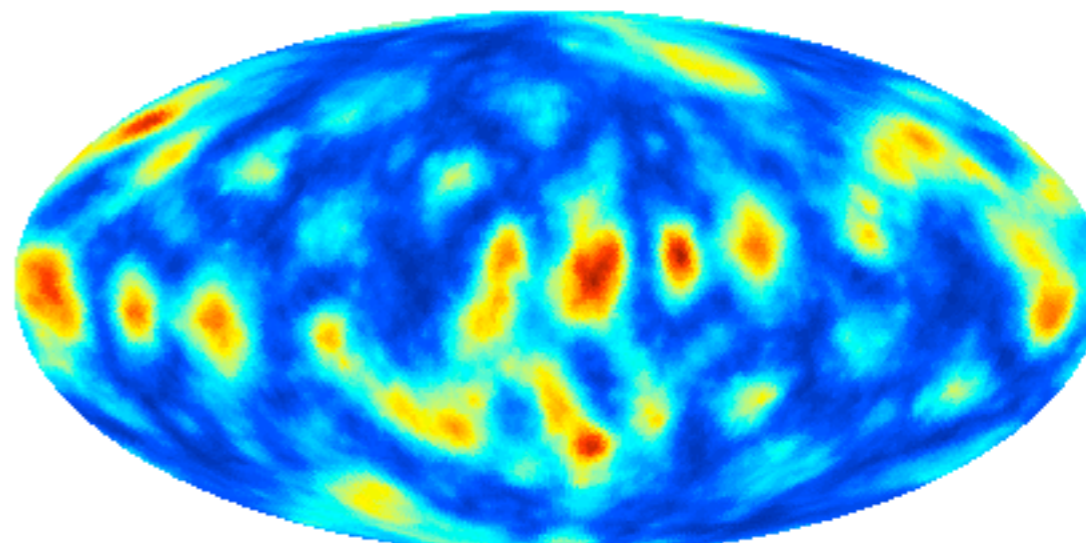
-23.23 21.35



-0.37 0.44



-23.00 21.00

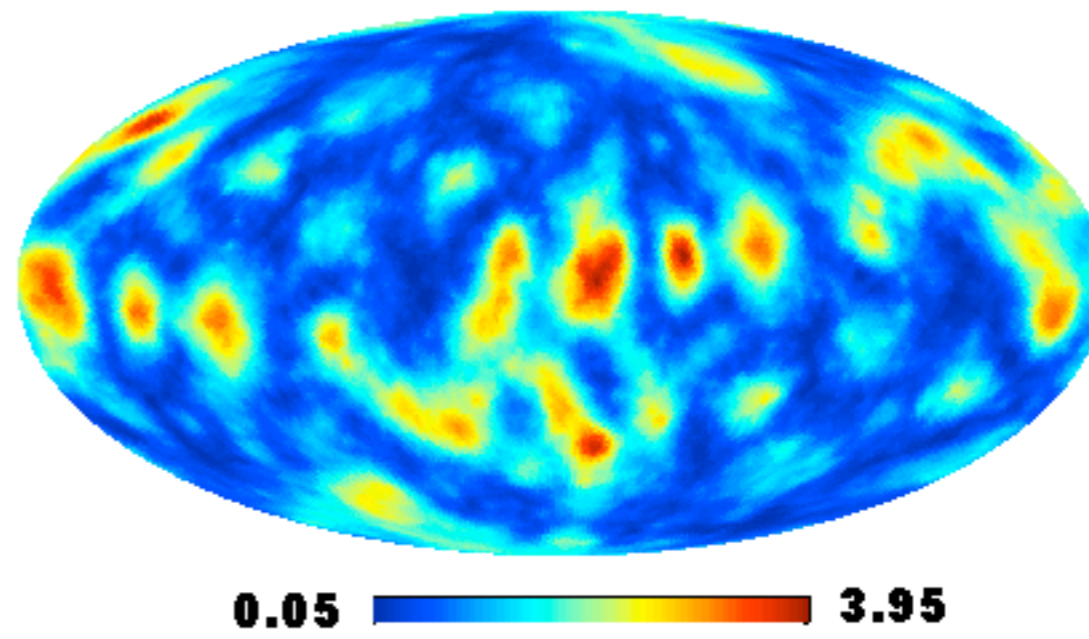
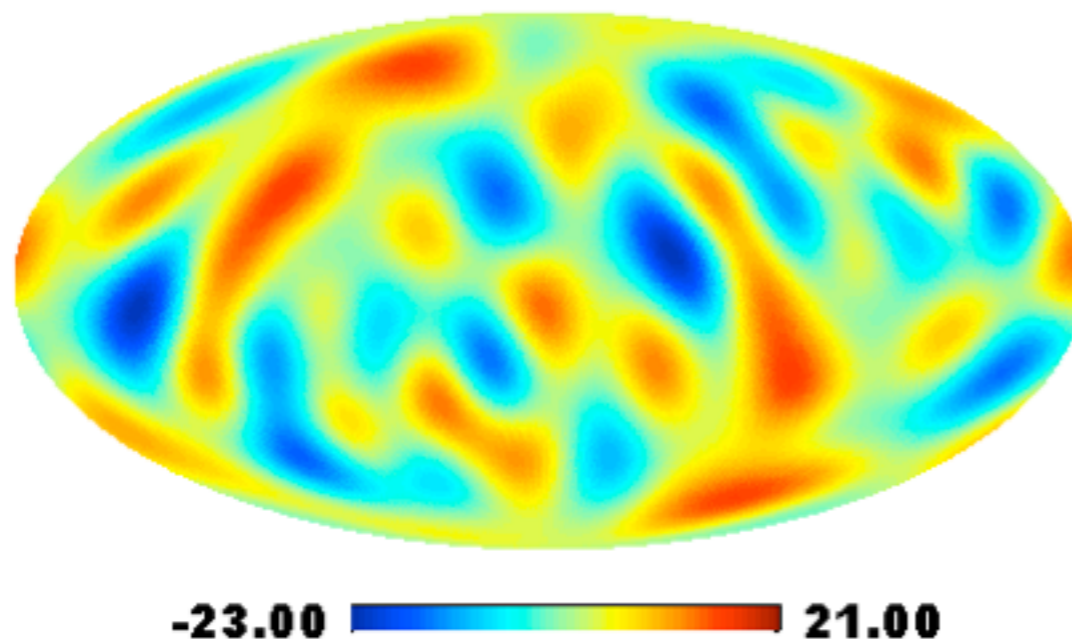
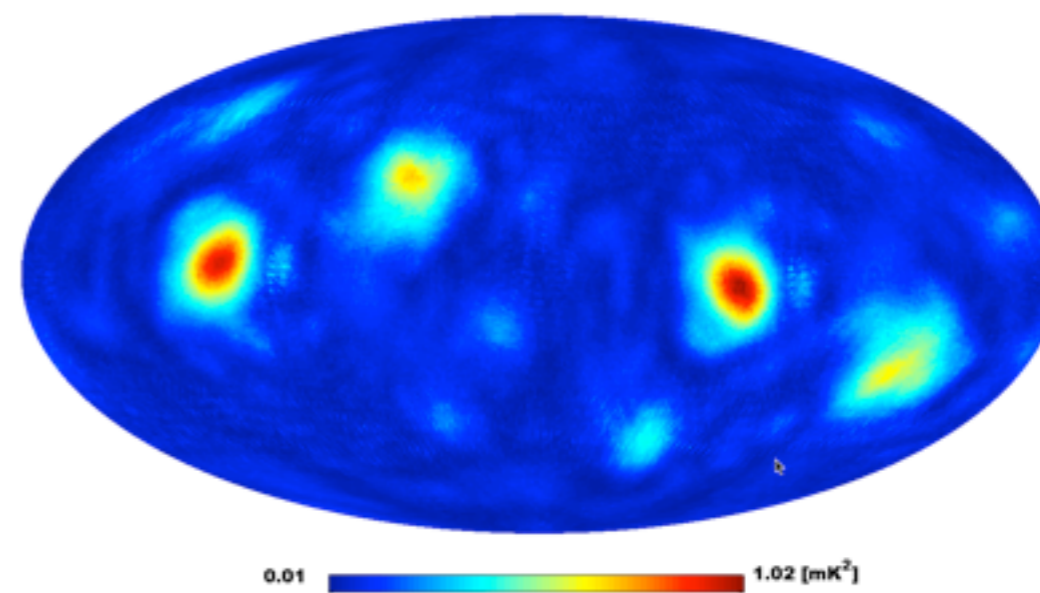
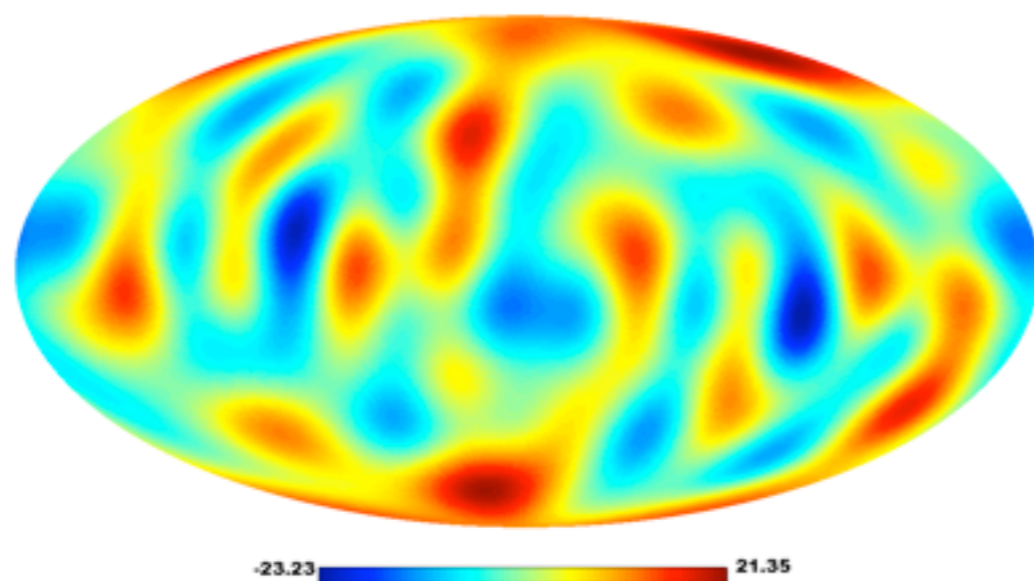


0.05 3.95

Simulation vs. data: Moving PIP

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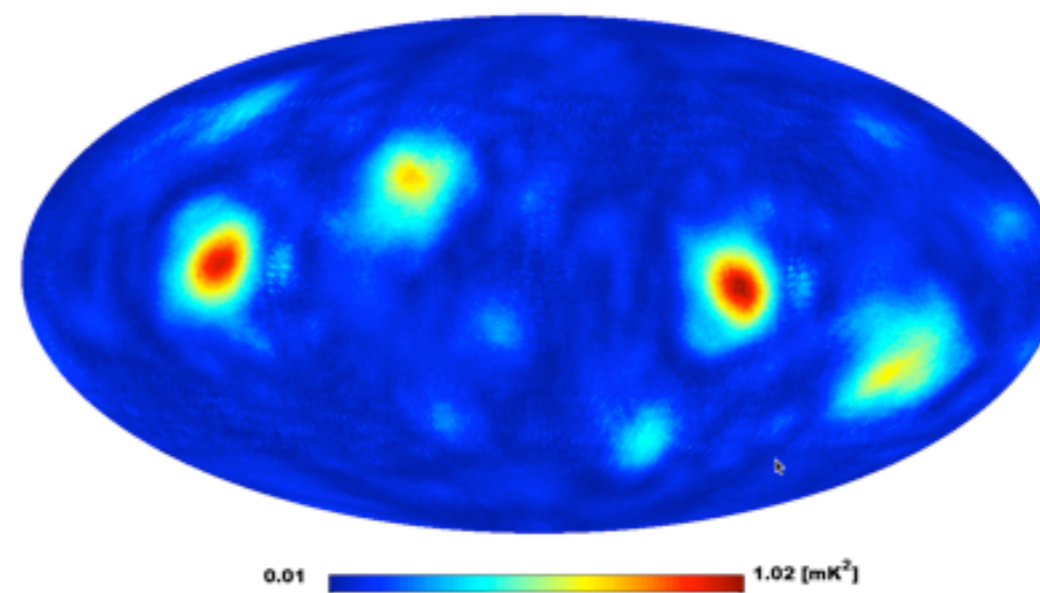
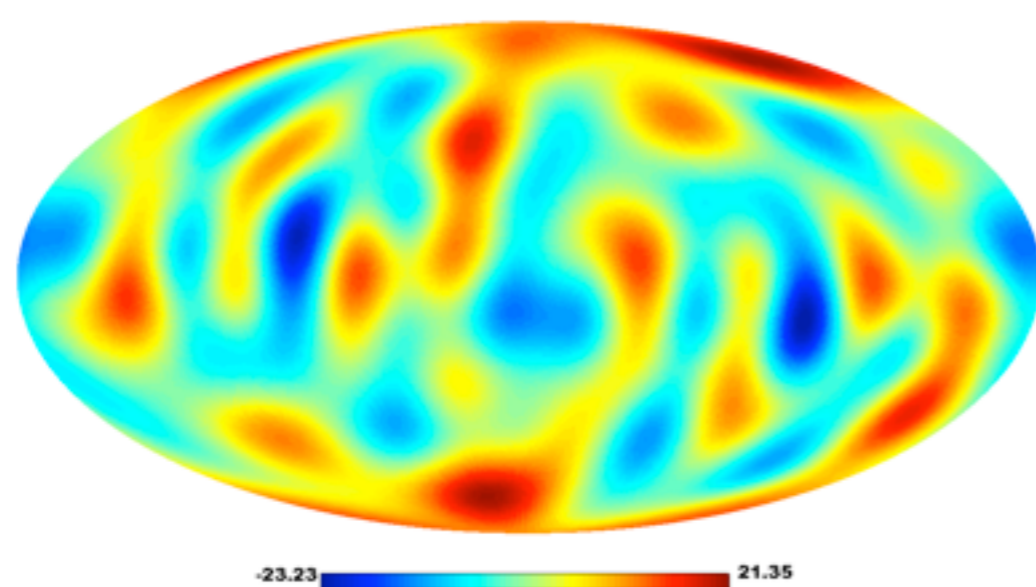
Simulating LCDM + a moving PIP (with fine-tuned location and velocity):



Simulation vs. data: Moving PIP

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Simulating LCDM + a moving PIP (with fine-tuned location and velocity):

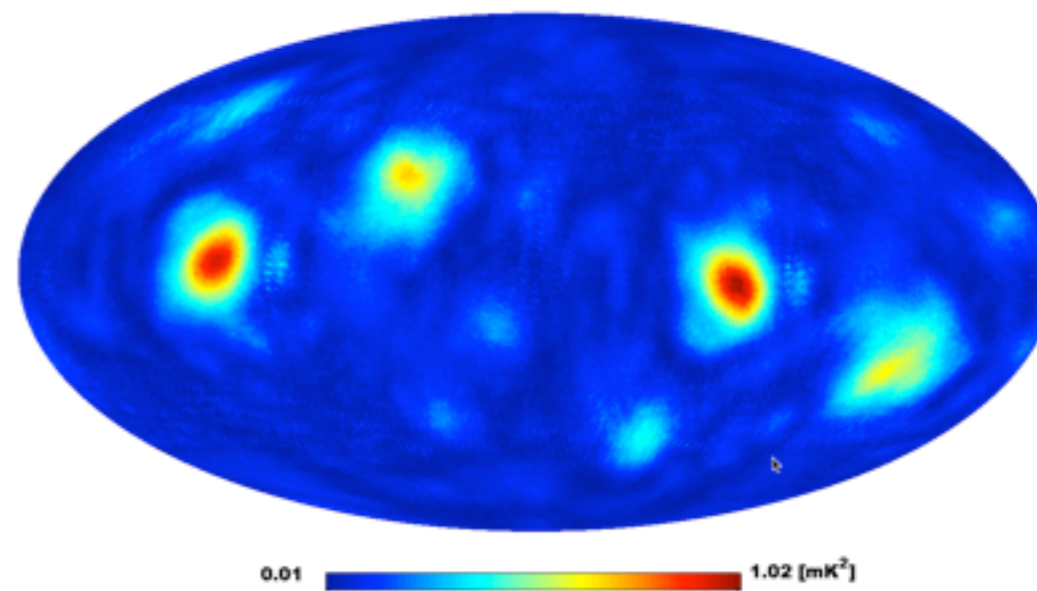
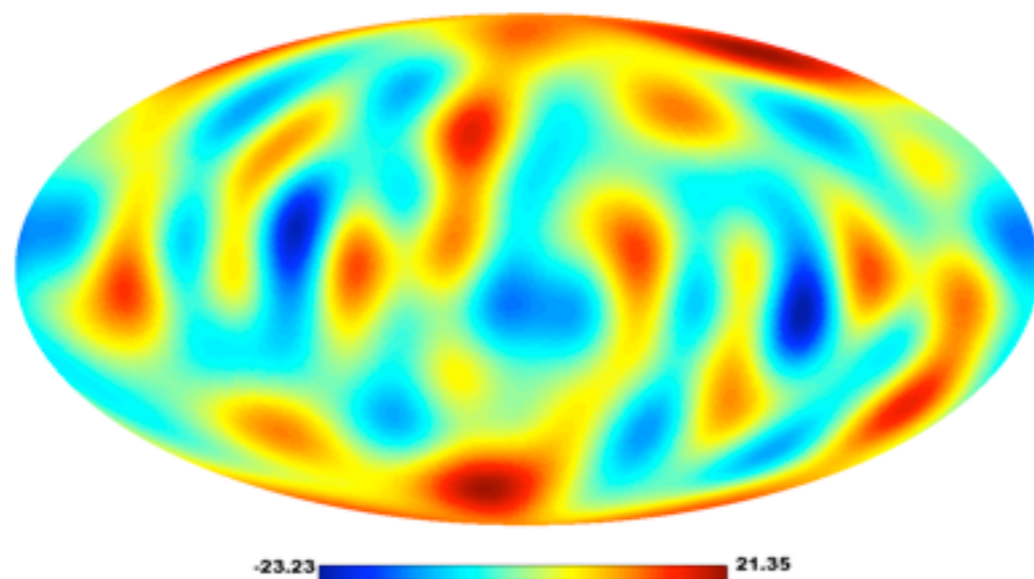


Real data:

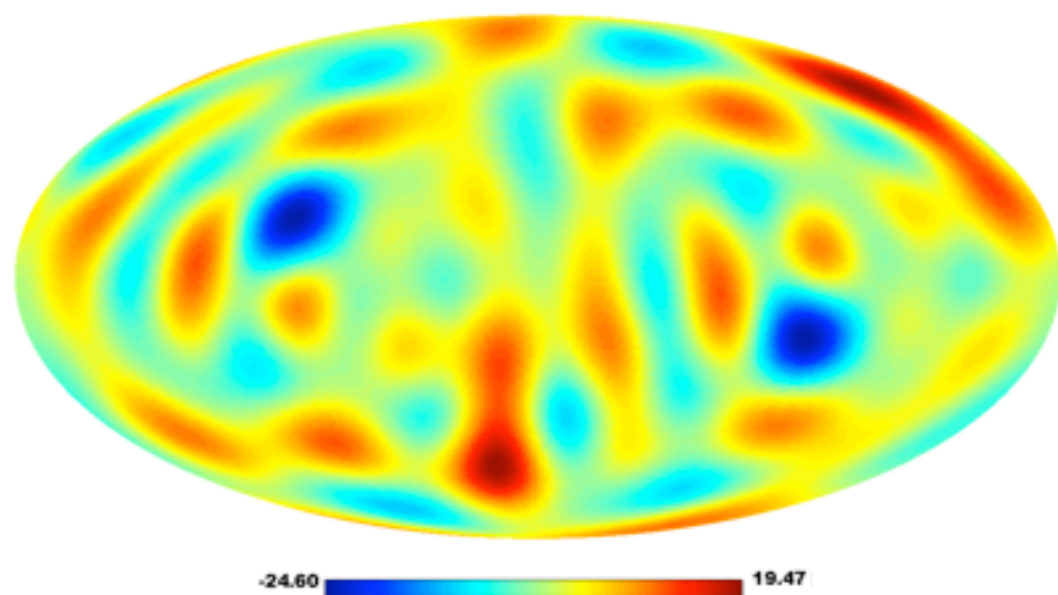
Simulation vs. data: Moving PIP

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Simulating LCDM + a moving PIP (with fine-tuned location and velocity):



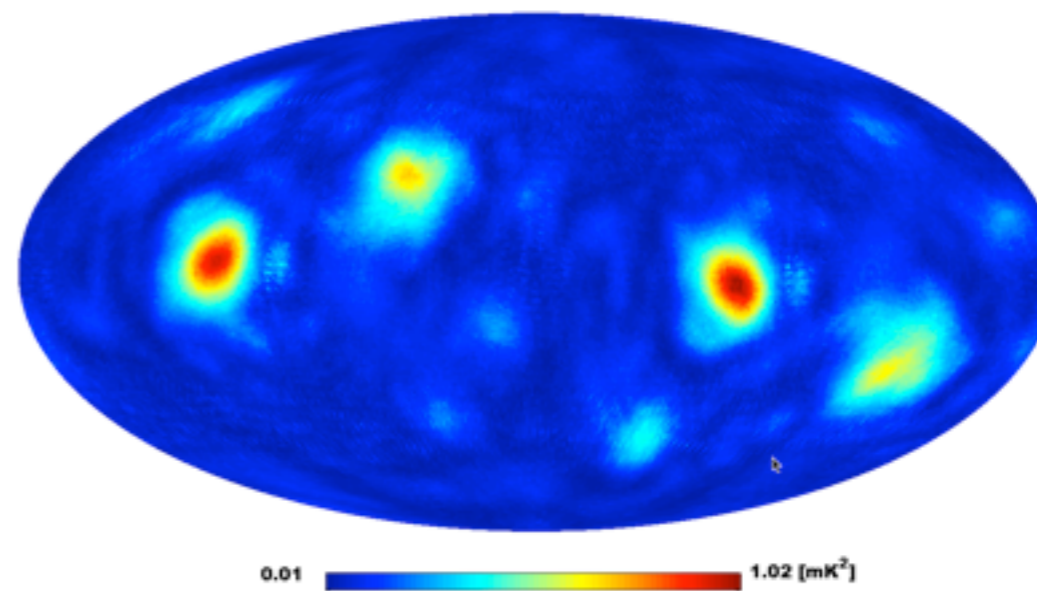
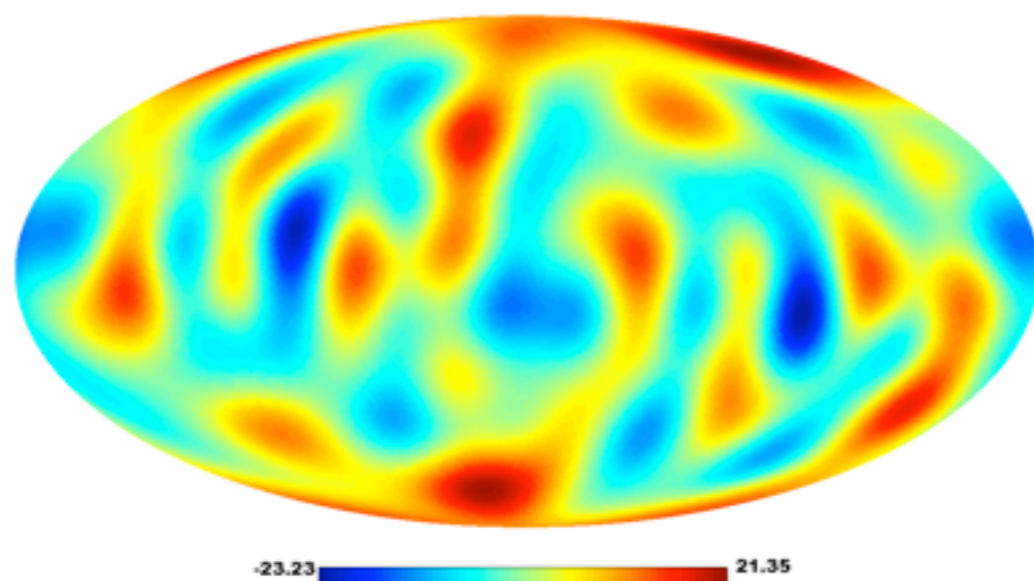
Real data:



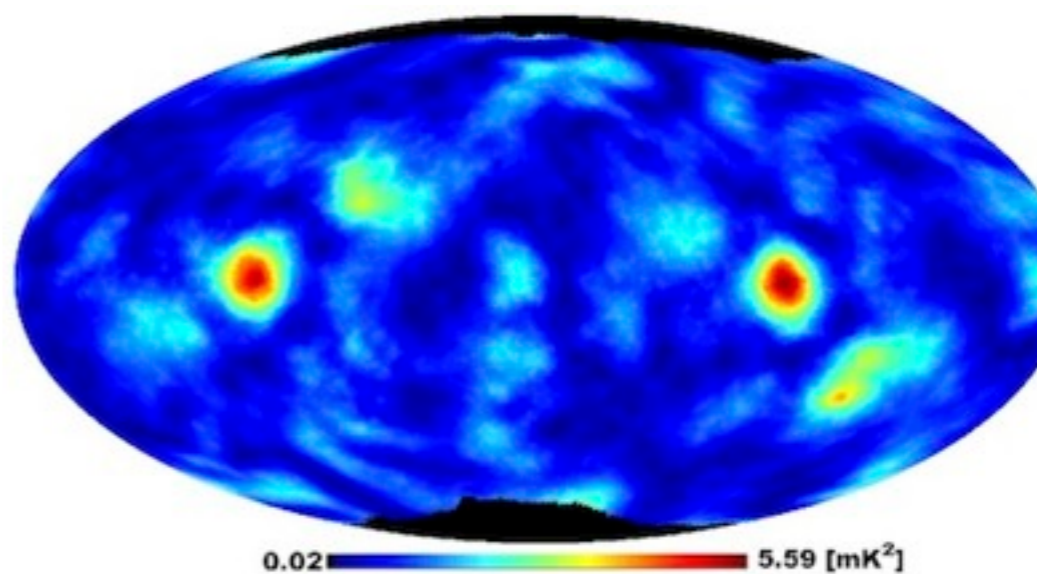
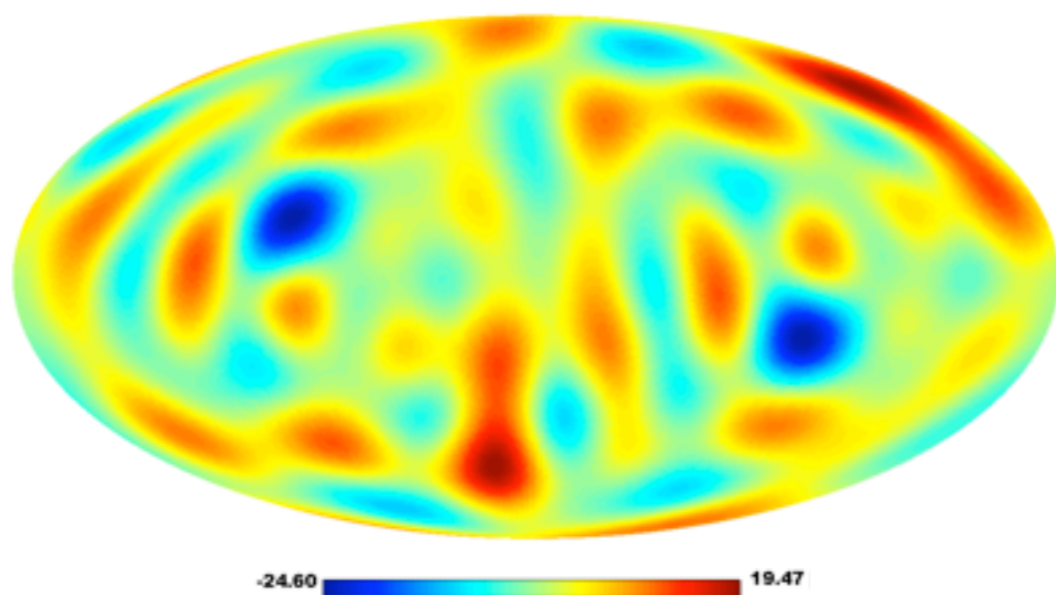
Simulation vs. data: Moving PIP

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Simulating LCDM + a moving PIP (with fine-tuned location and velocity):



Real data:

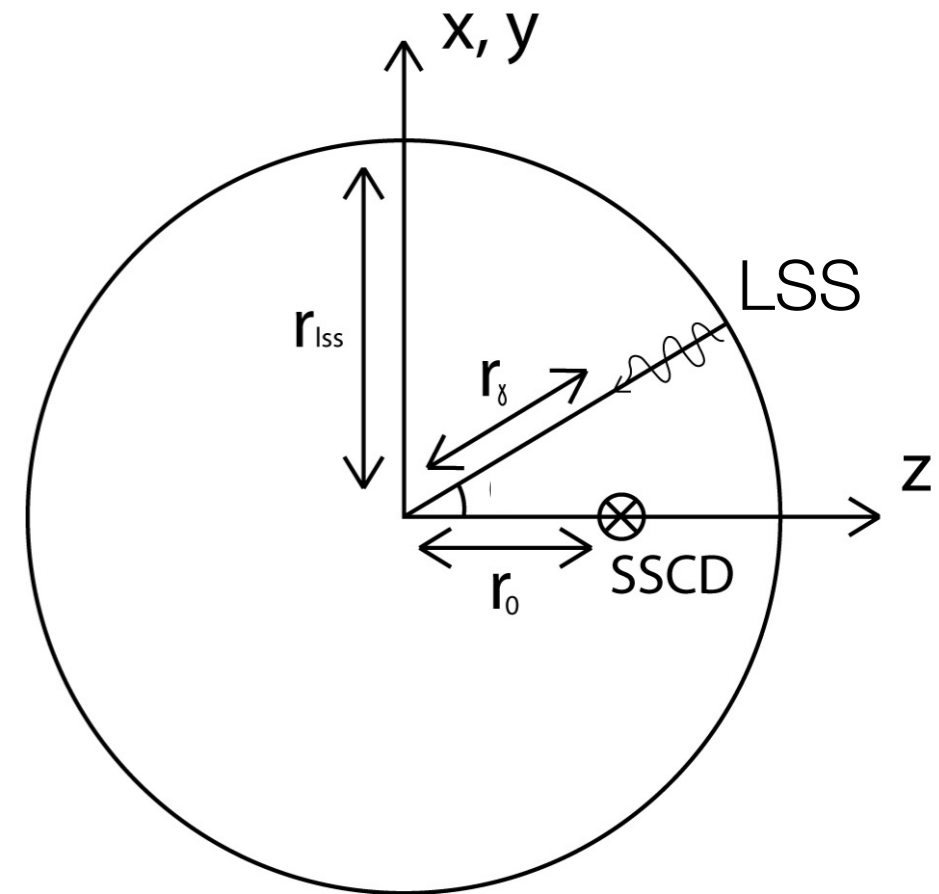


Constraining Parameter Space

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Constraining Parameter Space

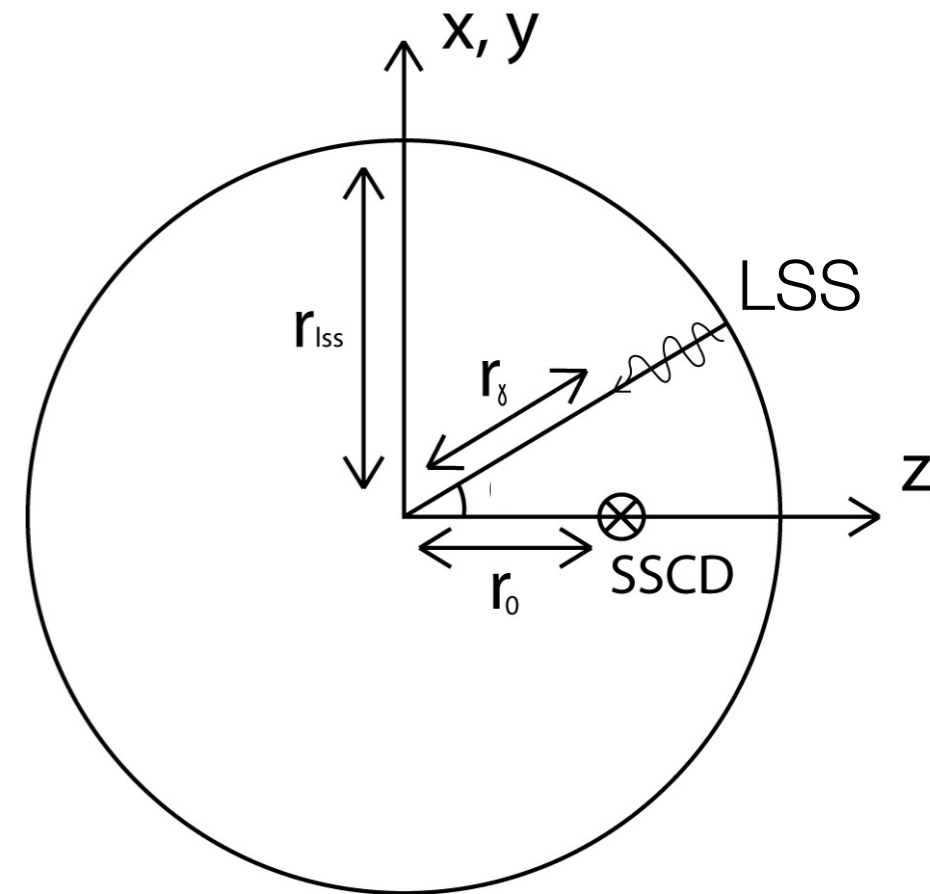
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Constraining Parameter Space

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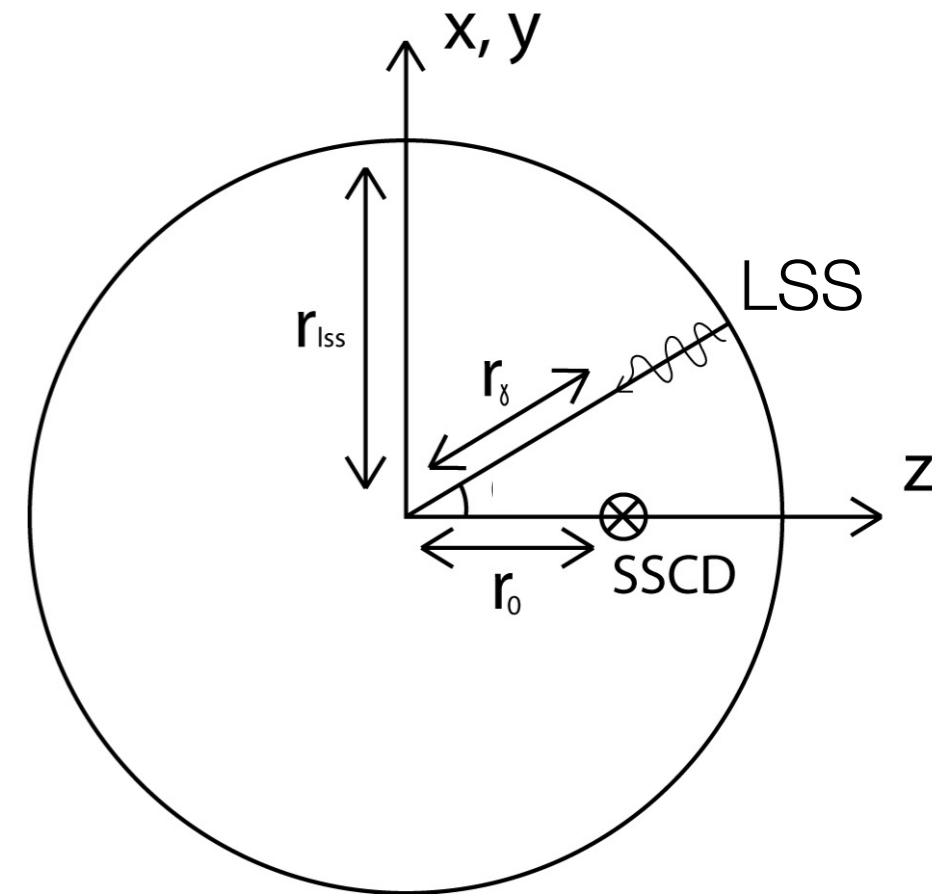
- Constrain the model with magnitude and location of:



Constraining Parameter Space

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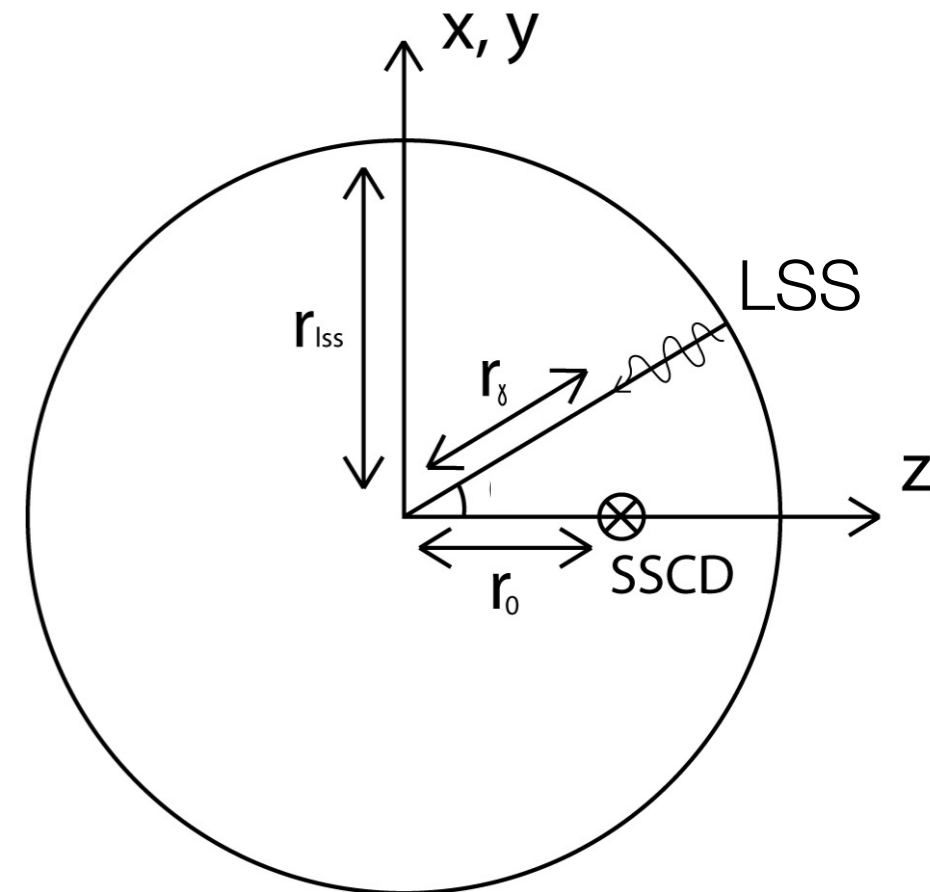
- Constrain the model with magnitude and location of:
 - Giant rings.
 - Bulk flow.



Constraining Parameter Space

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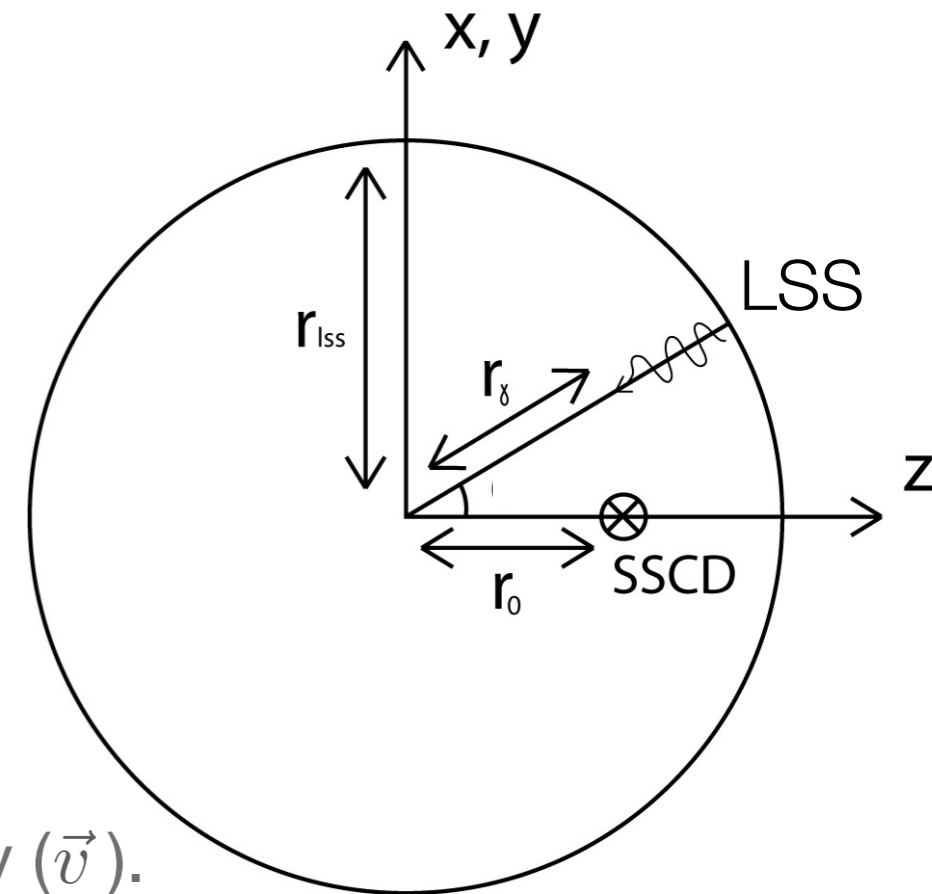
- Constrain the model with magnitude and location of:
 - Giant rings.
 - Bulk flow.
 - Odd parity.



Constraining Parameter Space

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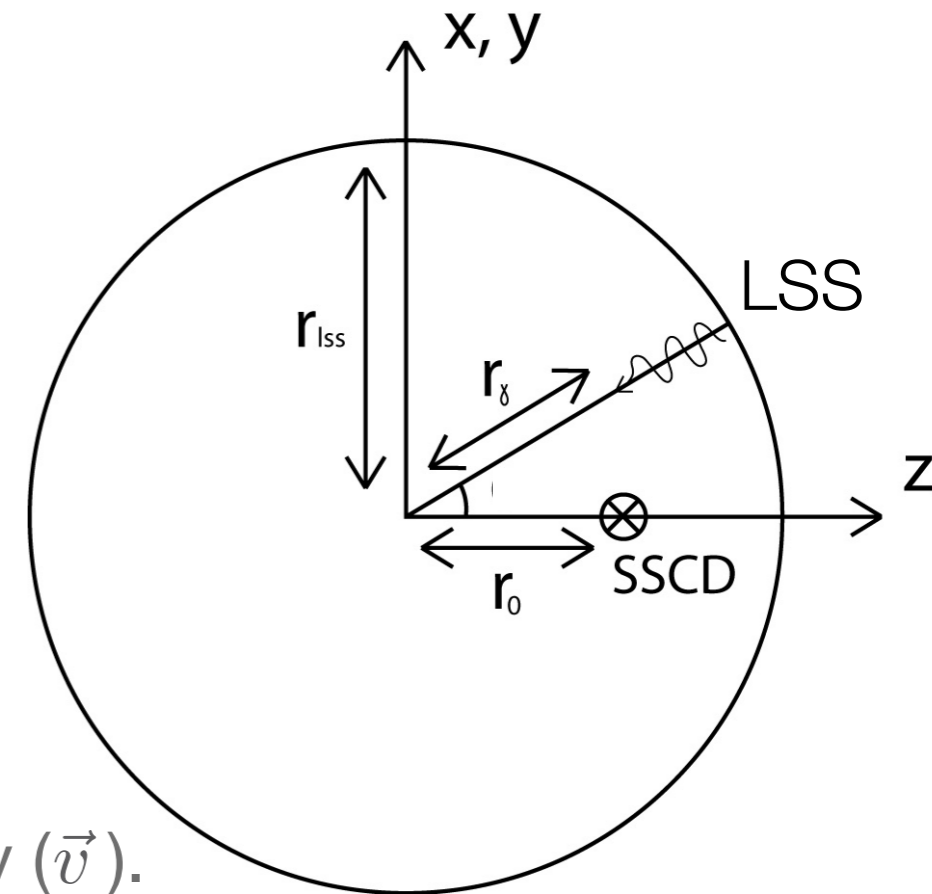
- Constrain the model with magnitude and location of:
 - Giant rings.
 - Bulk flow.
 - Odd parity.
- Free parameters: amplitude (λ), location (r_0), velocity (\vec{v}).



Constraining Parameter Space

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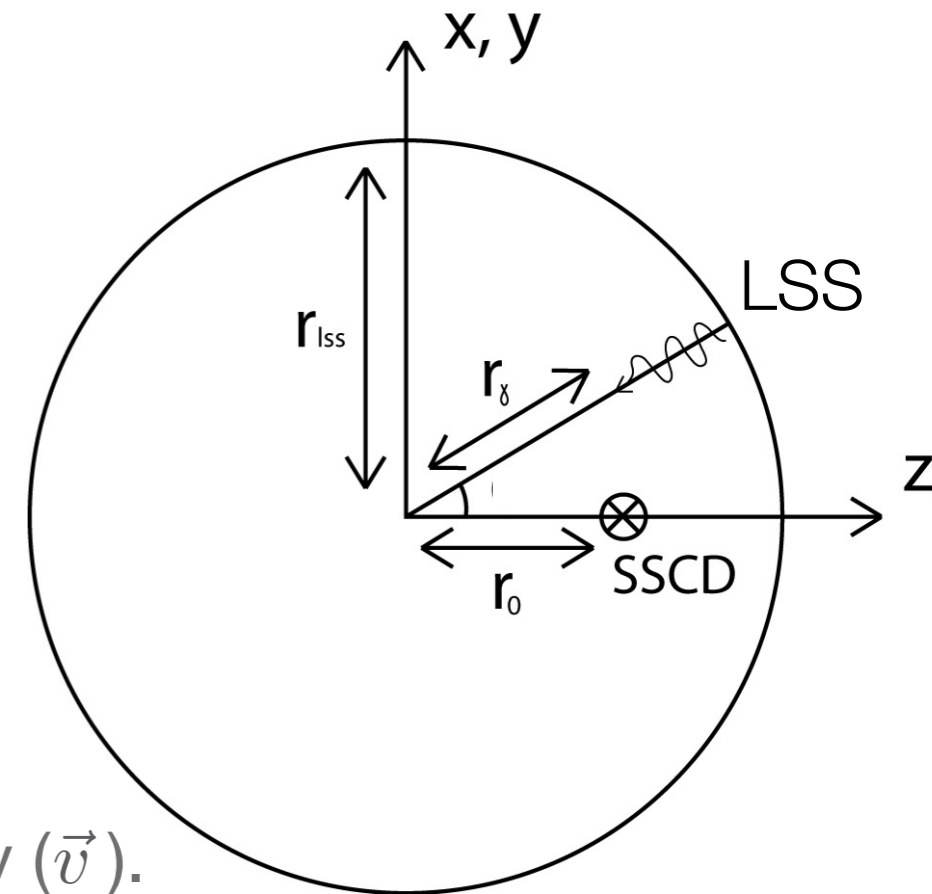
- Constrain the model with magnitude and location of:
 - Giant rings.
 - Bulk flow.
 - Odd parity.
- Free parameters: amplitude (λ), location (r_0), velocity (\vec{v}).
- Calculate SW + ISW for CMB, Peculiar Velocity for Local Group.



Constraining Parameter Space

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- Constrain the model with magnitude and location of:
 - Giant rings.
 - Bulk flow.
 - Odd parity.
- Free parameters: amplitude (λ), location (r_0), velocity (\vec{v}).
- Calculate SW + ISW for CMB, Peculiar Velocity for Local Group.
- Fit to data.



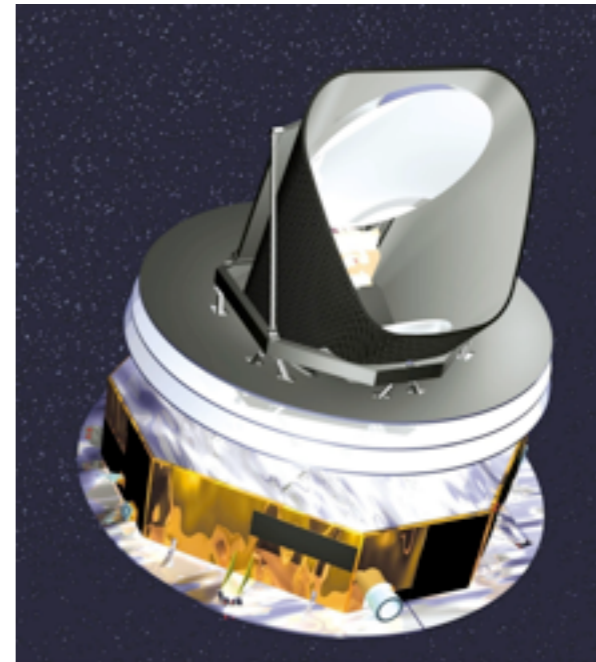
Future

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Future

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Using new observational data:

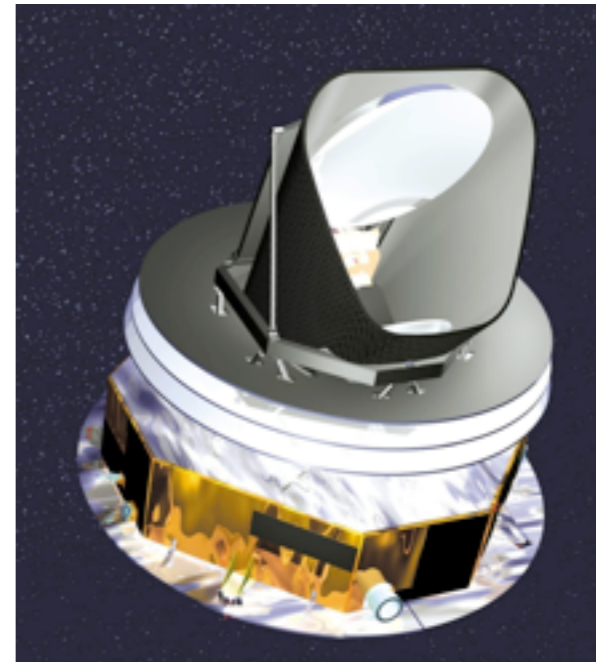


Future

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Using new observational data:

- Planck: T_{CMB} Anomalies.

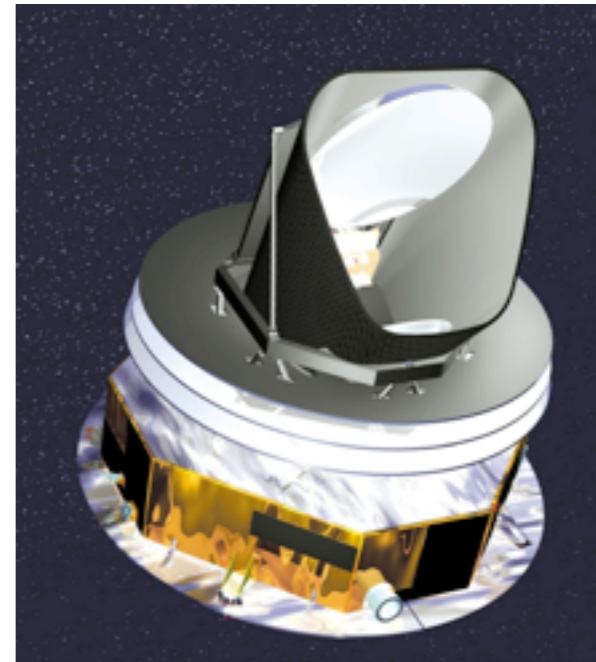


Future

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Using new observational data:

- Planck: T_{CMB} Anomalies.
- Planck: Weak Lensing.

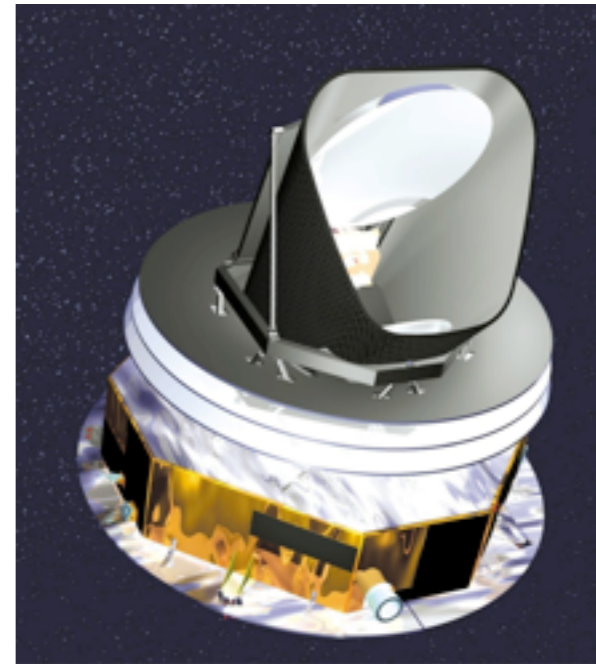


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Using new observational data:

- Planck: T_{CMB} Anomalies.
- Planck: Weak Lensing.
- Planck: BF from kSZ effect.

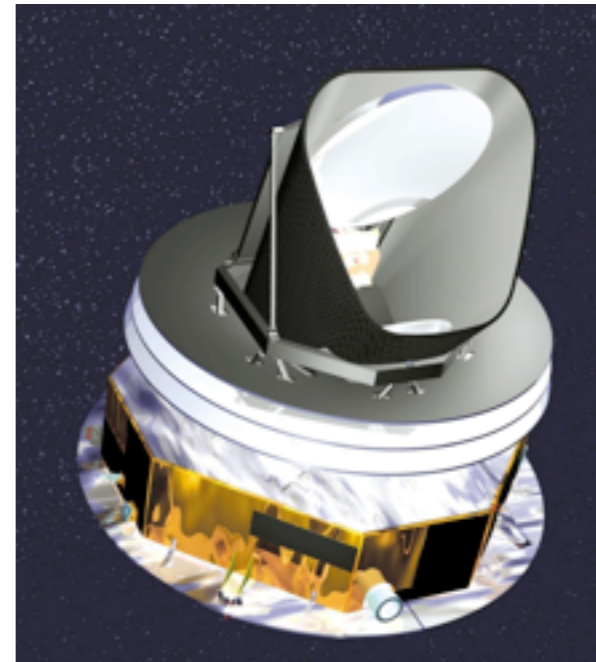


Future

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Using new observational data:

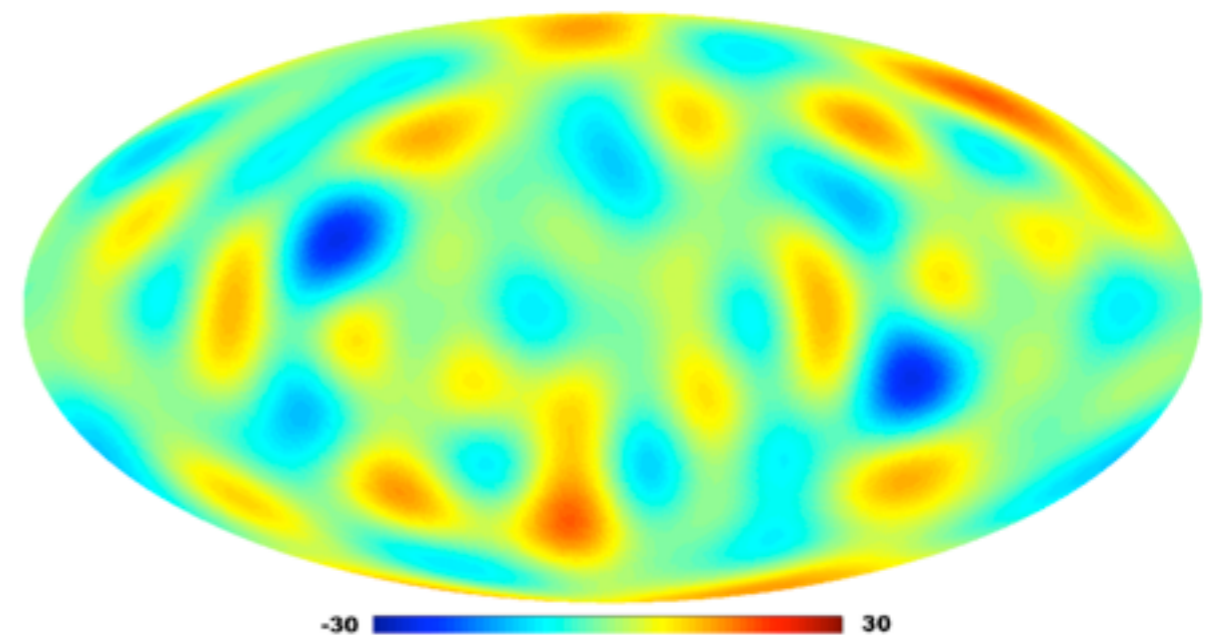
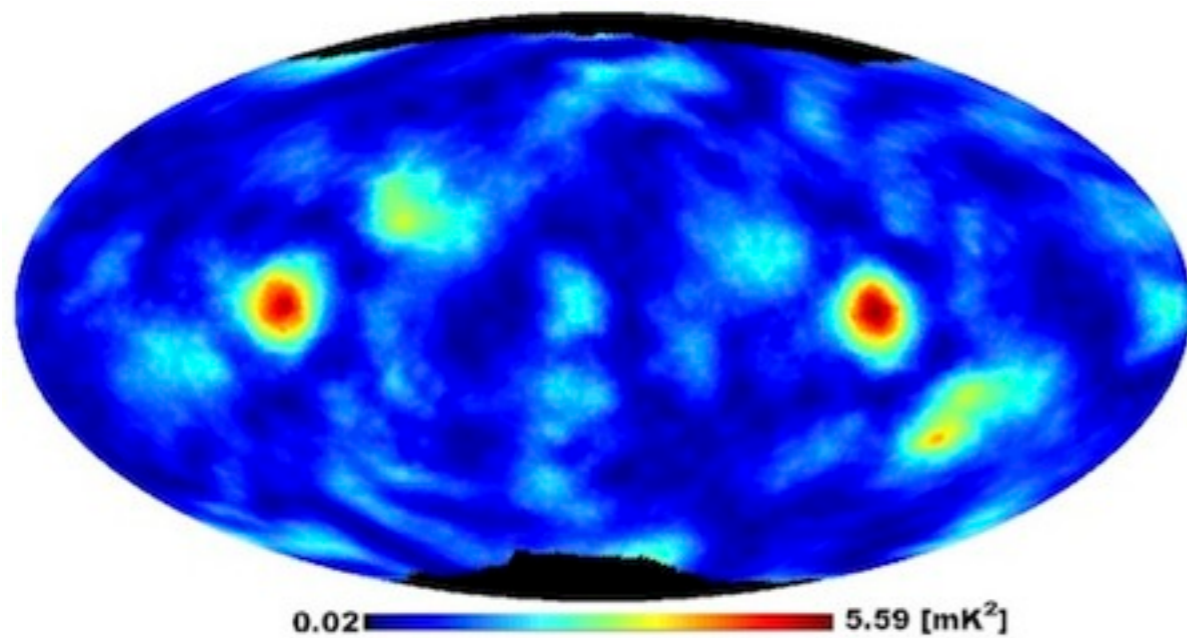
- Planck: T_{CMB} Anomalies.
- Planck: Weak Lensing.
- Planck: BF from kSZ effect.
- SNIa: BF converges at high- z ?



Supernova
Cosmology Project

Takeaway

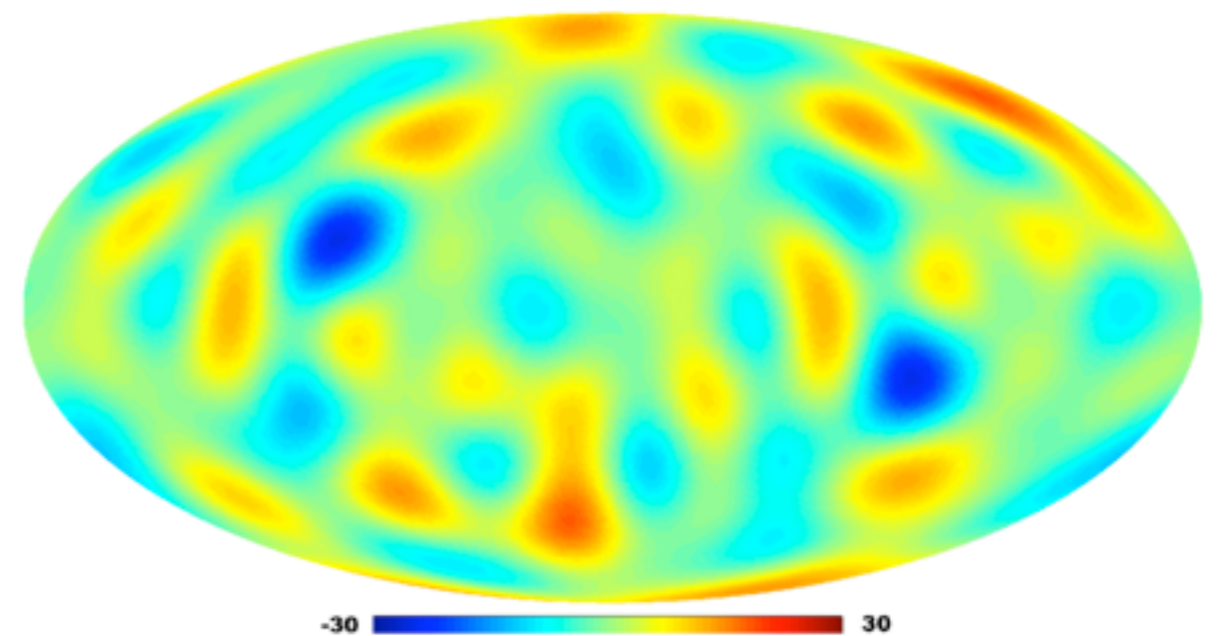
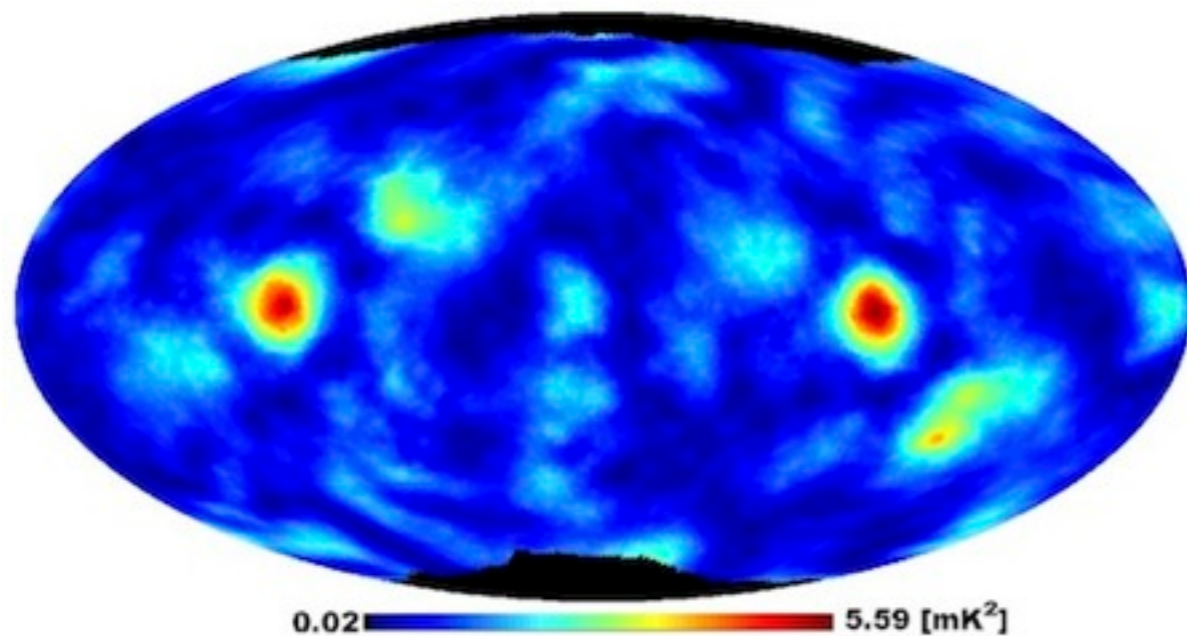
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Takeaway

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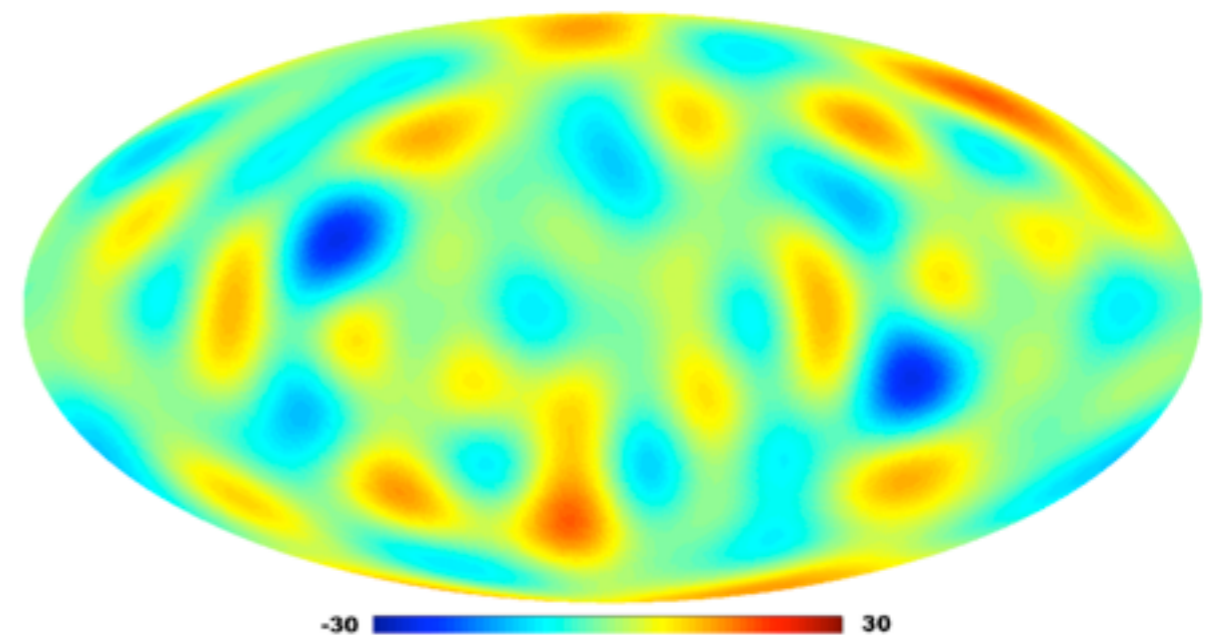
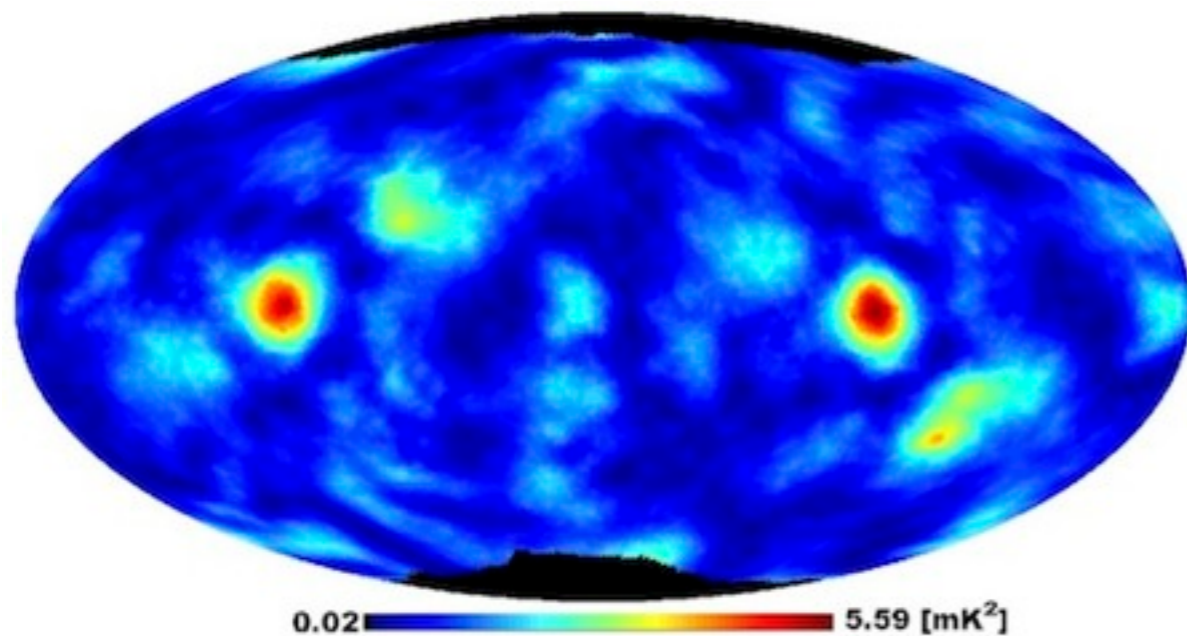
- Phenomenology of pre-inflationary physics: large scale anomalies.



Takeaway

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- Phenomenology of pre-inflationary physics: large scale anomalies.

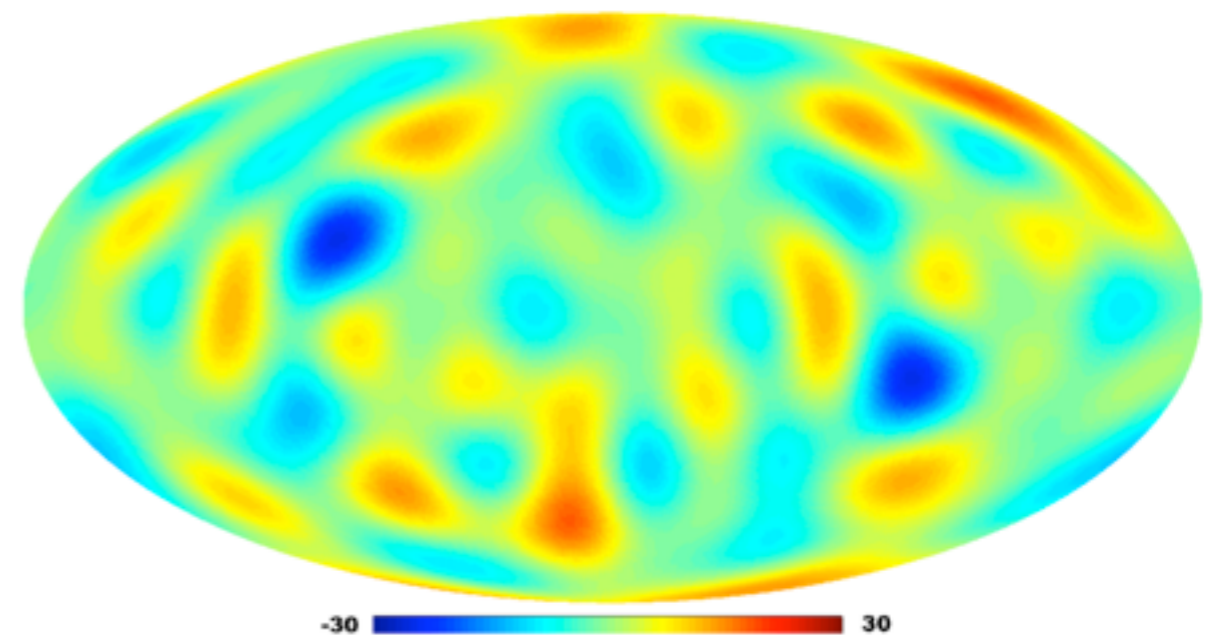
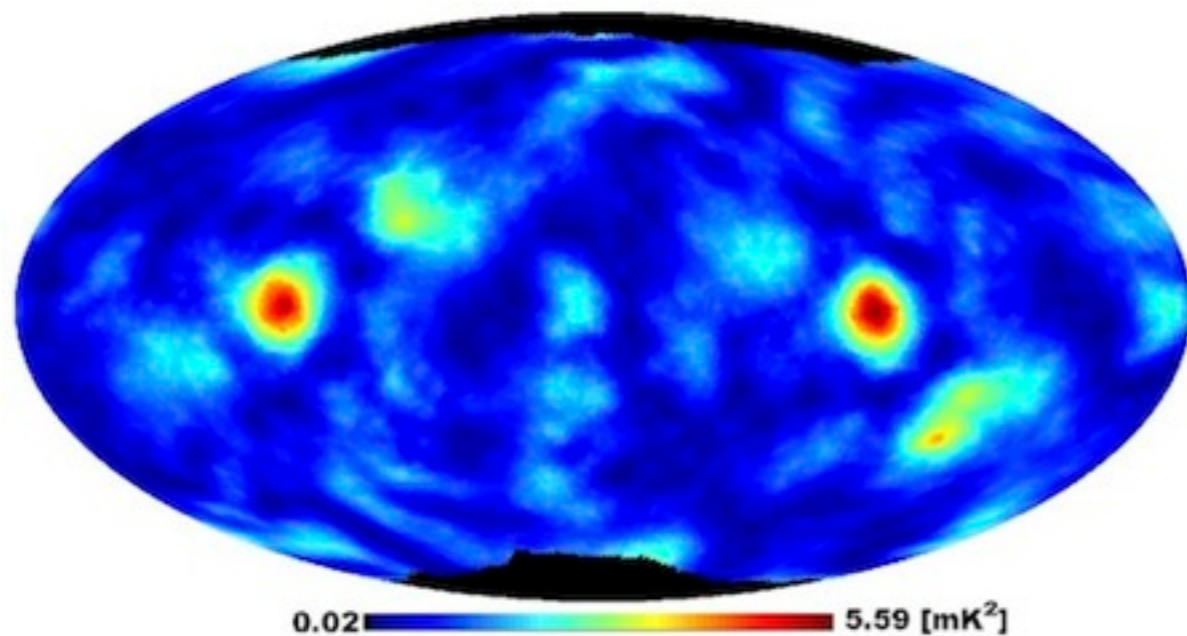


- Bubble Collisions

Takeaway

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- Phenomenology of pre-inflationary physics: large scale anomalies.

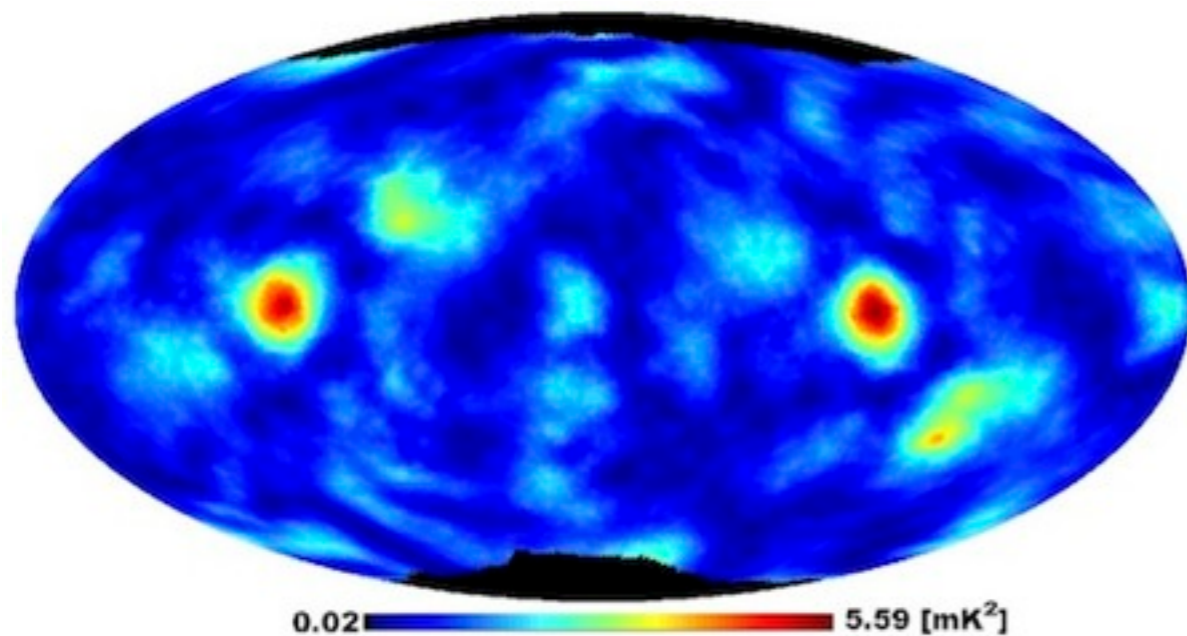


- Bubble Collisions
Produce rings profile?

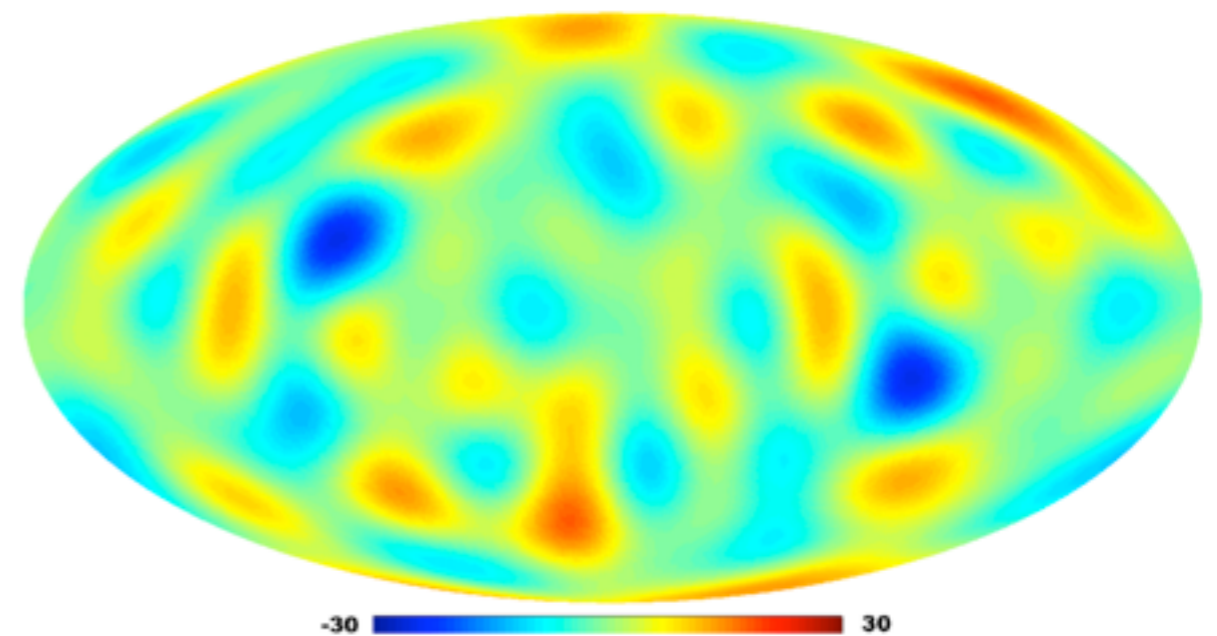
Takeaway

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- Phenomenology of pre-inflationary physics: large scale anomalies.



- Bubble Collisions
Produce rings profile?

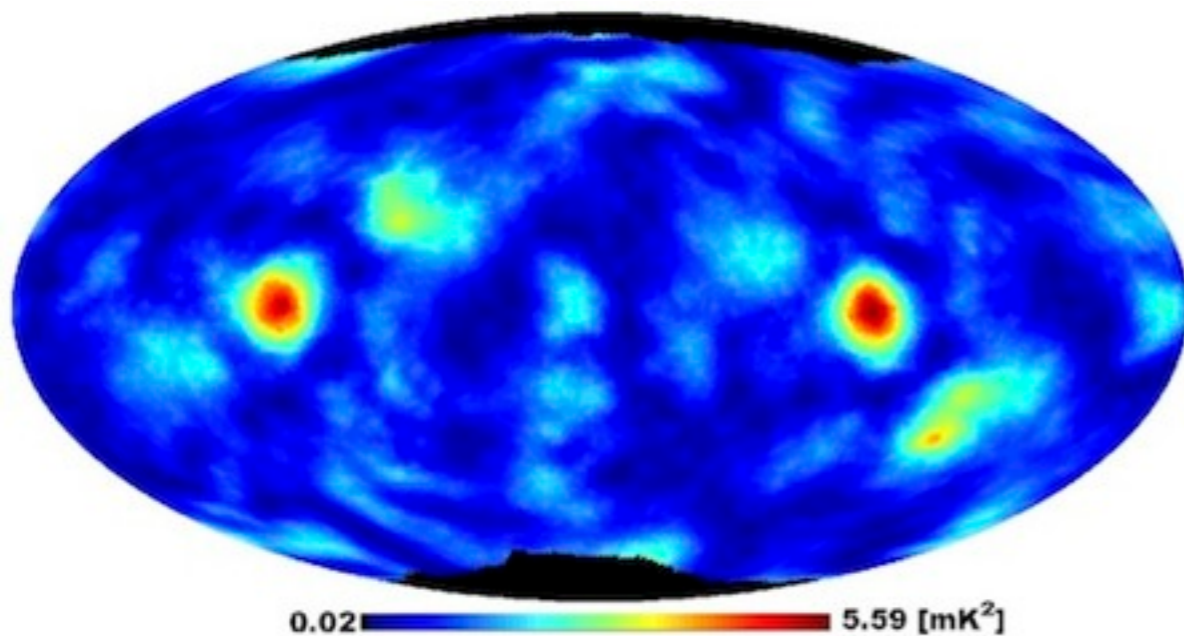


- Finite Universe Topologies

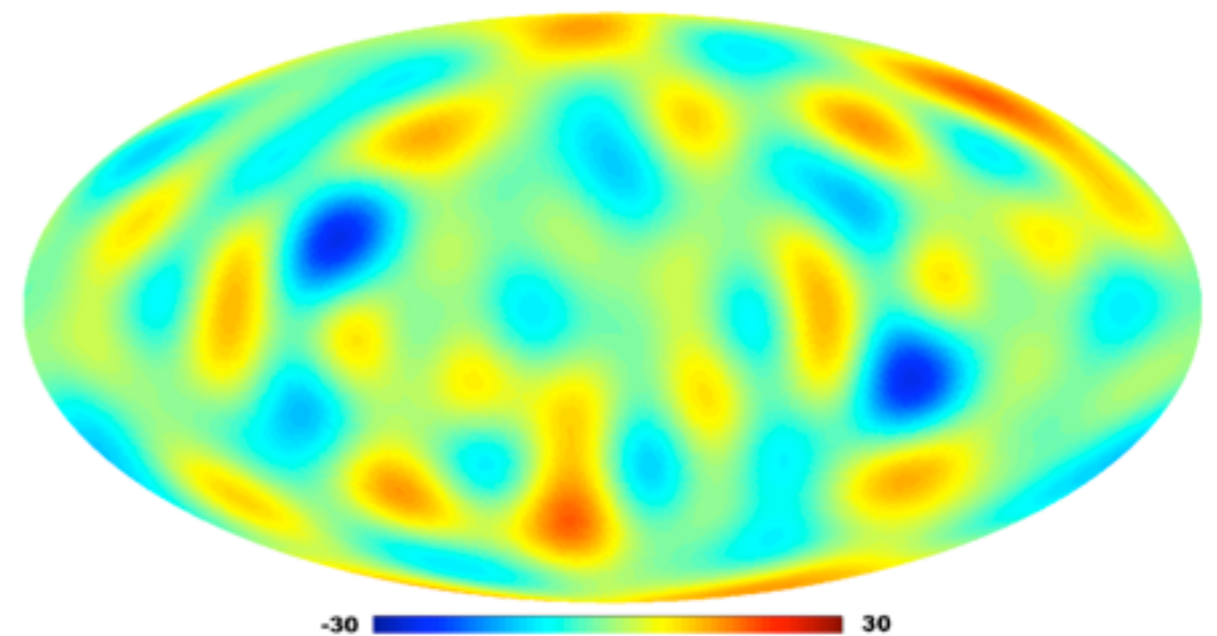
Takeaway

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- Phenomenology of pre-inflationary physics: large scale anomalies.



- Bubble Collisions
Produce rings profile?



- Finite Universe Topologies
Produce odd parity?

Thank You!

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